

# T Tauri candidates and accretion rates using IPHAS: method and application to IC 1396

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## ABSTRACT

The INT Photometric H-Alpha Survey (IPHAS) is a 1800 deg<sup>2</sup> survey of the Northern Galactic Plane, reaching down to  $r' \sim 21$ . We demonstrate how the survey can be used to (1) reliably select classical T Tauri star candidates and (2) constrain the mass accretion rates with an estimated relative uncertainty of 0.6 dex. IPHAS is a necessary addition to spectroscopic surveys because it allows large and uniform samples of accretion rates to be obtained with a precise handle on the selection effects.

We apply the method on a region of 7 deg<sup>2</sup> towards the HII region IC 1396 in Cepheus OB2 and identify 158 pre-main sequence candidates with masses between 0.2 and 2.0  $M_{\odot}$  and accretion rates between  $10^{-9.2}$  and  $10^{-7.0} M_{\odot} \text{yr}^{-1}$ . We find a power-law dependency between the stellar mass and the accretion rates with a slope of  $\alpha = 1.1 \pm 0.2$ , which is less steep than indicated by previous studies. We discuss the influence of method-dependent systematic effects on the determination of this relationship.

The majority of our sample consists of faint, previously unknown, low-mass T Tauri candidates (56 per cent between 0.2 and 0.5  $M_{\odot}$ ). Many candidates are clustered in front of three bright-rimmed molecular clouds, which are being ionized by the massive star HD 206267 (O6.5V). We discover a spatio-temporal gradient of increasing accretion rates, increasing Spitzer infrared excess, and younger ages away from the ionizing star, providing a strong indication that the formation of these clusters has been sequentially triggered by HD 206267 during the last  $\sim 1$  Myr.

**Key words:** stars: pre-main sequence, accretion, open clusters and associations: individual: IC 1396

## 1 INTRODUCTION

In the current picture of star formation, young solar-like stars grow by accreting gas from a circumstellar disc (Hartmann 1998). Whilst this picture is widely supported by observations and theory (e.g. Appenzeller & Mundt 1989; Hartmann et al. 1994), there remain considerable uncertainties with respect to the time-scales and mechanisms responsible for the removal of circumstellar disc material (e.g. Mayne & Naylor 2008), with fundamental implications for planet-formation models (Hester et al. 2004; Adams et al.

2004; Throop & Bally 2005; Hollenbach et al. 2005; Hillenbrand 2005). One of the main reasons why progress has been slow is that disc evolution is thought to be driven by a complex interplay between several physical effects, involving magnetospheric accretion, UV-photoevaporation, stellar winds, cosmic ray and X-ray ionization, dust coagulation, planet formation, and binary interactions (Hillenbrand 2008).

Another hindrance to progress is the distance of typical star forming regions. Most studies to date have concentrated on relatively small numbers of objects (several tens) in nearby regions (e.g., Taurus, Ophiuchus and Chamaeleon at  $\sim 150$  pc), because this is where high-quality spectra can

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be obtained. These regions cover only a limited range of stellar ages, and are not necessarily representative for the massive OB associations where the majority of solar-type stars are thought to form (e.g. Cygnus OB2, Wright et al. 2010). In fact, there is increasing evidence to suggest that our Solar System formed near a massive star (e.g. Hester & Desch 2005). Unfortunately nearly all OB regions are located beyond 500 pc, where high-resolution spectroscopy of large numbers of stars with solar-like masses becomes increasingly challenging (see e.g. Vink et al. 2008).

Whilst large and unbiased statistical samples of young stars across star-forming regions are required for a significant breakthrough in our understanding of star formation, such samples have only recently become available through infrared photometry (e.g. Evans et al. 2009) which traces the circumstellar *dust*. However, large samples are not yet available for optical (emission-line) studies which traces material in the *gas* phase, which dominates the circumstellar environment.

Over the past decade, gas accretion rates have been derived from UV and optical continuum emission produced by the impact shock (e.g. Gullbring et al. 1998; Herczeg & Hillenbrand 2008), as well as the shape and intensity of emission lines produced in the accreting gas (e.g. Hartigan & Kenyon 2003; Natta et al. 2004; Mohanty et al. 2005; Dahm 2008). The available measurements do not provide a particularly clear picture: at any given age, stars are observed with widely different accretion rates and disc properties (Calvet et al. 2000; Muzerolle et al. 2003; Sicilia-Aguilar et al. 2006b, 2010; Currie et al. 2009). For example, whilst roughly 50 per cent of stars appear to end accretion and lose their inner discs by the age of 1 Myr, some small fraction of stars appear to continue accreting beyond 10 Myr (Haisch et al. 2001; Armitage et al. 2003; Mamajek et al. 2004).

De Marchi et al. (2010) recently proposed a strategy to obtain statistical samples of mass accretion rates without the need for spectroscopy. The authors combined HST broadband V/I photometry with narrow-band H $\alpha$  imaging to select a large number of pre-main sequence candidates in the Large Magellanic Cloud, and to determine their mass accretion rates. Although the lack of spectroscopic follow-up might compromise the accuracy of their results for individual objects, their work produced a large and homogeneous sample which allows a sound statistical appraisal to be made.

In this paper, we introduce a similar approach based on r'/i'/H $\alpha$  photometry obtained from the IPHAS survey of the Northern Galactic Plane. In brief, we set-out to determine mass accretion rates from the (r'-i'/r'-H $\alpha$ ) colour-colour diagram, whilst simultaneously estimating stellar ages and masses from the (r'/r'-i') colour-magnitude diagram. Because we consider objects in a star-forming region for which some spectroscopic results are already in the literature, we can examine how our photometric approach fares in comparison with more conventional spectroscopic methods.

As a testbed, we study the HII region IC 1396, ionized by the O-star HD 206267. A comprehensive review of IC 1396 and its associated open cluster Trumpler 37 is given in Kun et al. (2008). In brief, the region is located at the edge of the Cepheus OB2 association at a distance of  $\sim 900$  pc (Contreras et al. 2002). It has previously been the target of photographic H $\alpha$  objective prism surveys (Kun 1986; Kun & Pasztor 1990; Balazs et al. 1996). However, these sur-

veys were not deep enough ( $R \lesssim 15-16$ ) to uncover the low-mass T Tauri population ( $R > 16$ ), because slit-less spectroscopy becomes increasingly impractical at faint magnitudes in crowded fields. The IPHAS digital narrow-band imaging is better suited for this purpose.

The central part of IC 1396 has been the target of an optical survey by Sicilia-Aguilar et al. (2004, 2005, 2006a,b, 2010), who uncovered a rich population of T Tauri stars using photometric candidate selection combined with follow-up spectroscopy. Our work extends this existing survey, as our photometric approach allows us to probe 2 magnitudes deeper and study an area that is a factor four times larger (including some of the unexplored outskirts of IC 1396).

We identify 158 accreting pre-main sequence candidates from a database of  $\sim 1$  million objects, with a contamination rate of less than 15 per cent and well-defined detectability limits. The sample allows us to investigate the dependency of accretion rates on stellar age and mass, and shows a spatio-temporal gradient which, we will argue, provides evidence for sequentially triggered star formation.

The paper is organized as follows: in §2 we introduce the IPHAS survey and explain how T Tauri candidates are selected from the colour-colour diagram. In §3 we show how the ages, masses and accretion rates are determined, and in §4 we explain the reduction of archival infrared photometry. In §5 we present the results and in §6 we use existing literature to show that the vast majority of our candidates are genuine T Tauri stars. A careful analysis of the errors is presented in §7. Finally, in §8 we discuss the implications of our results and §9 contains the conclusions.

## 2 METHOD: SELECTING T TAURI CANDIDATES USING IPHAS COLOURS

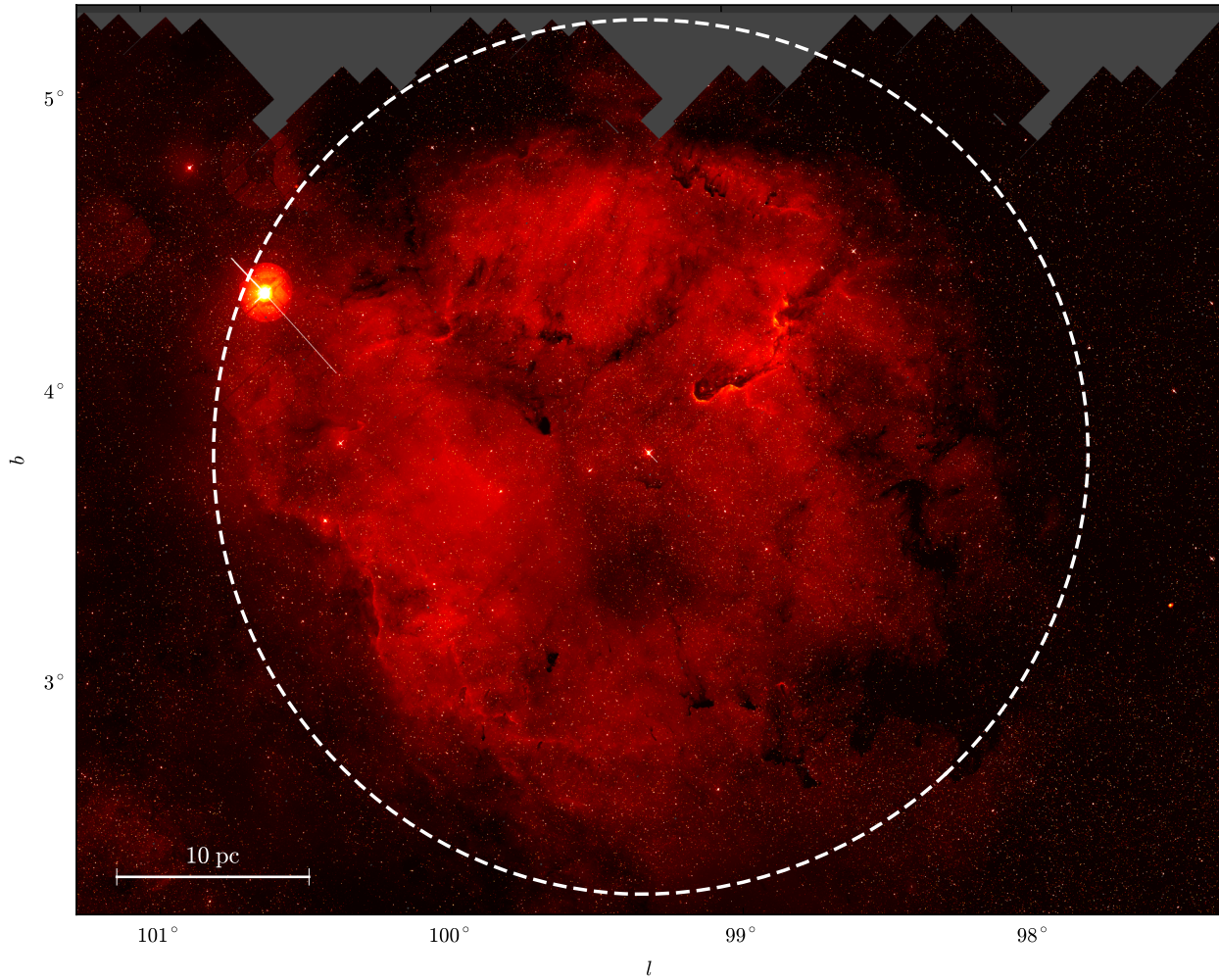
### 2.1 IPHAS observations of IC 1396

A comprehensive overview of the IPHAS survey is given in Drew et al. (2005); González-Solares et al. (2008) and on the project website<sup>1</sup>. Briefly, IPHAS is a 1800 deg<sup>2</sup> photometric survey of the Northern Galactic Plane ( $30^\circ \lesssim \ell \lesssim 220^\circ$ ,  $-5^\circ \lesssim b \lesssim +5^\circ$ ), carried out using the Wide Field Camera (WFC) on the 2.5-meter Isaac Newton Telescope (INT) in La Palma. Photometry was obtained using a narrow-band H $\alpha$  filter and the broadband Sloan r' and i' filters. All data are then pipeline processed at the Cambridge Astronomical Survey Unit (CASU; Irwin 1985, 1999). This routinely includes photometric calibration based on nightly observations of standard star fields. However we have been able to draw upon the results of a first-pass global calibration of IPHAS survey data (Miszalski et al, in preparation), thereby significantly reducing field-to-field magnitude shifts.

