

Measurement of the $^{90,92}\text{Zr}(p,\gamma)^{91,93}\text{Nb}$ reactions for the nucleosynthesis of elements around $A=90$

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Abstract

Cross section measurements of the reactions $^{90,92}\text{Zr}(p,\gamma)^{91,93}\text{Nb}$ were performed using the NSCL SuN detector at the University of Notre Dame. These reactions are part of the nuclear reaction flow for the synthesis of the light p nuclei. For the $^{90}\text{Zr}(p,\gamma)^{91}\text{Nb}$ reaction the new measurement resolves the disagreement between previous results. For the $^{92}\text{Zr}(p,\gamma)^{93}\text{Nb}$ reaction the present work reports the first measurement of this reaction cross section. Both reaction cross sections are compared to theoretical calculations and a very good agreement with the standard NON-SMOKER model is observed.

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I. INTRODUCTION

The origin of the heavy elements remains one of the overarching questions for the nuclear astrophysics community. The general concepts of heavy element nucleosynthesis were introduced in 1957 by the famous B²FH publication [1], however, 55 years later, many of the details of the responsible astrophysical processes are still not well understood.

A small fraction of the heavy elements is located on the neutron deficient side of the valley of stability and forms the group of the so-called p nuclei [2, 3]. This group consists of 35 stable isotopes in the mass region between Se (Z=34) and Hg (Z=80) that cannot be synthesized by the two neutron-capture-induced processes (s- and r-process) (e.g. [4]). Their synthesis requires a different astrophysical mechanism traditionally called the “p process”. It is not yet clear whether the p process is a single or multiple independent astrophysical scenarios. A number of astrophysical settings have been proposed for the site of the production of the p nuclei, such as type II supernovae (SNII) [5, 6], supernovae type Ia (e.g. [7] and references therein), the ν p process in the neutrino driven winds of SNII [8, 9], the rp-process in matter accreted on the surface of neutron stars [10] and others [2, 3].

The most favored scenario, and the one that has been investigated the most to date, takes place in SNII when the shock front passes through the O/Ne-rich layers of the massive star [5, 6]. In this scenario, the main reactions responsible for the formation of the p nuclei are neutron, proton and alpha-particle photodisintegrations together with a possible contribution to the light p nuclei from proton-captures at the higher temperatures. Due to the dominance of photodisintegration reactions, this process is also called “ γ process” and it takes place at temperatures between 1.5 and 3.5 GK. This scenario, although able to produce p nuclei in the whole mass range, is not able to reproduce the abundance patterns observed in solar system samples. Large underproductions relative to the solar abundance pattern occur in the lighter masses, especially for ^{92,94}Mo and ^{96,98}Ru. Many authors investigated the possible sources of uncertainties in the predicted abundance distributions ([2, 3] and references therein) and it was shown that these come both from uncertainties in the stellar conditions as well as the nuclear physics input. The present work focuses on the nuclear physics aspect of p-process nucleosynthesis via the investigation of the ^{90,92}Zr(p, γ)^{91,93}Nb reactions, in the region around the puzzling ^{92,94}Mo and ^{96,98}Ru.

Astrophysical calculations were performed using the post-processing code available in

NucNet tools [11], a suite of nuclear reaction codes developed at Clemson University. In the calculations the γ process was investigated using a full nuclear reaction network for a type II supernova explosion when the shock front passes through the O/Ne layer of a $25 M_{\odot}$ star. The calculations were performed in a multi-layer model (11 layers) using the seed distribution of a pre-explosion $25 M_{\odot}$ star. The seed distribution and temperature and density profiles were taken as described in Ref. [12]. In this model the calculations show that the p nuclei are produced mainly through photodisintegration reactions, as expected; however, for the lighter p nuclei there is an important contribution coming from (p,γ) reactions that take place in the inner-most layers, i.e. highest temperature/density regions of the O/Ne layer. This contribution was already observed in the sensitivity study by Rapp *et al.* [12] and was explored in recent studies of the $^{74}\text{Ge}(p,\gamma)^{75}\text{As}$ reaction [13, 14]. Similarly, the two reactions studied in the present work could have a direct contribution to the production of $^{92,94}\text{Mo}$ via two consecutive (p,γ) reactions from $^{90,92}\text{Zr}$. These (p,γ) reactions could have a contribution at the highest temperatures, especially if the reaction rates are significantly higher than the ones currently used in astrophysical calculations presented in the JINA REACLIB database [15].