The observations of IC 1396 were obtained during several bright but mostly photometric nights in August 2004. Fields that suffered from nights of poor seeing were repeated in subsequent years (2005-2009), resulting in a final average seeing of  $0.9 \pm 0.1''$ . We used the MONTAGE<sup>2</sup> toolkit to create an H $\alpha$  mosaic of the 120 IPHAS fields towards the region,

<sup>1</sup> <http://www.iphas.org>

<sup>2</sup> <http://montage.ipac.caltech.edu>



**Figure 1.** H $\alpha$ -mosaic of IPHAS observations towards IC 1396 at the edge of Cepheus OB2. The white dashed line shows the region studied in this paper, with a radius of 1.5° centered on the massive star system HD 206267 (O6.5V), which is the main ionizing source of the region. Grey areas near the top indicate the upper edge of the IPHAS survey.



**Figure 2.** Close-up color mosaic of IC1396a ( $l = 99^\circ, b = +4^\circ$ ), commonly known as the Elephant's Trunk Nebula, composed from the H $\alpha$  (red channel),  $r'$  (green) and  $i'$  (blue) bands. Red and yellow colors indicate regions with the most intense H $\alpha$  emission.

shown in Fig. 1. A close-up of the Elephant’s Trunk Nebula, which is part of the region, is shown in Fig. 2.

The contrast of the mosaic has been stretched to bring out the strong  $H\alpha$  background emission in the region. We draw the reader’s attention to the ability of the CASU pipeline to track variations in the background emission on scales of 20-30 arcsec. This allows the stellar flux to be accurately separated from the background in the vast majority of cases. Any objects which, nevertheless, suffer from incorrect background subtraction due to very rapidly spatially varying regions of high nebulosity are flagged as part of the pipeline’s morphological classification. In brief, the pipeline derives a curve-of-growth for each detected object from a series of aperture flux measures with different aperture radii. In the case of incorrect background correction, the curve-of-growth will not show the asymptotic behaviour expected of stellar sources and so will be classified as non-stellar. We confirmed the effectiveness of this feature using manual inspection of pipeline photometry in multiple nebulous regions.

## 2.2 Quality criteria

The IPHAS database contains more than 1 million photometric detections in our region of study (indicated by the dashed circle in Fig. 1). We narrowed down the catalogue to 404 975 sources using the following strict quality requirements:

- (i) the magnitude must be in the range  $13 < r' < 20$ ;
- (ii) the star must be morphologically classified as strictly stellar in all bands (“class = -1”);
- (iii) the difference in coordinates between the matched detections in the  $H\alpha$ ,  $r'$  and  $i'$  images must not be bigger than  $0.1''$ .

Criterion (i) avoids faint and saturated sources. Although the IPHAS photometry goes down to a  $3\sigma$ -depth of  $r' = 21 \pm 1$ , we choose to limit our study to  $r' < 20.0$  to keep the photometric uncertainties small and to avoid spurious detections.

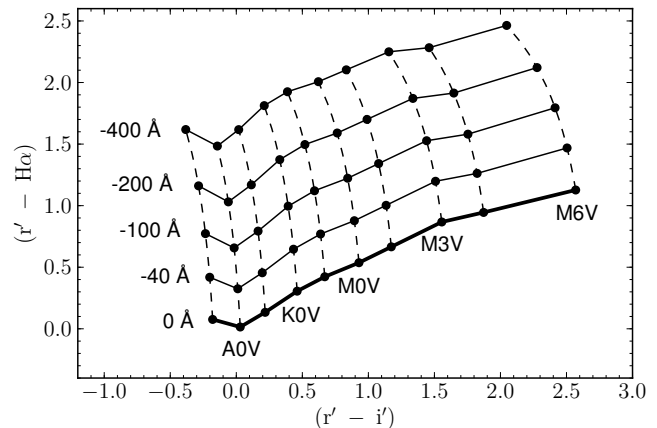
Criterion (ii) avoids saturated and extended objects or problems resulting from flux-contamination by nearby neighbours. The criterion also deals with the issue of nebulous  $H\alpha$  background emission. As noted in the previous section, background-contaminated photometry will be flagged as non-stellar and removed this way. We note that this has the effect of removing a small number of interesting – but unreliable – objects.

Criterion (iii) avoids spurious mismatches in crowded fields. The CASU pipeline allows matches up to 2.5 arcseconds apart between different filters. This criterion needs to be tightened considerably to avoid mismatches near the faint limit of the survey.

## 2.3 Quantifying the $H\alpha$ excess from IPHAS colours

### 2.3.1 Colour simulations

As described by Drew et al. (2005), the  $(r'-H\alpha)$  colour of a point source is an indicator of the strength of the  $H\alpha$  line relative to the red continuum. This may flag the object as an emission line star or, for most stars, provide a measure



**Figure 3.** Simulated colours for  $E(B-V)=0$  (unreddened stars) with increasing levels of  $H\alpha$  emission. The tracks are tabulated for different amounts of reddening in Appendix A.

of intrinsic colour. The  $(r'-i')$  colour reflects a combination of interstellar reddening and intrinsic colour.

Stars with  $H\alpha$  in emission show higher  $(r'-H\alpha)$  values than other stars of a similar spectral type. However, there is no straightforward analytical relationship between the  $(r'-H\alpha)$  colour and the strength of the  $H\alpha$  line, in part because the  $H\alpha$  band falls inside the  $r'$  band. A given amount of  $H\alpha$  excess emission will affect both magnitudes in a way that depends on the underlying SED and reddening of the star.

For the purpose of estimating the  $H\alpha$  excess from IPHAS data, we created a set of simulated colour tracks for stars with increasing levels of  $H\alpha$  emission equivalent width (EW) and different underlying SEDs. A similar grid was presented by Drew et al. (2005, table 4) but did not extend down in effective temperature to stars of K and M type (common among T Tauri stars).

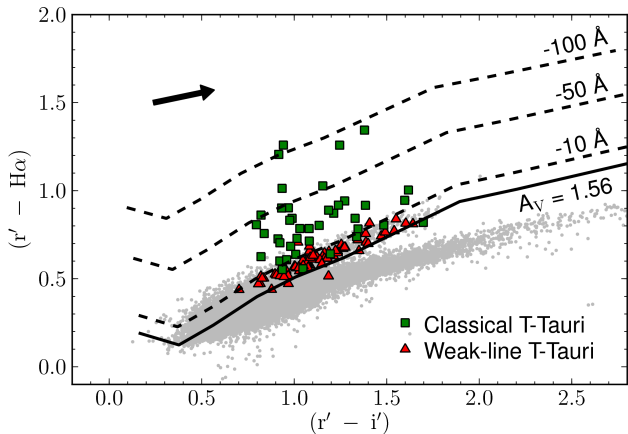
To simulate the colours we adopted the same procedure as described in detail by Drew et al. In brief, colours are generated from the library of stellar SEDs due to Pickles (1998), using a standard extinction law  $R_V = 3.1$  in the form given by Howarth (1983), and taking into account the relevant filter transmission profiles and the CCD response curve available from the WFC website<sup>3</sup>. We adopted as our reference Vega spectrum that due to Bohlin (2007). The resulting grid is tabulated in Appendix A and plotted in Fig. 3.

### 2.3.2 Validation of the simulated colours against spectroscopic observations of T Tauri stars

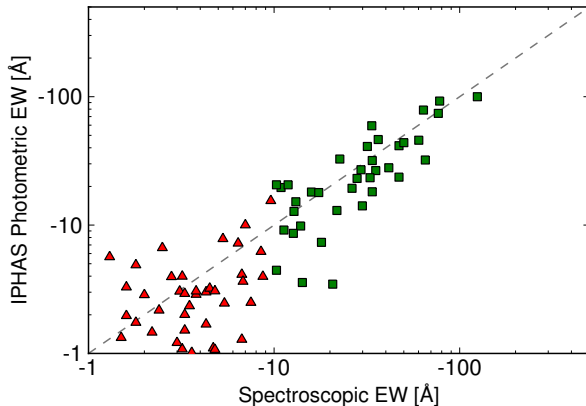
To test our grid of simulated colours, we took the sample of known T Tauri stars in IC 1396 from Sicilia-Aguilar et al. (2005, hereafter referred to as SA05). In their work, the authors used MMT/Hectospec multifiber spectroscopy to follow-up candidates selected using multi-epoch broadband photometry. The sample includes both weak-line T Tauri stars (WTTS) with  $H\alpha$  emission weaker than  $-10 \text{ \AA}$  EW, as

<sup>3</sup> <http://www.ing.iac.es/Astronomy/instruments/wfc/>





**Figure 4.** IPHAS colours of known T Tauri stars in IC 1396 from Sicilia-Aguilar et al. (2005). Green squares are classical T Tauri stars (CTTS) with spectroscopic  $EW_{H\alpha}$  stronger than  $-10 \text{ \AA}$ . Red triangles are weak-line T Tauri stars (WTTS) with  $EW_{H\alpha}$  weaker than  $-10 \text{ \AA}$ . The solid line shows the simulated main sequence curve at the mean reddening of the cluster ( $A_V = 1.56$ ). Dashed lines show the position of stars at increasing levels of  $H\alpha$  emission as predicted by our simulated tracks. Grey dots show field stars in the region. The arrow shows the reddening shift for an M0V-type object being reddened from  $A_V = 0$  to  $A_V = 1.56$  (note that the true reddening tracks are curved in a way that depends on the SED and the amount of reddening, see Drew et al. 2005)



**Figure 5.** Comparison of IPHAS photometric  $EW_{H\alpha}$  with spectroscopic values from Sicilia-Aguilar et al. (2005). The grey dashed line shows the unity relation. The scatter is thought to be dominated by natural  $H\alpha$  emission variability.

well as classical T Tauri stars (CTTS) with emission stronger than  $-10 \text{ \AA}$  EW.

Out of the 118 probable T Tauri members confirmed by Hectospec observations, 109 have a reliable counterpart in the IPHAS database. (The other 9 objects were classified as extended objects by the pipeline; 6 objects because they appear to suffer from flux contamination due to a very close neighbour, while 3 objects appear to suffer from a rapidly spatially varying background.)

In Fig. 4 we show the IPHAS colours of these 109 objects superimposed on the grid of simulated lines of constant  $H\alpha$  EW referred to an appropriately reddened main sequence (zero line emission at  $A_V = 1.56 \pm 0.55$ , which

is the typical reddening found by SA05). We notice that the classical T Tauri stars are well separated from the field stars (shown as grey points): most are above the  $-10 \text{ \AA}$  EW boundary as predicted by the drawn grid lines. In contrast the weak-lined stars fall within the main stellar locus, blending in with normal less-reddened stars. The fact that reddening raises the EW threshold for the clean detection of emission line stars is a recognised property of the IPHAS colour-colour plane (see Drew et al. 2005).

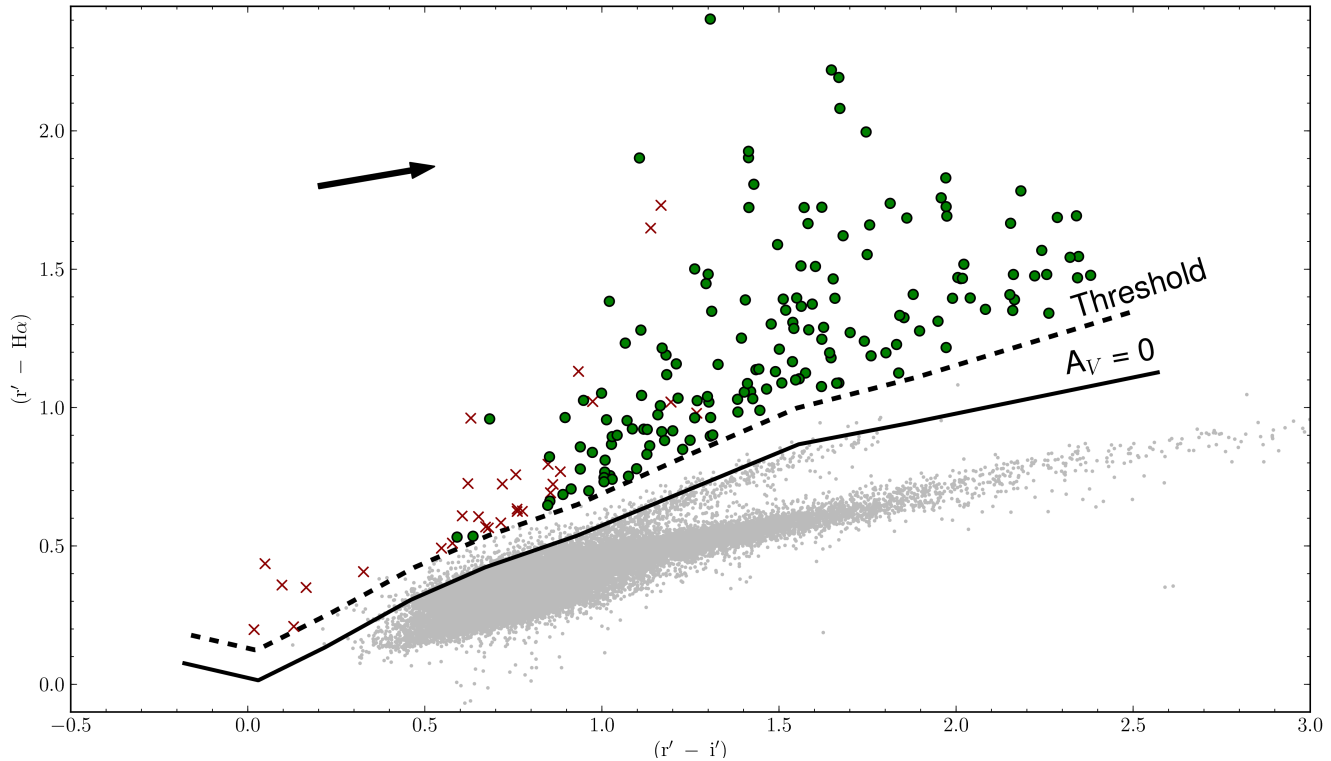
One weak-lined object, named 73-537 in SA05, can be seen to fall somewhat below the simulated main sequence (solid line in Fig. 4). The aberrant position is explained by the high reddening of the object,  $A_V = 3.3$ , which is an outlier in terms of reddening compared to the rest of the sample.

To validate the grid in more detail, we interpolated the tracks to derive  $H\alpha$  EWs for the known T Tauri objects. These values are then plotted against the spectroscopic values from SA05. The comparison is shown in Fig. 5. We find a strong correlation between the photometric and spectroscopic estimates, albeit with a large scatter on the order of 5 to  $10 \text{ \AA}$ .

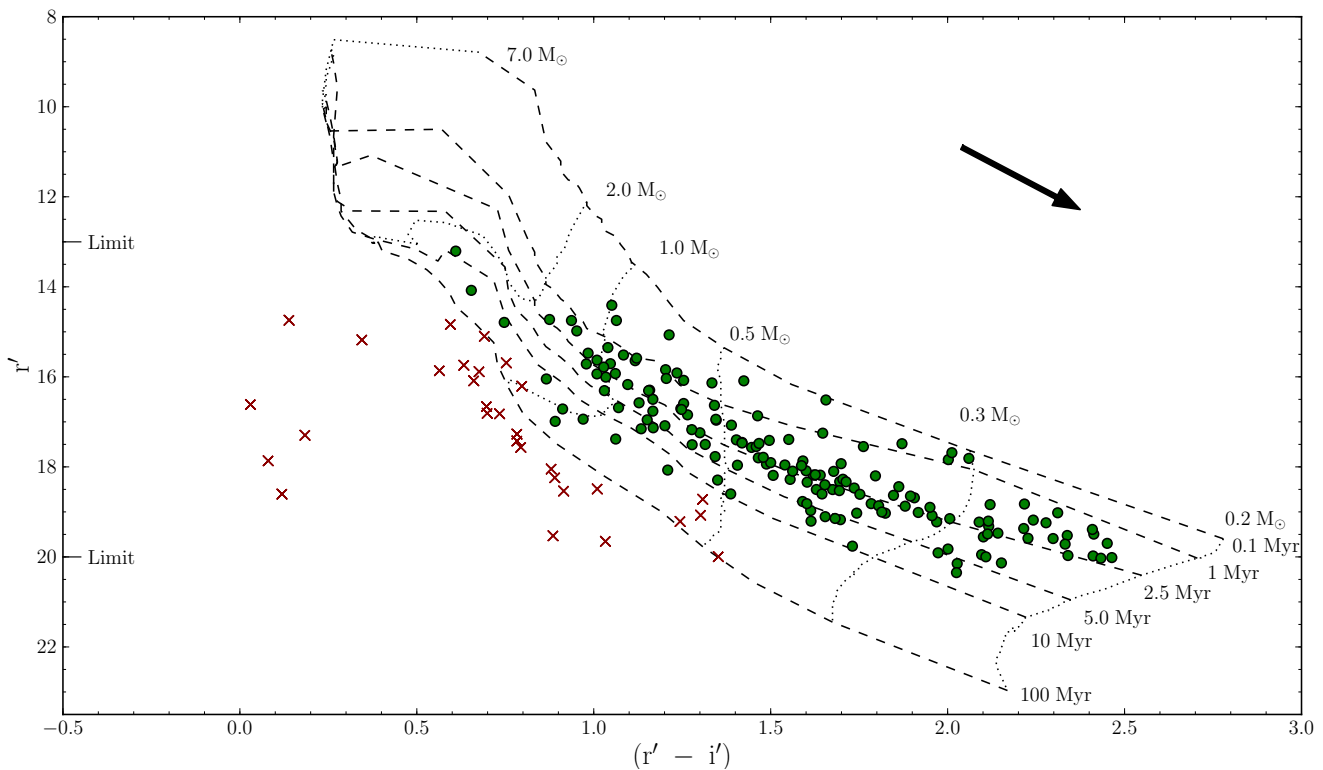
The scatter is explained by two reasons. First, photometric errors and reddening deviations will introduce large relative errors for weak-lined objects which are positioned close to the reddened main sequence line. This dominates the scatter observed for objects with lines weaker than  $-10 \text{ \AA}$ . Second, net  $H\alpha$  emission in T Tauri stars is known to show large natural variations on timescales as short as days (e.g. Fernandez et al. 1995; Alencar et al. 2001, Costigan et al. in preparation). For a subset of the objects considered here, Sicilia-Aguilar et al. (2006a, hereafter referred to as SA06) obtained high-resolution MMT/Hectospec spectroscopy at a second epoch. Compared to the Hectospec data, the EWs show a scatter of 33 per cent for emission EWs exceeding  $-20 \text{ \AA}$  (fig. 10 in SA06). In our sample we observe a similar scatter of 37 per cent for lines stronger than  $-20 \text{ \AA}$  which suggests that the spread is dominated by intrinsic variations.

We confirmed this by recomputing photometric equivalent widths, this time taking into account the individual reddening values  $A_V$  for each object from SA05, and find that the scatter decreased only marginally. Knowledge of the individual reddening values does not significantly improve the estimated line strengths here, because the spread in reddening is reasonably small for the T Tauri objects in IC 1396 ( $\sigma_{A_V} = 0.55$ , SA05).

We conclude that the simulated grid and IPHAS photometry can be used to identify emission-line objects and constrain their  $H\alpha$  excess with a sufficient degree of confidence for the purpose of our work, which is to constrain the  $H\alpha$  excess for a very large sample of stars - over a million objects in the field towards IC 1396 - in a homogeneous way which is not biased by a pre-selected list of observing targets. To study a similar amount of objects using traditional spectroscopy would require so much observing time that it is effectively impossible.



**Figure 6.** Results of the T Tauri candidate selection towards IC 1396. Green circles show accepted candidates. Red crosses show stars which were rejected on the basis of their position in the  $(r' / r'-i')$  colour-magnitude diagram (Fig. 7). The solid line shows the unreddened main sequence curve from our grid of simulated colours, while the dashed line shows the selection threshold which serves to remove chromospherically active foreground dwarfs (see text). Grey dots show 50 000 random field stars having accurate photometry ( $\sigma_{r'} < 0.01$ ). The arrow shows the reddening shift for an M0V-type object being reddened from  $A_V = 0$  to  $A_V = 1.56$ .



**Figure 7.** Position of the accepted (green circles) and discarded (red crosses) candidates in the colour-magnitude diagram. Dashed and dotted lines show isochrones and mass tracks from the Siess et al. (2000) model. The tracks have been placed at the mean reddening and distance modulus of the cluster, and the stars have been corrected for the  $r'$ -band excess due to the  $H\alpha$  line excess. Objects are discarded if they are not located above the 100 Myr isochrone with  $3\sigma$  confidence (see text). The arrow shows the reddening vector for  $A_V = 1.56$

## 2.4 Selection strategy

### 2.4.1 *T Tauri selection threshold*

We now make a homogeneous selection of T Tauri candidates based on their position above the main locus of field stars in the  $(r'-i')/r'-H\alpha$  diagram. This requires a selection threshold to be established relative to the upper edge of the main locus (cf. grey dots in Fig. 4), which is where unreddened foreground main-sequence stars are located.

The threshold is chosen such that the foreground stars are excluded, even when they are chromospherically active. We adopt the traditional CTTS threshold which requires EWs stronger than  $-10 \text{ \AA}$  for early- and intermediate-type stars. For late-type stars, which are known to show intense chromospheric activity long after their formation, we adopt the empirical thresholds found by Barrado y Navascues & Martin (2003) ranging from  $-11 \text{ \AA}$  for spectral type M2V ( $r'-i' \cong 1.2$ ) up to  $-24 \text{ \AA}$  for type M6V ( $r'-i' \cong 2.6$ ). The resulting threshold is shown as a dashed line in Fig. 6.

We stress that the threshold is drawn relative to the *unreddened* main sequence track. Stars at the typical reddening of IC 1396 will naturally be located slightly lower in the diagram, and will need to show an excess stronger than roughly  $-25 \text{ \AA}$  at all spectral types to fall above the threshold. This level of excess is the completeness limit of our work, which is to be discussed in §6.

We then select all sources which are located above the threshold at the  $3\sigma$  confidence level:

$$(r' - H\alpha)_{\text{star}} - (r' - H\alpha)_{\text{threshold}} > 3\sigma, \quad (1)$$

where  $\sigma$  is determined from the photometric uncertainty:

$$\sigma = \sqrt{\sigma_{(r'-H\alpha)}^2 + m^2 \cdot \sigma_{(r'-i')}^2}, \quad (2)$$

with  $m$  the local slope of the threshold ( $\bar{m} = 0.5$ ).

Fig. 6 shows the result of the selection procedure. We find a total of 188 unique objects located confidently above the selection threshold. In the case of multiple photometric detections due to overlapping or repeated CCD pointings, we averaged the photometry and show the object only once. We note that no extremely low-mass candidates are found beyond  $(r'-i') > 2.5$ , which is in agreement with the low density of such objects in IPHAS as previously found by Valdivielso et al. (2009).

### 2.4.2 *Rejecting candidates on the basis of colour-magnitude*

T Tauri stars are not the only  $H\alpha$ -emitting objects in the Galactic Plane. Other stellar objects that may show emission include evolved massive stars (e.g., Be stars, Wolf-Rayet stars, luminous blue variables), evolved intermediate-mass stars (e.g., Mira variables, unresolved planetary nebulae), and interacting binaries (e.g., cataclysmic variables, symbiotic stars). We refer to Corradi et al. (2008) for a discussion on the position of these objects in the IPHAS diagram.

Although it is not possible to confirm T Tauri stars solely on the basis of IPHAS photometry, these other classes of objects show significantly lower surface densities towards known star-forming regions, where faint red young stars have been shown to dominate the local IPHAS diagrams (Witham

et al. 2008, figs. 2-3). Moreover, the scarce but often luminous nature of these objects implies that they will most likely appear as distant background objects, at large distance moduli that may push them below the sequence of foreground T Tauri stars in a colour-magnitude diagram.

This effect is demonstrated in Fig. 7, where we placed the selected candidates in the  $(r'/r'-i')$  diagram. The objects are superimposed on the model evolutionary tracks and isochrones due to Siess et al. (2000), which are placed at the mean reddening and distance modulus of the region (discussed in the next section). A number of blue objects are seen to fall near or below the 100 Myr isochrone (which is a good approximation of the ZAMS for solar mass stars). Their position cannot be explained by extinction, because the reddening vector runs nearly parallel to the isochrones. Instead, these objects are likely to be background objects not related to IC 1396.

We remove these suspicious objects from our sample by requiring that all T Tauri candidates are located confidently ( $3\sigma$ -level) above the 100 Myr isochrone for the region. From our 188 initial candidates, 30 are discarded on this basis (shown as red crosses). We note that this criterion, combined with the faint magnitude limit ( $r' < 20$ ), will not allow us to detect the older population for objects below  $\sim 0.4 M_{\odot}$ , and will somewhat bias the survey towards the nearest part of IC 1396 (where the distance modulus is smaller)

This leaves us with a final sample of 158 T Tauri candidates (green circles), which are listed in Appendix B1 together with their photometry. In §6 we will show that the majority of these candidates are likely to be genuine members because they show spatial clustering and infrared colours consistent with T Tauri objects.