In the literature there are two previous measurements of the $^{90}\text{Zr}(p,\gamma)^{91}\text{Nb}$ reaction. One was a thick-target measurement using the activation technique by Roughton *et al.* [16]. The authors were able to extract an astrophysical reaction rate from their thick-target measurements without the need to first calculate the reaction cross section, which, as mentioned in their paper, would introduce additional uncertainties. The second measurement was performed in-beam by Laird *et al.* [17] by measuring the γ transitions populating the ground state, or the first excited metastable state of the final nucleus ^{91}Nb . The individual contributions of the γ transitions were summed to give the total (p,γ) cross section. Since the two works did not produce the same final quantity it is hard to compare them directly. Nevertheless, comparison to the same theoretical calculations (section IV) show that the two measurements are not in agreement. For this reason, in the present work an independent measurement of the $^{90}\text{Zr}(p,\gamma)^{91}\text{Nb}$ reaction cross section was performed. The results of the two previous measurements and of the present one will be discussed in section IV. The $^{92}\text{Zr}(p,\gamma)^{93}\text{Nb}$ reaction has not been measured before.

The present paper is organized as follows: the experimental setup and the technique are presented in section II, the results for the two reactions of interest are shown in section

III, the discussion of these results, comparison to theoretical calculations and astrophysical implications are presented in section IV and the paper is completed with the conclusions in section V.

II. EXPERIMENTAL

The experiment was performed at the FN Tandem Accelerator of the University of Notre Dame. A proton beam at energies between 2.0 and 5.0 MeV (uncertainty ≈ 2 keV) interacted with the two self-supporting and isotopically enriched Zr targets. The target thickness was measured using the Rutherford Back-Scattering (RBS) technique at the Hope College Ion Beam Analysis Laboratory (HIBAL). For the RBS measurements, a 2.95 MeV alpha beam was used. After impinging on the target the scattered alpha particles were detected at 168.2° in a Si surface barrier detector with a 40 keV energy resolution. The RBS spectra were simulated using the software SIMNRA [18], where all the parameters of the experiment were fixed and the only free parameter was the target thickness. The measured thickness was $969(48)$ $\mu\text{g}/\text{cm}^2$ for the ^{90}Zr target and $960(47)$ $\mu\text{g}/\text{cm}^2$ for the ^{92}Zr one. The two targets were highly enriched in the isotope of interest (98(1)%). The beam current on target was at most 22 nA to minimize possible dead-time effects. The total accumulated charge was between 1 and 10 μC .

The cross section measurements were performed using the newly commissioned Summing NaI(Tl) (SuN) detector of the National Superconducting Cyclotron Laboratory, Michigan State University. SuN is a barrel shaped NaI(Tl) scintillator, 16 inches in diameter and 16 inches in height with a bore hole along its axis. A detailed description of SuN, the electronics setup, and data acquisition system can be found in Ref. [19].

The targets were mounted in the center of the detector where the large angular coverage and high detection efficiency of SuN allowed for the use of the γ -summing technique [19, 20]. In this technique the capture of a proton by the target nucleus is measured by summing the γ rays emitted during the de-excitation of the produced nucleus. The energy of the entry state of the produced nucleus is $E_\Sigma = Q + E_{cm}$, where Q is the Q-value of the reaction and E_{cm} is the center-of-mass energy of the initial system. In the summing technique the γ spectrum is expected to show the so-called “sum peak”, located at the energy E_Σ , in the high energy region.