In comparison with the list of confirmed T Tauri members from SA06, we have 34 objects in common (mostly at bright magnitudes), while 61 objects were not previously discovered (mostly at faint magnitudes). The remaining 63 objects fall entirely outside the area that was studied by SA06. Only 3 objects showing emission stronger than  $-30 \text{ \AA}$  EW in SA06 were not recovered by our selection (to be discussed in §6).

In contrast, none of the discarded objects are clustered or appear in the work by SA06. These objects are listed separately in Appendix E1 because they are likely to be interesting emission-line objects, but they are not considered in what follows.

## 3 ANALYSIS: STELLAR AND ACCRETION PROPERTIES

Having identified a population of T Tauri candidates, we now estimate the ages, masses and accretion rates. The principal assumption made in this step is that all objects are placed at the mean distance  $\bar{d} = 870 \pm 80 \text{ pc}$  (Contreras et al. 2002) and mean reddening  $\bar{A}_V = 1.56 \pm 0.55$  (SA05).

The assumption of a fixed reddening is reasonable because SA05 found that the statistical distribution of  $A_V$  for T Tauri stars in the region is nearly Gaussian, with  $\sim 76$  per cent of the objects in the  $\pm 1\sigma$  region, and only three objects – related to the globules – showing extinctions beyond  $+3\sigma$ . Moreover, SA05 did not find an evident spatial variation of the extinction. This can be considered consistent with the

earlier finding by Morbidelli et al. (1997) that most of the reddening towards optical stars in IC 1396 is fairly uniform and of foreground origin.

Because the distribution of reddening values towards IC 1396 is constrained, the uncertainties that follow from the assumption of a fixed reddening are well-understood and will be discussed in §7.

### 3.1 Age, mass and radius

We estimate ages, masses and radii from the ( $r'/r'-i'$ ) diagram in Fig. 7 using the evolutionary tracks by Siess et al. (2000). We adopt the Siess tracks because they were found to be most consistent in a comparative study between different models by Lawson & Feigelson (2001). Moreover, these tracks were also adopted in SA05, so by adopting these tracks we are able to compare our results more directly with the existing study.

The Siess model tracks are available for the Cousins ( $R_c/R_c - I_c$ ) diagram based on temperature-colour conversions due to Kenyon & Hartmann (1995). In the absence of temperature-colour conversions for the Vega-based photometric system used by IPHAS, we adopt the tracks for Cousins and convert them to ( $r'/r'-i'$ ) using the transformations as determined for the WFC by the Cambridge Astronomical Survey Unit<sup>4</sup>:

$$r' = R_c + 0.275 \cdot (R_c - I_c) + 0.008 \quad (3)$$

$$(r' - i') = 1.052 \cdot (R_c - I_c) + 0.004 \quad (4)$$

These transformations are derived from multi-epoch observations of standards due to Landolt (1992). The measured uncertainty of these transformations ( $\sim 0.1^m$ ) is several times better than the errors estimated for the position of evolutionary tracks ( $0.5^m$  to  $1.0^m$ , Lawson & Feigelson 2001).

Having obtained model tracks in the appropriate system, we shift the tracks according to the assumed distance and reddening for the region. Previous observations have shown optical stars in the region to follow a standard extinction law  $R_V = 3.1$  (Morbidelli et al. 1997), allowing us to adopt the reddening relations due to Schlegel et al. (1998):

$$A_{r'} = 0.843 \times A_V \quad (5)$$

$$A_{i'} = 0.639 \times A_V \quad (6)$$

We note that the reddening vector runs nearly parallel to the low-mass isochrones. Uncertainties in the extinction will therefore only have a limited effect on the age determination.

As a final step, an automated routine determines the two closest isochrones and mass tracks for each object and interpolates the model properties using an inverse square distance law. The routine also corrects for the  $r'$ -band excess due to the  $H\alpha$  line, which falls inside the  $r'$  band. The corrections are based on the table of simulated colours, and range between  $r'+0.02$  ( $EW_{H\alpha} \cong -20 \text{ \AA}$ ) and  $r'+0.5$  ( $EW_{H\alpha} \cong -500 \text{ \AA}$ ).

The resulting ages, masses and radii for all candidates are listed in the Appendix C1 and histograms are plotted in Fig. 10. We note that the colour-magnitude diagram (Fig. 7)

appears to show a large scatter in the ages. Whether or not this indicates a true age spread will be discussed in §7.

### 3.2 Mass accretion rate

Having determined ages and masses, we now constrain the accretion properties of our candidates. We will adopt the method introduced by De Marchi et al. (2010).

In brief, mass accretion is thought to take place along magnetic field lines which act as channels connecting the disc to the star. The infalling gas is essentially on a ballistic trajectory, falling on to the star at close to free-fall velocities, producing a hot impact shock (Calvet & Gullbring 1998; Gullbring et al. 2000). The energy released in these shocks heats the infalling circumstellar gas, resulting in the broad  $H\alpha$  emission lines seen in classical T Tauri stars. This implies that the measured line luminosity  $L_{H\alpha}$  may be used to estimate the accretion luminosity  $L_{acc}$  and subsequently the mass accretion rate  $\dot{M}_{acc}$  (e.g. see Muzerolle et al. 1998; Kurosawa et al. 2006; Herczeg & Hillenbrand 2008).

#### 3.2.1 Line luminosity ( $L_{H\alpha}$ )

To obtain the line luminosity,  $L_{H\alpha}$ , we can safely assume that the  $H\alpha$  line falls entirely inside the breadth of the  $H\alpha$  filter ( $\text{FWHM} = 95 \text{ \AA}$ ).

To convert the  $H\alpha$  magnitude and equivalent width into luminosity, we recall that IPHAS magnitudes are calibrated in the Vega system. A specially tailored Kurucz 9400 K model calibration spectrum of Vega (Bohlin 2007) was used to estimate the average flux of Vega weighted by the  $H\alpha$  throughput curve:

$$\begin{aligned} \bar{F}_V(H\alpha) &= \frac{\Sigma T(\lambda) F_V(\lambda) \Delta\lambda}{\Sigma T(\lambda) \Delta\lambda} \\ &= 1.813 \times 10^{-9} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ \AA}^{-1} \end{aligned} \quad (7)$$

where  $F_V$  is the Vega model spectrum and  $T$  is the throughput curve, obtained by multiplying the transmission profile of the  $H\alpha$  filter with the CCD efficiency curve.

This number is then multiplied by the width of the filter to obtain the appropriate band-integrated reference flux, i.e.  $F_0 = 95 \cdot \bar{F}_V(H\alpha)$ . This is the flux that corresponds to the magnitude of Vega for which we adopt  $m(H\alpha) = 0.03$  (for historical and definitional reasons, the magnitude of Vega in the optical is not exactly zero, see Bessell et al. 1998).

The flux for any star with a given dereddened magnitude  $m(H\alpha)$  is then obtained from:

$$F_{band} = F_0 \cdot 10^{-0.4 \cdot [m(H\alpha) + 0.03]}, \quad (8)$$

where  $m(H\alpha)$  is dereddened using the optical-IR extinction law due to Howarth (1983):

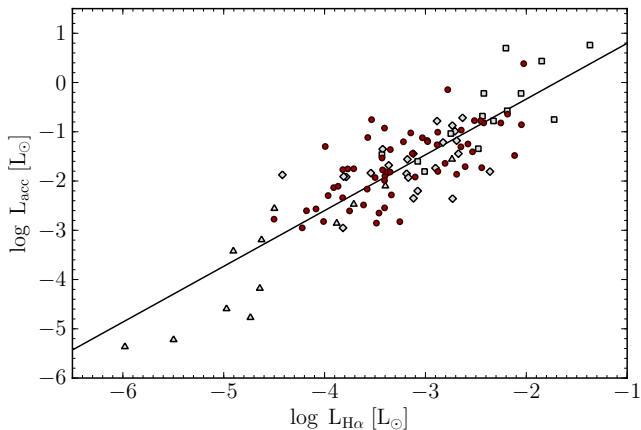
$$A_{H\alpha} = 0.796 \cdot A_V. \quad (9)$$

$F_{band}$  will contain both the continuum and the line flux, but we are only interested in the latter. We may separate the line flux using the equivalent width obtained earlier from the table of simulated colours:

$$\begin{aligned} F_{line} &= F_{band} - F_{continuum} \\ &= F_{band} \cdot \frac{-EW/95}{1 - EW/95}, \end{aligned} \quad (10)$$

<sup>4</sup> <http://www.ast.cam.ac.uk/~wfcsur/technical/photom/colours/>





**Figure 8.** Relationship between the H $\alpha$ -line luminosity and the accretion continuum luminosity. Filled circles represent known members of IC 1396 from Sicilia-Aguilar et al. (2010) where  $L_{H\alpha}$  is derived from spectroscopy and  $L_{acc}$  is based on U-band photometry. We also include literature values from (Hartigan & Kenyon 2003, open diamonds), (Herczeg & Hillenbrand 2008, open triangles) and (Dahm 2008, open squares). The solid line shows the linear fit to all values.

which follows from the definition of EW.

Finally, the line luminosity  $L_{H\alpha}$  is obtained by multiplying the continuum subtracted flux by  $4\pi d^2$ , where  $d$  is the adopted distance to the region. We note that  $L_{H\alpha}$  obtained in this way may include a small contribution from the [NII] forbidden emission lines at 6548 Å and 6584 Å, which fall inside the IPHAS H $\alpha$  filter. De Marchi et al. (2010) corrected for this effect by assuming that 2.4 per cent of the H $\alpha$  intensity is due to the [NII] lines. This correction is based on existing T Tauri line measurements from literature, which show the contribution to be in the range of 0 to  $\sim 5$  per cent. Although the effect is very small compared to other sources of error, it appears to be systematic, so we adopt it.

### 3.2.2 Accretion luminosity ( $L_{acc}$ )

As discussed previously, it is generally accepted that the energy  $L_{acc}$ , which is released in the accretion shock, goes in part towards heating and ionizing the circumstellar gas in the magnetosphere. This will cause the gas to radiate measurable emission lines, i.e. we expect a correlation between  $L_{acc}$  and  $L_{H\alpha}$ .

To study the relationship, we compiled a set of T Tauri objects for which measurements of both  $L_{acc}$  and  $L_{H\alpha}$  are available in the literature (Hartigan & Kenyon 2003; Herczeg & Hillenbrand 2008; Dahm 2008). In these works, high-resolution spectroscopy was used to measure  $L_{acc}$  from the blue continuum excess, while  $L_{H\alpha}$  is derived directly from the H $\alpha$  line. After removing lower and upper limit points, we plot the resulting set of 49 objects in Fig. 8 as open symbols, showing a clear trend.

We then extend the sample with known members of IC 1396 for which  $L_{acc}$  was recently determined by Sicilia-Aguilar et al. (2010) using deep U-band photometry and  $L_{H\alpha}$  being available from their earlier spectroscopic work. These additional 59 objects are shown as filled red circles.

The data suggests a relationship ( $L_{acc}$ )  $\propto$  ( $L_{H\alpha}$ ) $^\alpha$  with

$\alpha$  between 0.5 and 2.0. This is consistent with theoretical accretion models due to Muzerolle et al. (1998). In their work, the authors find a range of slopes, ranging from  $\alpha = 0.7$  (small magnetosphere) to  $\alpha = 1.0$  (larger magnetospheres) at moderate accretion rates ( $\dot{M} < 10^{-8} M_\odot \text{yr}^{-1}$ ;  $L_{acc} \lesssim 10^{-1} L_\odot$ ). At higher rates ( $\dot{M} > 10^{-8} M_\odot \text{yr}^{-1}$ ), the slope steepens to  $\alpha > 1$ , i.e. a smaller increase in  $L_{H\alpha}$  per unit increase of  $L_{acc}$ . This is because, at very high accretion rates, the densities reach a level where the magnetosphere becomes optically thick and the line flux will only emerge from the outer surface. In addition, increasing outflows may produce H $\alpha$  absorption.

In the absence of a sufficiently tight correlation to fit a nonlinear trend, we applied a linear regression and find:

$$\log L_{acc} = (1.13 \pm 0.07) \log L_{H\alpha} + (1.93 \pm 0.23), \quad (11)$$

where  $L_{H\alpha}$  and  $L_{acc}$  are given in solar units.

We note that there is a significant scatter in the relationship (rms = 0.54), which is also present for objects where  $L_{H\alpha}$  and  $L_{acc}$  were measured simultaneously. Indeed correlations of the accretion luminosities with line fluxes have previously been found to show a large amount of scatter (e.g. see Herczeg & Hillenbrand 2008), which is likely to be caused by a combination of effects (e.g. circumstellar absorption, uncertain extinction corrections, line emission not due to accretion). In §7 we will show that this scatter is the primary source of uncertainty for the accretion rates obtained in this work.

### 3.2.3 Mass accretion rate ( $\dot{M}_{acc}$ )

Having obtained an estimate for  $L_{acc}$ , we now proceed to convert this into an estimate for the mass accretion rate  $\dot{M}_{acc}$ . We will do this using the same equations as adopted by previous studies in the literature. In brief, the T Tauri accretion models (e.g. Hartmann et al. 1998) assume that the circumstellar disc is truncated at an inner radius  $R_{in}$  by the magnetosphere, and accretion proceeds along magnetic field lines. The infalling gas is essentially on a ballistic trajectory, falling on to the star at near free-fall velocities from the distance  $R_{in}$ . If we assume that the energy released in the impact ‘‘accretion’’ shock is reprocessed entirely in the accretion energy  $L_{acc}$ , we may write:

$$L_{acc} \simeq \frac{GM_\star \dot{M}}{R_\star} \left(1 - \frac{R_\star}{R_{in}}\right) \quad (12)$$

By adopting a standard value  $R_{in} = 5 \pm 2 R_\star$  (Gullbring et al. 1998), we obtain:

$$\dot{M}_{acc} \simeq 1.25 \frac{L_{acc} R_\star}{GM_\star} \quad (13)$$

Where  $R_\star$  and  $M_\star$  are taken from the evolutionary tracks in the ( $r'/r'-i'$ ) diagram and  $L_{acc}$  is obtained as described above. The resulting values for  $\dot{M}_{acc}$  are listed in Appendix C1 and plotted in Fig. 10.

## 4 ARCHIVAL INFRARED DATA

### 4.1 2MASS photometry

Photometry from the near-infrared 2MASS Point Source Catalog (Skrutskie et al. 2006) was retrieved through the

Virtual Observatory Conesearch protocol. We found a match for 152 out of 158 candidates within  $0.3''$  from the IPHAS position (for 105 objects the match is even better than  $0.1''$ , which is expected because the IPHAS astrometric calibration is based on the 2MASS catalog).

The remaining 6 candidates are part of close visual binaries. The centroid coordinates of these objects differ by more than  $1''$  between both databases, because the close neighbours were resolved by IPHAS but not by the 2MASS pipeline. The near-infrared magnitudes for these objects have not been included in our sample, which is tabulated in Appendix B1.

## 4.2 Spitzer photometry

We obtained wide-area Spitzer mid-infrared IRAC (3.6, 4.5, 5.8 and  $8\ \mu\text{m}$ ) and far-infrared MIPS ( $24\ \mu\text{m}$ ) images for the region from the Spitzer Heritage Archive<sup>5</sup>. The IRAC observations were taken in December 2003 (Program ID 37, PI: Giovanni Fazio), while the MIPS data were obtained in June 2004 (Program ID 58, PI: George Rieke).

The IRAC images cover a box of roughly 35 by 35 arcminutes in the central region (shown as a dashed-dotted rectangle in Fig. 9). 65 of our pre-main sequence candidates (41 per cent) fall inside this box, 62 of which are also covered by the MIPS data (which shows a different footprint).

Basic reduction of the IRAC and MIPS images into artifact-mitigated mosaics was performed by a recent version of the Spitzer Science Center Pipeline (versions S18.7 and S18.12). Photometry was then obtained using the *user-list single frame* option in the MOPEX software (version 18.4.0), which allows aperture photometry to be performed based on a user-supplied list of target coordinates.

For the IRAC mosaics, we used a  $3.6''$  (6 pixel) aperture radius, and a sky annulus from  $14.4$  to  $24''$  (24 to 40 pixels). The MIPS mosaic was reduced using a  $13''$  (5.3 pixel) aperture and a sky annulus from  $20$  to  $32''$  (8.2 to 13.1 pixels). The apertures were chosen to avoid flux contamination from nearby sources on one hand, while keeping the required aperture corrections as small as possible on the other hand.

Following the IRAC and MIPS Data Handbooks, we adopted aperture corrections of 1.112, 1.113, 1.125, 1.218 and 1.170 for the 3.6, 4.5, 5.8, 8.0 and  $24\ \mu\text{m}$  bands, respectively. The adopted zeropoints for conversion between flux densities and magnitudes were 280.9, 179.7, 115.0, 64.1 and  $7.14\ \text{Jy}$ .

Out of the 65 objects that fell inside the area covered by IRAC, all were reliably detected in the four bands at  $\text{SNR} \gg 10$ . 54 of these objects were also detected in the MIPS  $24\ \mu\text{m}$  data at  $\text{SNR} > 10$  (from the remaining objects, 8 were not reliably detected with MIPS and 3 fell just outside the footprint).

To ensure the reliability of the extracted magnitudes, all sources were inspected by eye. Four objects were removed from our sample as part of this step: three objects because they suffer from severe flux contamination by a very close or bright neighbour (IPHAS ID's J213647.63+572954.0; J213657.46+572730.3; J213657.67+572733.1) and one ob-

ject which was partially affected by a gap of missing data in the  $4.5\ \mu\text{m}$  band (J213816.87+573926.4).

Finally, for those objects for which Spitzer magnitudes were previously reduced by Sicilia-Aguilar et al. (2006b), we compared our magnitudes with the literature values and find an acceptable agreement within  $\sim 10$  per cent.

The resulting magnitudes for 61 objects are tabulated in Appendix D1 and plotted as colour-colour diagrams in Fig. 11. We repeated the same reduction steps for the list of WTTS objects from SA06 and included them in Fig. 11 (black crosses) for comparison.

## 5 RESULTS

The spatial distribution of our sample is shown on top of the  $\text{H}\alpha$  mosaic in Fig. 9. The size and shape of the symbols reflect the estimated accretion rates and ages. We note that most candidates are clustered in front of the bright-rimmed molecular clouds A, B and E (nomenclature from Pottasch 1956). In the final section of this paper we will argue that the location and dispersion of these clusters suggest that their formation was triggered by the central O-type star.

Histograms of the derived parameters are shown in Fig. 10. We note that the age distribution (panel b) shows a median age of 2.3 Myr. This is in good agreement with the median age of 2.6 Myr for the confirmed classical T Tauri members from SA06. It is also consistent with the expansion age of the gas derived from CO radial velocity measurements (2-3 Myr, Patel et al. 1995), but possibly younger than the upper main sequence turnoff age ( $\sim 4-7$  Myr Marschall et al. 1990; Contreras et al. 2002).

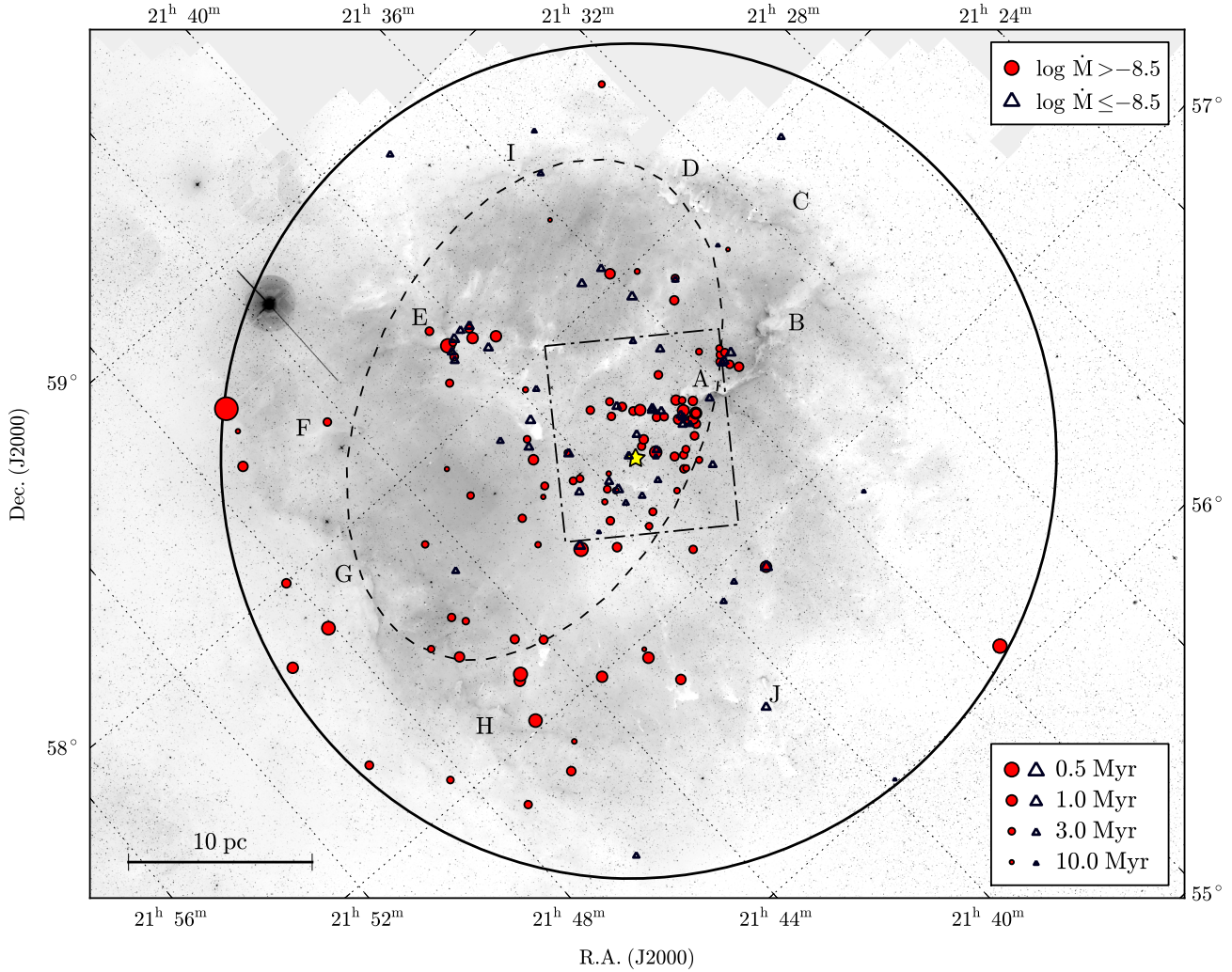
The mass distribution (panel c) shows the expected power-law increase towards lower masses. When compared with a Kroupa IMF, our sample shows a deficit of candidates lower than  $M_\star < 0.3 M_\odot$  and beyond  $M_\star > 1.5 M_\odot$ , which can be explained by the imposed magnitude limits.

The equivalent widths (panel d) range from  $-22$  to  $-525\ \text{\AA}$ , while the mass accretion rates (panel f) range from  $10^{-9.2}$  to  $10^{-7.0} M_\odot \text{yr}^{-1}$ . These ranges are consistent with values for strong accretors in other regions.

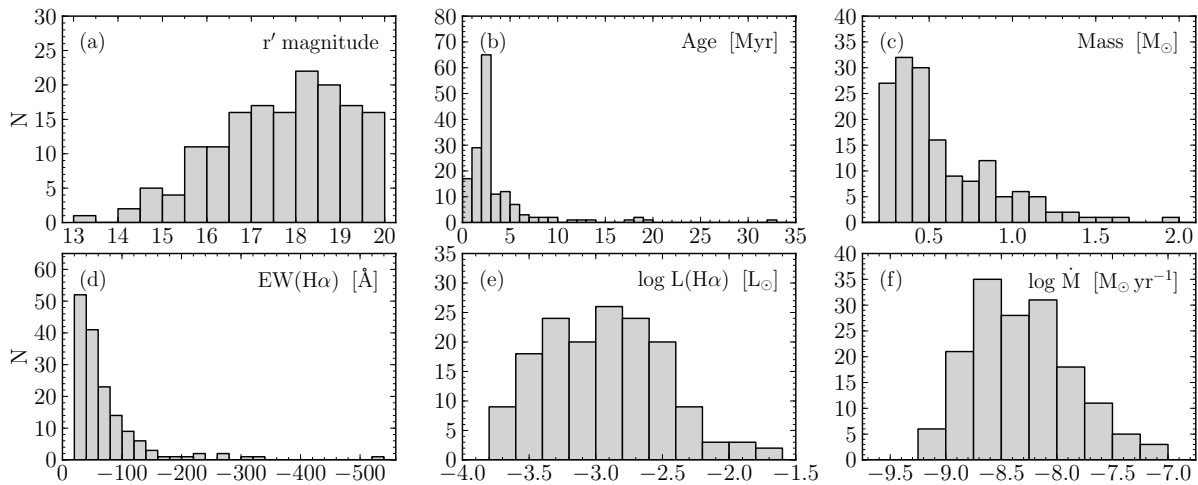
Infrared colour-colour diagrams are shown in Fig. 11. In the near-infrared 2MASS diagram, we note that many candidates are well separated from the general locus of stars due to an H-K excess (91 per cent showing  $\text{H-K} > 0.3$ ; 39 per cent showing  $\text{H-K} > 0.5$ ). The emission is commonly thought to be caused by emission from the inner dust rim of the accretion disc. Some objects are not separated from the main locus, which might give the impression that they do not have a disc. However, the accretion disc models due to Meyer et al. (1997) show that a near-infrared excess may disappear in the case of a large inner-hole size or an edge-on inclination angle, therefore the absence of an H-K excess does not necessarily imply the absence of a disc. (In fact we did not include 2MASS colours as part of our selection criteria for this reason.)