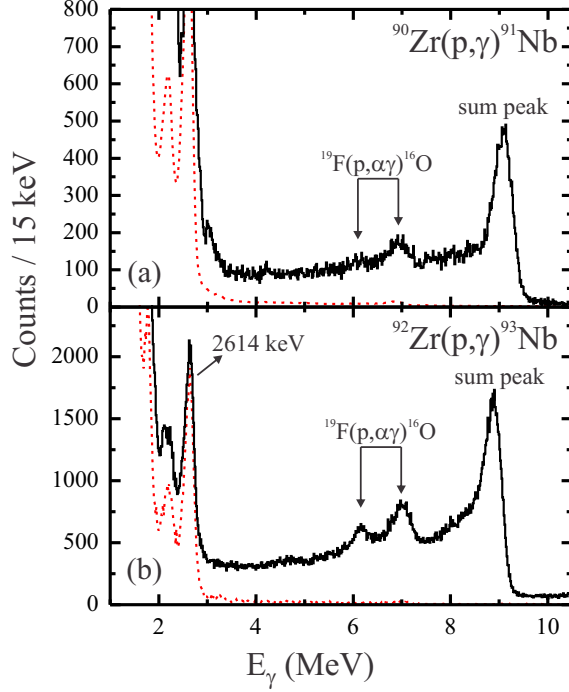


FIG. 1: (Color online) Spectra taken with the SuN detector for (a) the $^{90}\text{Zr}(p,\gamma)^{91}\text{Nb}$ reaction at proton beam energy of 4 MeV and (b) the $^{92}\text{Zr}(p,\gamma)^{93}\text{Nb}$ reaction at 2.8 MeV. In both figures the black-solid line represents the beam-on-target spectrum while the red-dashed line shows the normalized room background.

Typical spectra for the $^{90}\text{Zr}(p,\gamma)^{91}\text{Nb}$ and $^{92}\text{Zr}(p,\gamma)^{93}\text{Nb}$ reactions are shown in Fig. 1. The high energy part of the spectra is dominated by the sum peak for each of the reactions. For the case of the $^{90}\text{Zr}(p,\gamma)^{91}\text{Nb}$ reaction shown in Fig. 1a, the Q-value is 5.16 MeV and the beam energy was 4 MeV resulting in a sum peak at 9.09 MeV (assuming that the interaction takes place at the center of the target). For the case of the $^{92}\text{Zr}(p,\gamma)^{93}\text{Nb}$ reaction shown in Fig. 1b, the Q-value is 6.04 MeV and the beam energy was 2.8 MeV resulting in a sum peak at 8.78 MeV. In the region between 6 and 7 MeV the spectrum is contaminated by the γ rays from the reaction $^{19}\text{F}(p,\alpha\gamma)^{16}\text{O}$, which is a common contaminant reaction in proton induced measurements. The low energy region of the spectrum, below 3 MeV, is dominated by the room background γ rays from ^{40}K at 1461 keV and from the Th decay chain at 2614 keV. The presence of any beam-induced background will also contribute to the low-energy regions of the spectra. The only source of background in the high-energy region, where the sum

peak is observed, comes from cosmic rays. For all measurements presented in this work the cosmic ray background contribution was negligible.

III. RESULTS

The reaction cross section can be determined from the analysis of the sum peak in each of the spectra. The peak integration was based on 3σ limits on each side of the centroid of the sum peak and linear background subtraction. This method was chosen to be consistent with the efficiency calibration procedure described in detail in Ref. [19]. In both reactions presented in this work the final nucleus also has a long-lived metastable state, which will not be summed in the main sum peak. In this case, a second sum peak is expected in the spectra that will be located at an energy equal to the difference between the entry state and the metastable state [20, 21]. For the $^{90}\text{Zr}(p,\gamma)^{91}\text{Nb}$ reaction, the metastable state in ^{91}Nb is located at 104.6 keV with a half-life of 60.9 d [22]. For the $^{92}\text{Zr}(p,\gamma)^{93}\text{Nb}$ reaction, the metastable state in ^{93}Nb is located at 30.8 keV with a half-life of 16.1 y [22]. In both cases the metastable-state sum peak is expected to be overlapping with the ground-state sum peak and show a small extension on the low energy side, which was taken into account in the analysis.

Using the summing technique, the reaction cross section can be calculated from the analyzed spectra with:

$$\sigma = \frac{A}{N_A N_b} \frac{1}{\xi} \frac{I_\Sigma}{\epsilon_\Sigma} \quad (1)$$

where A is the atomic weight of the target nucleus (in amu), N_A is the Avogadro constant, N_b is the total number of beam particles impinging on the target, ξ is the target thickness, I_Σ is the sum-peak intensity and ϵ_Σ is the sum-peak efficiency. N_b was measured during the experiment using a current integrator, ξ was determined as mentioned before by means of RBS analysis, and I_Σ was the result of the sum-peak analysis described in the previous paragraph. The sum peak efficiency ϵ_Σ depends on the sum-peak energy but also on the multiplicity of the γ cascade. Two techniques have been developed in the literature for the experimental determination of ϵ_Σ , the ‘‘In/Out ratio’’ method [20] and the ‘‘hit pattern’’ technique [19]. In the present work the sum-peak efficiency was determined using the latter method, which was developed and optimized specifically for the SuN detector. In this method, the segmentation of SuN provides a sensitive tool for determining the γ -cascade