Indeed, the objects for which mid- and far-infrared Spitzer data is available all show colour excesses consistent with having accretion discs (panels b-c-d in Fig. 11). Five objects are seen to show strong infrared excess only at 8 and  $24\ \mu\text{m}$  but not at shorter wavelengths, indicating that dust is absent from the

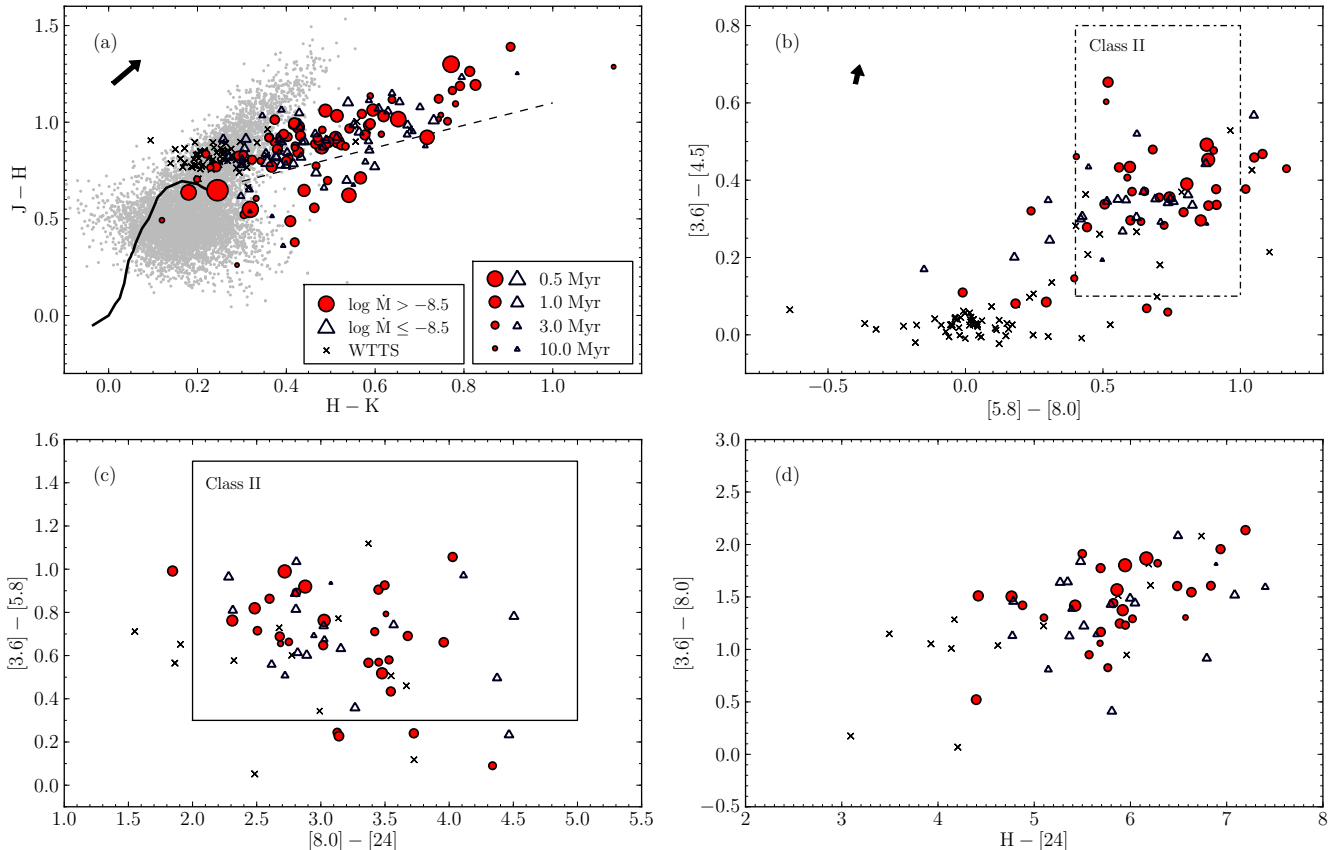
<sup>5</sup> <http://sha.ipac.caltech.edu/applications/Spitzer/SHA>



**Figure 9.** Spatial distribution of pre-main sequence candidates (circles and triangles). The size of the symbol denotes the age, the shape/colour denotes the mass accretion rate. The large solid circle shows the  $3^\circ$  diameter circular region considered in our work, centred on the O6.5V-type star HD 206267 (yellow star symbol). The background shows an inverted greyscale  $H\alpha$  mosaic of the region. The prominent bright-rimmed clouds in the region are indicated with letters A-J as assigned in Pottasch (1956) and Weikard et al. (1996). The expanding ring of molecular material on which the clouds are thought to be located (Patel et al. 1995) is indicated by the dashed ellipse. The central region covered by Spitzer data is indicated by the dashed-dotted rectangle. We note that a large number of candidates are clustered in front of the bright-rimmed clouds A and E.



**Figure 10.** Distribution of stellar and accretion parameters as tabulated in Appendix A-B.



**Figure 11.** Infrared colour-colour diagrams for the pre-main sequence candidates (red filled circles and blue open triangles) and weak-lined cluster members with  $EW_{H\alpha}$  weaker than  $-10 \text{ \AA}$  from Sicilia-Aguilar et al. (crosses). (a) Near-infrared colours for stars with “AAA”-quality data in the 2MASS database. Grey dots show a sample of 8000 field stars in the region. The dashed line shows the typical position of unreddened T Tauri stars determined by Meyer et al. (1997), while the solid line denotes the main sequence for dwarfs due to Bessell & Brett (1988). The arrow shows the reddening vector for  $A_V = 1$ . (b) Spitzer IRAC colours for candidates in the central region (see text). The dashed-dotted box shows the typical colours for T-Tauri stars as predicted by the models of D’Alessio et al. (2005) for  $\log \dot{M} = -8 M_{\odot} \text{ yr}^{-1}$ . The arrow shows the reddening vector for  $A_V = 3$ . (c) Spitzer IRAC colours combined with the MIPS [24] band. The solid box shows the typical colours for Class II objects found by Muzerolle et al. (2004). The objects which fall below the box are candidate Transition Objects. Note that most of the WTTS objects were not detected in the [24] band and are therefore not included in the diagram. (d) Combination of 2MASS, IRAC and MIPS colours. All diagrams show positive but statistically insignificant correlations between the colours, accretion rates and ages.

inner regions of the disk (i.e., candidate “transition objects”; J213709.09+572548.6, J213825.90+574204.8, J213830.30+573255.2, J213955.69+571638.2 and J214035.33+572309.6).

## 6 VALIDATION

### 6.1 Are our candidates genuine T-Tauri stars?

Several arguments allow us to assume that the vast majority of our candidates are genuine T-Tauri members of IC 1396:

- (i) all objects fall within the 0.1-100 Myr isochrones placed at the distance and extinction of the region;
- (ii) the stellar ages of 2-3 Myr agree with literature findings (SA06);
- (iii) the 2MASS near-infrared colours are consistent with the T Tauri locus by Meyer et al. (1997);
- (iv) the Spitzer mid- and far-infrared colours are consis-

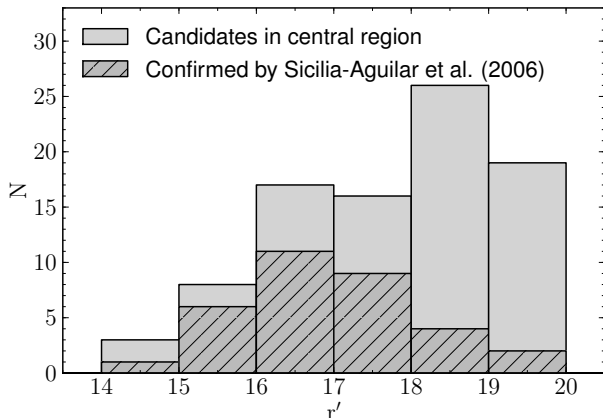
tent with accretion disc models by D’Alessio et al. (2005) and observations by Muzerolle et al. (2004);

- (v) the candidates are strongly clustered;

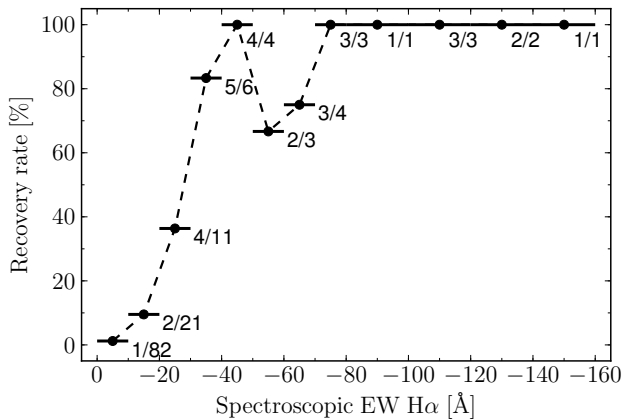
(vi) 49 of the objects are known in SIMBAD, all having types consistent with young stars (T Tau, IR or X-ray).

Moreover, as part of the ongoing IPHAS spectroscopic follow-up campaign, five of our candidates with very strong emission-lines have already been observed by Valdivielso, Bouy & Martin (submitted; objects J213528.42+575823.1, J213545.88+573640.1, J213938.83+575451.4, J214547.74+564845.8 and J214625.99+572828.9). These objects were not previously studied by Sicilia-Aguilar et al. All five spectra show M-type SEDs with strong  $H\alpha$  emission (ranging between  $-105$  and  $-360 \text{ \AA}$ , in agreement with our estimates) as well as emission from the Ca II infrared triplet, which is typical for classical T Tauri type stars.





**Figure 12.** Spectroscopic confirmation of T-Tauri candidates as a function of magnitude. Only objects in the centre of IC 1396, defined as the region within  $0.5^\circ$  from the ionizing star, are included in this plot. This corresponds to the region which was surveyed by SA06. The majority of objects brighter than  $r' < 18$  are spectroscopically confirmed T-Tauri stars. In addition, our work finds many new low-mass candidates at fainter magnitudes due to the greater depth of the IPHAS survey.



**Figure 13.** Fraction of members from Sicilia-Aguilar et al. (2006a) which are also recovered in our work, given as a function of the spectroscopic equivalent width. We included literature objects which are listed as “confirmed” or “probable” members and for which a low-resolution spectroscopic equivalent width is available (141 objects). We find that members with  $H\alpha$  lines weaker than  $-30 \text{ \AA}$  are not recovered because they fall below the selection threshold that we used.

## 6.2 Contamination rate and completeness

While we can assume that the vast majority of candidates are genuine members, it is likely that our method includes at least some contamination from foreground or background emission-line objects. We may constrain this contamination rate as follows.

The surface density of our sample is equal to  $22.3$  candidates/deg $^2$  ( $158$  candidates in  $7.1$  deg $^2$ ). The density is not uniform in the entire area. On one hand, the region within  $1^\circ$  from HD 206267 contains  $135$  candidates and has a density of  $43.0$  objects/deg $^2$ . On the other hand, the region just south of IC 1396 shows a density of only  $\sim 3.2$  objects/deg $^2$ . If we adopt the latter number as an upper limit for the rate of contamination, then our sample would

contain at most  $23$  non-members ( $15$  per cent), which is similar to the number of objects that fall below the T Tauri locus in the 2MASS diagram.

In reality, it is very likely that emission-line stars in the outskirts are genuine members of the surrounding Cepheus OB2 association and that the contamination rate is lower. We conclude that the sample is sufficiently clean to allow it to be used for initial statistical investigations.