multiplicity. The number of SuN segments that recorded a signal in each event depends on the number of emitted γ rays and on the sum-peak energy. Details on the “hit pattern” technique were given in Ref. [19]. The efficiencies for the two reactions of interest ranged from 21.0(2.1)% to 38.5(3.0)%.

The cross section results for the two reactions of interest are presented in Tables I and II. In both tables the first column presents the proton beam energy. The second column shows the center-of-mass energy with the assumption that the reaction is taking place at the center of the target. The energy uncertainty due to this assumption and due to a 2 keV uncertainty in the set beam energy is also presented in the second column. The third column presents the reaction cross section and the last column the astrophysical S-factor. The latter quantity was calculated as:

$$S = E\sigma(E)e^{2\pi\eta} \quad (2)$$

where E is the center-of-mass energy, $\sigma(E)$ is the reaction cross section and $\eta = Z_1Z_2e^2/\hbar v$ is the Sommerfeld parameter, with Z_1 and Z_2 the proton numbers of the interacting particles and v their relative velocity. The uncertainties presented in the two tables include: 5% for the target thickness estimate, 5% for the charge accumulation, of the order of 1% statistical uncertainty and approximately 10% uncertainty based on the efficiency analysis in [19].

IV. DISCUSSION

The γ process temperatures quoted in section I translate to astrophysically relevant energy ranges of 1.2 – 4 MeV for $^{90}\text{Zr}(p,\gamma)^{91}\text{Nb}$ and 1.2 – 3 MeV (i.e., below the neutron emission threshold) for $^{92}\text{Zr}(p,\gamma)^{93}\text{Nb}$ [23]. The data extend well into these ranges, allowing a detailed comparison to cross section predictions made for the calculation of astrophysical reaction rates.

The sensitivities of the cross sections and astrophysical S factors to variations of the averaged widths in the Hauser-Feshbach model are shown in Fig. 2a and 2b for the $^{90}\text{Zr}(p,\gamma)^{91}\text{Nb}$ and $^{92}\text{Zr}(p,\gamma)^{93}\text{Nb}$ reactions, respectively. The sensitivity Ω_{S_q} represents the impact of the variation of a model quantity q (e.g. nuclear level density or averaged widths) on the final

TABLE I: Cross sections and S-factors of the present work for the reaction $^{90}\text{Zr}(p,\gamma)^{91}\text{Nb}$.

| E_{beam} | $E_{c.m.}$ | σ | S-factor |
|------------|---------------|-----------------|--------------------|
| (keV) | (keV) | (mb) | (10^6 keV barn) |
| 5000 | 4925 ± 20 | 6.66 ± 0.99 | 1.65 ± 0.25 |
| 4800 | 4727 ± 20 | 5.77 ± 0.69 | 1.98 ± 0.24 |
| 4600 | 4529 ± 20 | 4.42 ± 0.60 | 2.15 ± 0.29 |
| 4400 | 4331 ± 21 | 3.65 ± 0.47 | 2.59 ± 0.33 |
| 4200 | 4132 ± 22 | 2.84 ± 0.39 | 3.01 ± 0.41 |
| 4000 | 3933 ± 23 | 1.80 ± 0.29 | 2.95 ± 0.47 |
| 3800 | 3735 ± 24 | 1.43 ± 0.18 | 3.73 ± 0.48 |
| 3600 | 3536 ± 24 | 0.97 ± 0.14 | 4.23 ± 0.60 |
| 3400 | 3338 ± 25 | 0.66 ± 0.08 | 5.00 ± 0.60 |
| 3200 | 3139 ± 26 | 0.39 ± 0.05 | 5.39 ± 0.68 |
| 3000 | 2940 ± 27 | 0.28 ± 0.04 | 7.54 ± 1.08 |
| 2800 | 2741 ± 28 | 0.17 ± 0.02 | 9.64 ± 1.19 |

result Ω (e.g. reaction rate, cross section or S -factor). The sensitivity is defined as:

$$\Omega_{S_q} = \frac{v_\Omega - 1}{v_q - 1} \quad (3)$$

where $v_\Omega = \Omega_{new}/\Omega_{old}$ is the change in Ω when the quantity q changes by a factor $v_q = q_{new}/q_{old}$. The sensitivity $\Omega_{S_q} = 0$ when no change occurs and $\Omega_{S_q} = 1$ when Ω changes by the same factor as q . More details about the sensitivity Ω_{S_q} can be found in [24].

In figures 2a and 2b it can clearly be seen that the proton width is determining the cross sections, and thus the reaction rates, at astrophysical energies. Above the astrophysical energy window, the cross sections become increasingly sensitive to the γ width and also to the neutron width once the energy is above the neutron-emission threshold. The neutron threshold for the $^{90}\text{Zr}(p,\gamma)^{91}\text{Nb}$ reaction is at 6.8 MeV and therefore the sensitivity to the neutron width is not observed in the energy region of Fig. 2a. For the $^{92}\text{Zr}(p,\gamma)^{93}\text{Nb}$ reaction, the neutron channel opens at 2.8 MeV with a clear increase in the neutron width sensitivity at that energy as shown in Fig. 2b.

TABLE II: Cross sections and S-factors of the present work for the reaction $^{92}\text{Zr}(p,\gamma)^{93}\text{Nb}$.

| E_{beam} | $E_{c.m.}$ | σ | S-factor |
|------------|---------------|-------------------|--------------------|
| (keV) | (keV) | (mb) | (10^6 keV barn) |
| 5000 | 4925 ± 22 | 0.209 ± 0.029 | 0.052 ± 0.007 |
| 4800 | 4726 ± 23 | 0.146 ± 0.023 | 0.050 ± 0.008 |
| 4600 | 4527 ± 24 | 0.140 ± 0.025 | 0.069 ± 0.012 |
| 4400 | 4329 ± 24 | 0.139 ± 0.021 | 0.099 ± 0.015 |
| 4200 | 4130 ± 25 | 0.117 ± 0.019 | 0.13 ± 0.02 |
| 4000 | 3931 ± 26 | 0.094 ± 0.016 | 0.15 ± 0.03 |
| 3800 | 3733 ± 27 | 0.107 ± 0.015 | 0.28 ± 0.04 |
| 3600 | 3534 ± 28 | 0.099 ± 0.015 | 0.44 ± 0.07 |
| 3400 | 3335 ± 29 | 0.072 ± 0.010 | 0.55 ± 0.08 |
| 3200 | 3136 ± 30 | 0.115 ± 0.018 | 1.62 ± 0.25 |
| 3000 | 2937 ± 31 | 0.484 ± 0.069 | 13.41 ± 1.91 |
| 2800 | 2738 ± 32 | 0.405 ± 0.075 | 23.82 ± 4.40 |
| 2600 | 2538 ± 34 | 0.215 ± 0.035 | 29.31 ± 4.83 |
| 2400 | 2339 ± 36 | 0.093 ± 0.012 | 32.70 ± 4.22 |
| 2200 | 2139 ± 38 | 0.036 ± 0.005 | 37.68 ± 5.38 |
| 2000 | 1939 ± 40 | 0.016 ± 0.002 | 57.22 ± 7.52 |

The cross section measurements of the present work are compared to theoretical calculations using the codes NON-SMOKER [25] and TALYS [26] in Figs. 3 and 4 for the $^{90}\text{Zr}(p,\gamma)^{91}\text{Nb}$ and $^{92}\text{Zr}(p,\gamma)^{93}\text{Nb}$ reactions, respectively. As mentioned in section I, for the reaction $^{90}\text{Zr}(p,\gamma)^{91}\text{Nb}$ there were two previous measurements that were not in agreement. The cross section measurement by Laird *et al.* [17] are presented in the main panel of Fig. 3 in open circles while the reaction rates presented by Roughton *et al.* [16] are shown in the inset of Fig. 3. In both Figs. 3 and 4 the solid line presents theoretical calculations from the code NON-SMOKER [25]. These calculations represent the suggested reaction rates in the JINA REACLIB database [15]. The dashed and dotted lines in both figures present theoretical calculations performed with the code TALYS [26]. The TALYS calculations were