A further test to validate our sample is to compare our objects with the members confirmed by SA06. In Fig. 12 we plot the total number of IPHAS candidates in the central region, defined as  $\sim 0.5^\circ$  from HD 206267, and indicate the number of these objects which have been found previously. We find that  $18$  out of  $28$  of our candidates brighter than  $r' < 17$  have already been found and confirmed. At fainter magnitudes, we find many new candidates. This can be explained from the fact that the photometric selection of candidates for spectroscopic follow-up by Sicilia-Aguilar et al. concentrated on stars with  $V=15-19$  ( $R \approx 14-18$ ), while we considered objects up to  $r' = 20$ .

To estimate the completeness of our sample, Fig. 13 shows the fraction of objects from SA06 which were successfully retrieved by our work. We find that for objects with  $EW_{H\alpha}$  stronger than  $-30 \text{ \AA}$ ,  $24$  out of  $27$  objects were successfully recovered ( $89$  per cent). From the remaining  $3$  objects, one is heavily affected by source confusion and was rejected by our criteria because the pipeline classified it as “non-stellar” (SA06 ID 13-1877). The two other objects (14-1017 and 92-393) fall just below the selection threshold, which can be explained by a small decrease in the line intensity due to natural variability.

We conclude that, compared to the extensive spectroscopic survey, our method appears to be equally complete for objects with lines stronger than  $-30 \text{ \AA}$ . This is consistent with the fact that the selection threshold lies near the  $-25 \text{ \AA}$  boundary at the cluster reddening (compare Fig. 4 and 6).

In brief, our comparison with the literature indicates that we obtained a reliable and nearly equally complete picture of strong accretors in the region. More precise detectability limits as a function of mass and age will be given in §8.

## 7 UNCERTAINTIES IN THE DERIVED PARAMETERS

Our stellar parameters suffer from uncertainties caused by errors in the input data and the physical assumptions. In what follows we will discuss two types:

(i) *non-systematic uncertainties*: the fact that individual stars may have a true extinction or distance which deviates from the assumed mean values, thereby introducing errors in the derived parameters relative to the other objects in the sample;

(ii) *systematic uncertainties*: the fact that the assumed mean distance/extinction values or the evolutionary tracks may be shifted in a way that affects the derived parameters for all stars in a systematic direction.

In what follows we will also discuss the effect of unresolved binaries and investigate the age spread.

**Table 1.** Summary of the non-systematic uncertainties. The mean standard deviations ( $\bar{\sigma}$ ) are computed from the individual  $1\sigma$ -errors given for each object in Table C1.

Source	Simulated uncertainty	$\bar{\sigma}(\text{Age})$ [Myr]	$\bar{\sigma}(\text{Mass})$ [ $M_{\odot}$ ]	$\bar{\sigma}(\text{Radius})$ [ $R_{\odot}$ ]	$\bar{\sigma}(\text{EW}_{\text{H}\alpha})$ [ $\text{\AA}$ ]	$\bar{\sigma}(\log L_{\text{H}\alpha})$ [ $L_{\odot}$ ]	$\bar{\sigma}(\log \dot{M})$ [ $M_{\odot} \text{ yr}^{-1}$ ]
Photometry	$\bar{\sigma}(r') = 0.01$	0.18	0.01	0.04	<b>4.16</b>	0.02	0.04
Extinction	$\sigma(A_V) = 0.55$	<b>0.78</b>	<b>0.10</b>	<b>0.17</b>	3.89	<b>0.19</b>	0.20
Distance	$\sigma(d) = 20 \text{ pc}$	0.20	0.01	0.05	0.00	0.02	0.04
$L_{\text{H}\alpha}$ - $L_{\text{acc}}$ relation	$\sigma(\log L_{\text{acc}}) = 0.54$	-	-	-	-	-	<b>0.54</b>
Combined	All of the above	0.89	0.10	0.18	5.97	0.20	0.59

**Table 2.** Summary of the simulated systematic uncertainties. Each column shows how shifting one of the input assumptions affects the mean values of the stellar parameters in our sample.

Source	Simulated shift	$\mu(\text{Age})$ [Myr]	$\mu(\text{Mass})$ [ $M_{\odot}$ ]	$\mu(\text{Radius})$ [ $R_{\odot}$ ]	$\mu(\text{EW}_{\text{H}\alpha})$ [ $\text{\AA}$ ]	$\mu(\log L_{\text{H}\alpha})$ [ $L_{\odot}$ ]	$\mu(\log \dot{M})$ [ $M_{\odot} \text{ yr}^{-1}$ ]
Extinction shift	$\bar{A}_V + 0.5$	-0.00	+0.10	+0.11	-3.71	+0.18	+0.17
"	$\bar{A}_V - 0.5$	+0.20	-0.08	-0.08	+3.33	-0.18	-0.16
Distance shift	$\bar{d} + 100 \text{ pc}$	-0.71	+0.00	+0.17	-	+0.09	+0.15
"	$\bar{d} - 100 \text{ pc}$	+1.10	-0.01	-0.14	-	-0.11	-0.15
Model tracks shift	$r' + 1.0; i' + 0.5$	-1.98	-0.26	+0.19	-	-	+0.27
"	$r' - 1.0; i' - 0.5$	+4.43	+0.29	-0.07	-	-	-0.25

## 7.1 Non-systematic uncertainties

Constraining the uncertainties using analytical error propagation would be complex and error-prone. Instead we opt to investigate the uncertainties using a Monte Carlo approach which works as follows.

For each star in our sample, we generated 10 000 sets of artificial ‘‘clones’’. Each clone is perturbed by randomly adding Gaussian noise to the input photometry and to the physical assumptions (distance, extinction,  $L_{\text{H}\alpha}$ - $L_{\text{acc}}$  regression). For the photometry, the noise is added according to the Poissonian errors listed in Table B1. For the error in the extinction, we assume the literature value  $\sigma(A_V) = 0.55$  (SA05). The adopted error in distance is the approximate diameter of the region:  $\sigma(d) = 20 \text{ pc}$ . Finally, the error in the  $L_{\text{H}\alpha}$ - $L_{\text{acc}}$  regression is the rms scatter we found previously:  $\sigma(L_{\text{acc}}) = 0.54$ .

After generating a set of clones, we then recompute the stellar parameters for each clone; resulting in a set of 10 000 estimates per parameter per star. These sets can be considered to be samplings of probability distributions. We computed the mean and 1-sigma standard deviation<sup>6</sup> from each of these sets, which are the numbers given for each object in Table C1.

We summarized the uncertainties in Table 1 by showing the mean 1-sigma uncertainty for the entire sample per output parameter and per source of error. The summary shows that the determination of the stellar parameters suffers mostly from the uncertainty in the extinction. However,

the final uncertainty in the accretion rate is dominated by the scatter in the  $L_{\text{H}\alpha}$ - $L_{\text{acc}}$  regression. Because individual uncertainties add quadratically to form the combined uncertainty (i.e.,  $\sigma_{A+B} = \sqrt{\sigma_A^2 + \sigma_B^2}$  if A and B are independent), a single largest source of uncertainty tends to dominate the total uncertainty, therefore a better estimate of the extinction would not significantly improve the accuracy of the accretion rates.

We conclude that the accretion rates are constrained within an uncertainty of  $\sim 0.6$  dex relative to the other stars in the sample, mainly due to the scatter in the relationship between  $L_{\text{H}\alpha}$  and  $L_{\text{acc}}$ . The ages are constrained within  $\sim 1$  Myr and masses within  $\sim 0.1 M_{\odot}$ .

## 7.2 Systematic uncertainties

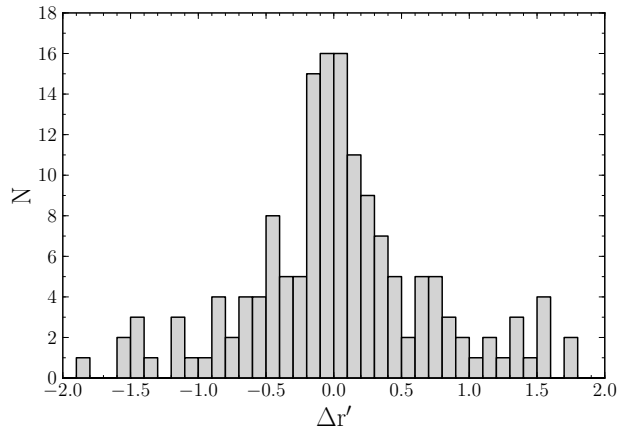
We investigated the effect of systematic uncertainties by introducing shifts in:

- (i) the assumed mean extinction ( $A_V \pm 0.5$ );
- (ii) the assumed mean distance ( $d \pm 100 \text{ pc}$ );
- (iii) the position of the evolutionary tracks.

We recomputed the stellar parameters with each of these shifts introduced, and summarized the results in Table 2. We find that a shift in the assumed extinction shifts the masses by  $\pm 0.1 M_{\odot}$ , while a shift in the distance shifts the ages by  $\pm 1$  Myr.

For simulating errors in the model track positions, we adopted the conclusions by Lawson & Feigelson (2001) who found that the position of mass tracks and isochrones in colour-magnitude diagrams differ by up to  $\sim 1.0$  mag and  $\sim 0.5$  colour between different models. Our simulations show that shifting the Siess et al. (2000) tracks downwards by  $(r', r'-i') + (1.0, 0.5)$  makes the ages on average 2 Myr younger and masses  $0.3 M_{\odot}$  lighter. Shifting the tracks upwards by

<sup>6</sup> In the case of the stellar age, where the probability distribution is heavily skewed, we list the median age along with the 25% and 75% quartiles. The logarithm of the age does have a symmetric uncertainty distribution, but we choose to print the linear age for readability.



**Figure 14.** Spread in  $r'$  magnitude around the 2.5 Myr isochrone shown in Fig. 3. It is unclear whether the scatter can entirely be explained by photometric variability and uncertainties, or if a genuine age spread exists in the region.

( $r', r'-i'$ ) – (1.0, 0.5) yields ages which are 4 Myr older and masses  $0.3 M_{\odot}$  higher. The ages appear to be very uncertain for this reason, and may perhaps be wrong by a factor two in an absolute sense (see also Naylor 2009; Baraffe et al. 2009).

### 7.3 Unresolved binaries

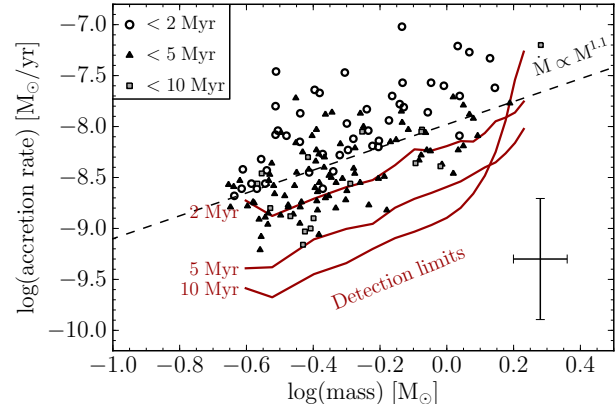
We cannot exclude that a fraction of our candidates are very close unresolved binaries (e.g. having a separation of less than  $\sim 1''$ , i.e.  $\sim 1000$  AU). We investigated this effect by simulating what happens when the luminosity of a star is over-estimated by a factor two while the colour remains unchanged (the case of an unresolved equal-mass binary system). We find that such case leads to an age estimate which is on average 4 Myr too young and mass accretion rates which are 0.5 dex too low.

We did not include this effect in the simulations because the fraction of unresolved binaries is highly uncertain. Observations by Bouwman et al. (2006) suggest that close binaries lose their disk significantly faster than single stars, which means that the number of binaries in our sample could be very low. If this is not the case, unresolved binaries may be the single most important source of error in our work.

### 7.4 Is the spread in ages real?

The estimated ages of our candidates range between 0.2 and 32 Myr, with a median of 2.3 and standard deviation 4.0. The spread is significantly larger than the standard deviation of  $\sim 0.9$  which is predicted to occur due to the combined uncertainties in the photometry, extinction and distance (Table 1). Does this mean that a true age spread exists in the region?

We must take into account that T Tauri stars have been observed to change in brightness by as much as one magnitude on timescales of only days (e.g. Eiroa et al. 2002), which is thought to be explained by stellar spots, variable accretion, or variable obscuration by circumstellar dust. In the colour-magnitude diagram presented earlier (Fig. 3), we found that the objects are widely scattered around the 2.5 Myr isochrone, with 79 per cent deviating



**Figure 15.** Distribution of mass accretion rates as a function of stellar mass for our T-Tauri candidates. We show candidates younger than 2 Myr (open circles), 5 Myr (filled triangles) and 10 Myr (filled squares). The solid red lines show the detectability limits as a function of age and mass. The dashed line shows a linear regression through the data, which is heavily affected by the detectability limits.

by  $\Delta r' \geq 0.1$  from the isochrone and 37 per cent showing  $\Delta r' \geq 0.5$  (Fig. 14). This is slightly larger than the variability of stars in this region recently reported by Sicilia-Aguilar et al. (2010) based on multi-epoch V band magnitudes, where the authors found 65 per cent showing a variability  $\geq 0.1$  mag and 16 per cent  $\geq 0.5$  mag. This indicated that the observed variability may explain a significant part of the age spread, although perhaps not entirely.