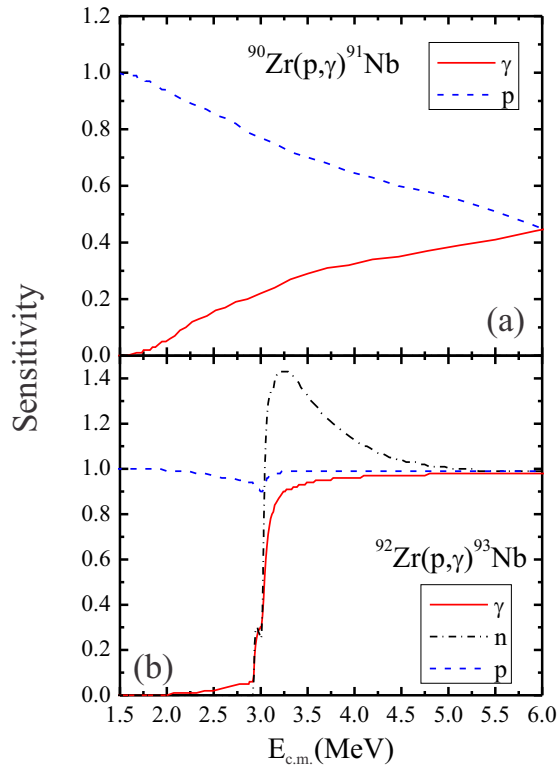


FIG. 2: (Color Online) Sensitivity of the reaction $^{90}\text{Zr}(p,\gamma)^{91}\text{Nb}$ (a) and $^{92}\text{Zr}(p,\gamma)^{93}\text{Nb}$ (b) to the various widths in the Hauser-Feshbach model. The solid (red) line corresponds to the sensitivity to the γ width, the dashed (blue) line shows the sensitivity to the proton width and the dot-dashed (black) line shows the sensitivity to the neutron width.

performed using default parameters, except for the Optical Model Potential (OMP) and the Nuclear Level Densities (NLD). The semi microscopic OMP of Bauge, Delaroche, and Girod [27] was used, which is a reparametrization of the standard microscopic OMP of Jeukenne, Lejeune, and Mahaux (JLM) [28] to cover a wider energy range of scattering data. The JLM potential was used in the NON-SMOKER calculations presented here. TALYS offers a variety of options for the description of the Nuclear Level Densities. In the present work, the use of the different NLDs gave similar results and for this reason only the lowest and highest cross section calculations are presented in the figures, as limits of these calculations. For both reactions the highest limit (dashed lines) was the result of using the NLD from the Generalized Superfluid Model (GSM) [29, 30], while the lower limit (dotted lines) was calculated using the microscopic NLD from Hilaire's table (HT) [31].

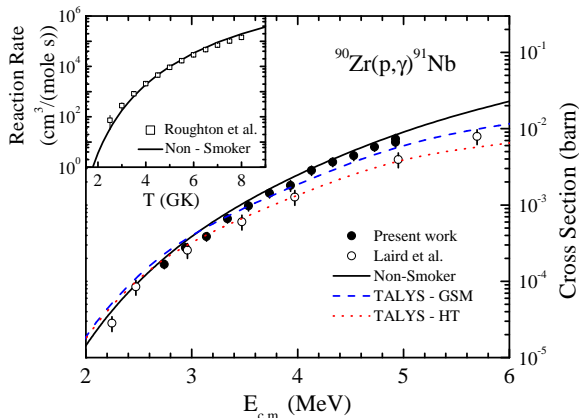


FIG. 3: (Color Online) Main figure: cross section of the reaction $^{90}\text{Zr}(p,\gamma)^{91}\text{Nb}$ from the present work (black dots) and from Laird *et al.* [17] (open circles). Standard NON-SMOKER theoretical calculations [25] are shown in the solid (black) line and two TALYS calculations [26] using different Nuclear Level Densities (NLD) are shown in the dashed (blue) and dotted (red) lines. See text for details. The inset shows a comparison of the reaction rate extracted by Roughton *et al.* [16] (open squares) compared to the standard NON-SMOKER calculation.

In both reactions, a good agreement between the results of the present work and the three calculations shown in Figs. 3 and 4 is observed at the lowest energies, in the regions where the sensitivity is dominated by the proton width. The good agreement with the NON-SMOKER calculation remains also at higher energies, with significant deviations only present in the $^{92}\text{Zr}(p,\gamma)^{93}\text{Nb}$ reaction above the neutron threshold. Larger discrepancies are observed when comparing to the calculations using the TALYS code, although the calculation using the GSM nuclear level densities can reproduce the experimental data fairly well. Overall, NON-SMOKER yields a better description of the data for both reactions. Since, it reproduces the experimental data well within the uncertainties, the astrophysical reaction rates obtained with NON-SMOKER are also confirmed. Therefore we refrain from giving rate tables but just refer to [25, 32].

In comparing the results of the present work with previous experimental results, it is observed that the reaction rates obtained by Roughton *et al.* [16] shown in the inset of Fig. 3 are in good agreement with the standard NON-SMOKER calculations which can describe the results of the present work as well. A larger discrepancy is observed between the present

results and the ones from Laird *et al.* [17], especially at the higher energies. As mentioned in section I, the latter authors used the method of γ angular distribution measurements to extract the reaction cross section. This method relies on the detection of all individual γ rays contributing to the cross section and presents the risk of missing low intensity γ rays that are below the detection (and background) limits of the experimental setup. With the technique used in the present work we were able to extract average γ multiplicities, $\langle M \rangle$, as described in [19]. For the $^{90}\text{Zr}(p,\gamma)^{91}\text{Nb}$ reaction, the average multiplicity increases with beam energy getting to a value of approximately $\langle M \rangle = 3$ for proton beam energy of 5 MeV. This result supports the observation that the γ -angular distribution data of Laird *et al.* are in good agreement with our measurement at the low beam energies and the deviation increases as the beam energy and average multiplicity are increased.

The astrophysical calculations performed in the present work, as outlined in Sec. I, did not show any significant contribution of the studied reactions to the production of the p nuclei $^{92,94}\text{Mo}$. Since the results confirmed the NON-SMOKER calculations that were already used in the astrophysical model, no change was observed in the final abundances and the puzzle of the underproduction of the neutron deficient molybdenum and ruthenium isotopes remains.

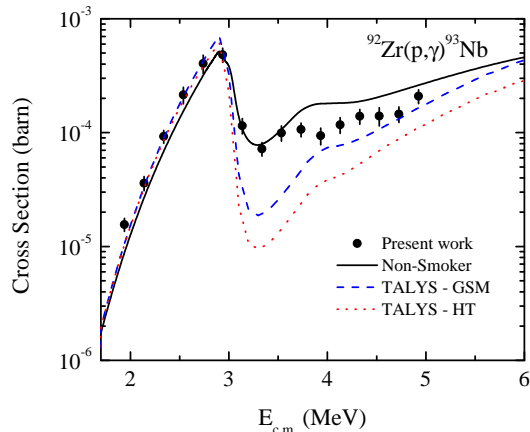


FIG. 4: (Color Online) Cross section of the reaction $^{92}\text{Zr}(p,\gamma)^{93}\text{Nb}$ from the present work (black dots) compared to three theoretical calculations. Standard NON-SMOKER calculations [25] are shown in the solid (black) line and two TALYS calculations [26] using different Nuclear Level Densities (NLD) are shown in the dashed (blue) and dotted (red) lines. See text for details.

V. CONCLUSIONS

The present work reports on two cross section measurements of the reactions $^{90}\text{Zr}(p,\gamma)^{91}\text{Nb}$ and $^{92}\text{Zr}(p,\gamma)^{93}\text{Nb}$. The measurements were performed using the NSCL SuN detector at the University of Notre Dame in the energy range between 2.0 and 5.0 MeV. This energy range covers the majority of the Gamow window for the two reactions. The results were found to be in good agreement with one of the previous measurements for the $^{90}\text{Zr}(p,\gamma)^{91}\text{Nb}$ reaction. The results were also in excellent agreement with the predictions of the code NON-SMOKER, the calculations of which are commonly used in astrophysical models. Somewhat larger discrepancies were observed when using the code TALYS for the cross section calculations. This work provides a confirmation of the reaction rates provided by the code NON-SMOKER and used in the JINA REACLIB database for the two reactions studied. However, no generalization for other reactions can be inferred from the present work.

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