If the effect of unresolved binaries play a significant role, it is likely that the combined effect of uncertainties and variability does explain the entire age spread. However, it will always be difficult to exclude that perhaps a small true age spread – heavily masked by the noise – is hiding in the large apparent spread.

## 8 DISCUSSION

In this section we provide a discussion on the implications of our results. First we investigate how our mass accretion rates depend on age and mass, then we discuss the implications for the history of IC 1396 and the role of the massive star HD 206267.

### 8.1 Accretion rates as a function of mass and age

#### 8.1.1 The $\dot{M}_{\text{acc}}-M_{\star}$ relationship

Fig. 15 shows the distribution of mass accretion rates as a function of stellar mass. We fitted the  $\dot{M}_{\text{acc}} \propto M_{\star}^{\alpha}$  power-law relationship and find a slope  $\alpha = 1.1 \pm 0.2$  (all objects) or  $\alpha = 1.3 \pm 0.2$  (objects younger than 2 Myr), although with a large spread in  $\dot{M}_{\text{acc}}$  for any  $M_{\star}$ .

The slope is less steep than recently claimed in literature, where different authors found the accretion rate to correlate roughly with the square of the stellar mass;  $\alpha = 1.8 - 2.3$  (Muzerolle et al. 2003; Mohanty et al. 2005; Natta et al. 2006; Herczeg & Hillenbrand 2008) and recently as steep as  $\alpha = 3.1$  (Fang et al. 2009).

The apparent inconsistency with the literature may be explained by two reasons. First, our dataset does not include very low mass stars or brown dwarfs, which form a significant part of the mass range considered in several of the previous studies. We do not rule out that our slope would have been steeper if we had been able to detect such objects at the distance of IC 1396. Likewise, we note that our data covers only  $\sim 2$  magnitudes in mass accretion rates.

Second, following Clarke & Pringle (2006) and Mayne & Harries (2010), the determination of the slope is affected by method-dependent systematic effects, namely:

- (i) detectability limits in the  $(\dot{M}_{\text{acc}}, M_{\star})$  plane;
- (ii) intrinsic correlations between  $M_{\star}/R_{\star}/L_{\star}$ .

An advantage of our photometric method is that it is straightforward to derive its detectability limits, because all objects were selected in a homogeneous way from a precisely defined threshold in the  $(r^{\prime}-i^{\prime}/r^{\prime}-H\alpha)$  plane. The detectability limits derived from this threshold, computed for objects of different ages, are shown as red solid lines in Fig. 15. We find that the limits correlate with the mass in a similar way as the accretion rates ( $\alpha \simeq 1.2$ ); i.e. it is easier to detect smaller accretion luminosities around lower mass (less luminous) stars.

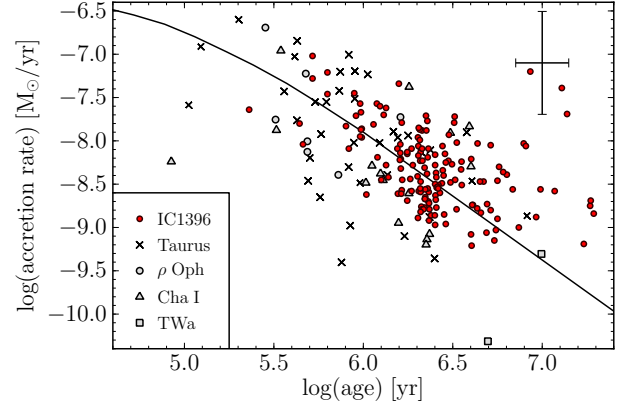
It is worth noting that the compilation by Natta et al. (2006), which shows a steeper dependency ( $\alpha = 1.8 \pm 0.2$ ), is also affected by a steeper slope in the detection limit ( $\alpha \simeq 2$ , see Clarke & Pringle 2006). The selection effects of the other studies, however, are less clear.

The second effect suggested by Clarke & Pringle (2006) are the tight natural correlations between  $M_{\star}/R_{\star}/L_{\star}$  for stars of a similar age. Because we derive all stellar parameters using a single method (namely, model tracks in the IPHAS diagrams), the uncertainties of these parameters are not independent. Using the Monte Carlo simulations discussed previously, we were able to determine the covariance between the mass and the accretion rate for each star (using the set of 10 000 clones disturbed with errors). We find strong positive covariances, i.e. the uncertainty in the accretion rate tends to be correlated and have the same sign as the uncertainty in the stellar mass. The mean covariance-to-variance ratio is:

$$\left\langle \frac{\text{covariance}(\log \dot{M}_{\text{acc}}, \log M_{\star})}{\text{variance}(\log M_{\star})} \right\rangle = 1.7. \quad (14)$$

This ratio can be interpreted as follows: in a sample of (equal-mass) stars where the apparent mass range would be entirely due to uncertainty, we would expect to retrieve a default slope  $\alpha = 1.7$  driven purely by the effect of correlated uncertainties. Fortunately, the mass range considered in this paper is several times larger than the estimated mass uncertainty, therefore the impact of the effect is limited. Our simulations suggest that the effect accounts for an overestimation of the slope by  $\sim 10$  per cent.

We conclude that we do not find evidence to suggest that accretion rates correlate with the square of the stellar mass in the range  $0.2-2.0 M_{\star}$ . We stress, however, that we do not rule out such relationship from this dataset. We do emphasize the importance of including detection limits and covariance estimates in the discussion.



**Figure 16.** Distribution of mass accretion rates as a function of stellar age for our T-Tauri candidates (red filled circles), compared with literature values collected by Muzerolle et al. (2000). The solid line indicates the evolution for models of viscous disc evolution due to Hartmann et al. (1998). The average 1-sigma error is shown in the top right.

### 8.1.2 The $\dot{M}_{\text{acc}}$ -Age relationship

Figure 16 shows the distribution of accretion rates as a function of stellar age, together with spectroscopic results for other regions taken from literature. We find  $\dot{M}_{\text{acc}} \propto \text{Age}^{-0.9 \pm 0.1}$  for our sample, which is in excellent agreement with the trend for literature data:  $\dot{M}_{\text{acc}} \propto \text{Age}^{-1.0 \pm 0.2}$  (Muzerolle et al. 2000), although slightly less steep than expected theoretically (Hartmann et al. 1998).

However, the  $\dot{M}_{\text{acc}}$ -Age relationship is subject to the same statistical effects discussed in the previous section. First, in Fig. 15 we already showed that lower mass accretion rates are more easily detected around older (less luminous) stars. Second, intrinsic correlations in the parameter space yield a negative covariance-to-variance ratio:

$$\left\langle \frac{\text{covariance}(\log \dot{M}_{\text{acc}}, \log \text{Age})}{\text{variance}(\log \text{Age})} \right\rangle = -0.6. \quad (15)$$

i.e. in a sample of coeval stars where the age spread is dominated by errors, we expect to find a default relationship  $\dot{M}_{\text{acc}} \propto \text{Age}^{-0.6}$ .

The effect can easily be understood from the fact that the age and the radius are both estimated from the same colour-magnitude diagram. For example, when the magnitude of an object is over-estimated (e.g. due to an error in the photometry or distance), the object will be located higher up in the colour-magnitude diagram near a younger isochrone, which will lead to an under-estimated age and at the same time an over-estimated radius (because  $\log R_{\star}$  is roughly proportional to  $\log \text{Age}$  in the model tracks). In turn, the larger radius will automatically lead to an over-estimated accretion rate (cf. Equation 13).

We have previously shown that, unlike masses, the model ages are highly uncertain and the age spread is likely to be explained, at least for a significant part, by uncertainties and variability. Therefore the effect of correlated uncertainties is likely to play a very significant role. The same effect is likely to occur in spectroscopy-based studies as well.

We conclude that the  $\dot{M}_{\text{acc}}$ -Age relationship for a single region must be interpreted with great care, especially where



the age spread is uncertain. Large samples of accretion rates across regions, spanning a large range of truly different ages, are essential before conclusions may be drawn on the relationship.

## 8.2 Role of the ionizing star HD 206267 (O6.5V)

### 8.2.1 The influence of massive stars

It is often claimed that radiation and stellar winds from O-type stars may have a profound effect on nearby star formation. On one hand, a massive star may disrupt star formation by dispersing the local molecular cloud and removing circumstellar material from newly formed stars (*UV photoevaporation*, e.g. Johnstone et al. 1998). On the other hand, expanding ionization fronts may act to compress the gas around the periphery of an HII region which then collapses to form stars (*collect and collapse*, Elmegreen & Lada 1977), or trigger star formation in a pre-existing cloud (e.g. *radiatively driven implosion*, Kessel-Deynet & Burkert 2003).

While the effect of photoevaporation has been observed using direct imaging (so-called *proplyd objects*, O'dell et al. 1993; O'dell & Wen 1994), a proof of triggered star formation is harder to obtain. Merely observing the presence of young stars in molecular clouds near hot stars is *no solid evidence*, because it is not clear whether the birth of the stars was triggered by the hot star, or whether stars were forming in these clouds regardless. Moreover, photoevaporation may act to disperse the circumstellar environment of protostars as a function of their distance, and therefore introduce an apparent evolutionary gradient which may incorrectly be interpreted as an age gradient (e.g. discussed by Beltrán et al. 2009).

The relatively simple nature of IC 1396, where the UV radiation field is dominated by the massive star system HD 206267 (O6.5V), makes it a good region to study the influence of a hot star. We note that the region also contains an O9V and several B-type stars (Simonson 1968; Garrison & Kormendy 1976), but these objects are thought to produce one (O9.5V) to five (B3V) orders of magnitude less ionising radiation compared to an O6.5V-type (Panagia 1973; Martins et al. 2005).

In what follows, we discuss the hypothesis that star formation in IC 1396 is triggered, and test it by investigating the properties of our candidates as a function of their distance from HD 206267. We discover an age gradient which, we will argue, provides evidence for triggered star formation.

### 8.2.2 Hypothesis: is star formation in IC 1396 triggered?

IC 1396 is known to contain multiple dense molecular clouds (e.g. BRC 32-42, Sugitani et al. 1991) containing a total mass between 9000 and 12000  $M_{\odot}$  (Weikard et al. 1996; Froebrich et al. 2005). The discovery of red IRAS point sources associated with these clouds first suggested that stars are actively forming here (Sugitani et al. 1991; Schwartz et al. 1991). More recently, Spitzer and Chandra observations uncovered a large number of embedded protostars (e.g. Reach et al. 2004; Sicilia-Aguilar et al. 2006b; Getman et al. 2007; Ikeda et al. 2008; Mercer et al. 2009; Choudhury et al. 2010), while optical studies have found Herbig Haro outflows (Reipurth et al. 1997; Froebrich et

al. 2005). These results indicate the presence of very young YSO-type objects ( $<1$  Myr, Evans et al. 2009) in the region, although none of them are part of our sample because they are not detected in the optical.

Using CO radial velocities, Patel et al. (1995) found that the clouds in the region are part of an expanding ring of molecular material centred roughly around HD 206267 (dashed ellipse in Fig. 9). In their picture, the massive star was part of a first population which formed 3 to 4 Myr ago when the parent molecular cloud was compressed and flattened by turbulence (e.g., a supernova in the surrounding Cepheus OB2 association). Subsequently, wind or radiation from HD 206267 caused leftover molecular material to be swept up in the expanding ring that is observed today, and is thought to be the site of a second generation of star formation.

This picture is consistent with our results, because many of our youngest objects are clustered just in front of the bright-rimmed clouds A, B and E (see Fig. 9). The position of these objects between the clouds – known to contain protostars – and the massive star, suggests that they may have formed at an earlier stage when the shock front was located closer to the hot star. The typical radial velocity dispersion in the region ( $\sim 1-2$  pc Myr $^{-1}$ , SA06<sup>7</sup>) suggests that our candidates may indeed still be located close to the place where they once formed.

### 8.2.3 Test: does our sample show an age gradient?

Whether the clusters of stars observed in our sample have been triggered is a testable hypothesis. If star formation is triggered, we expect that:

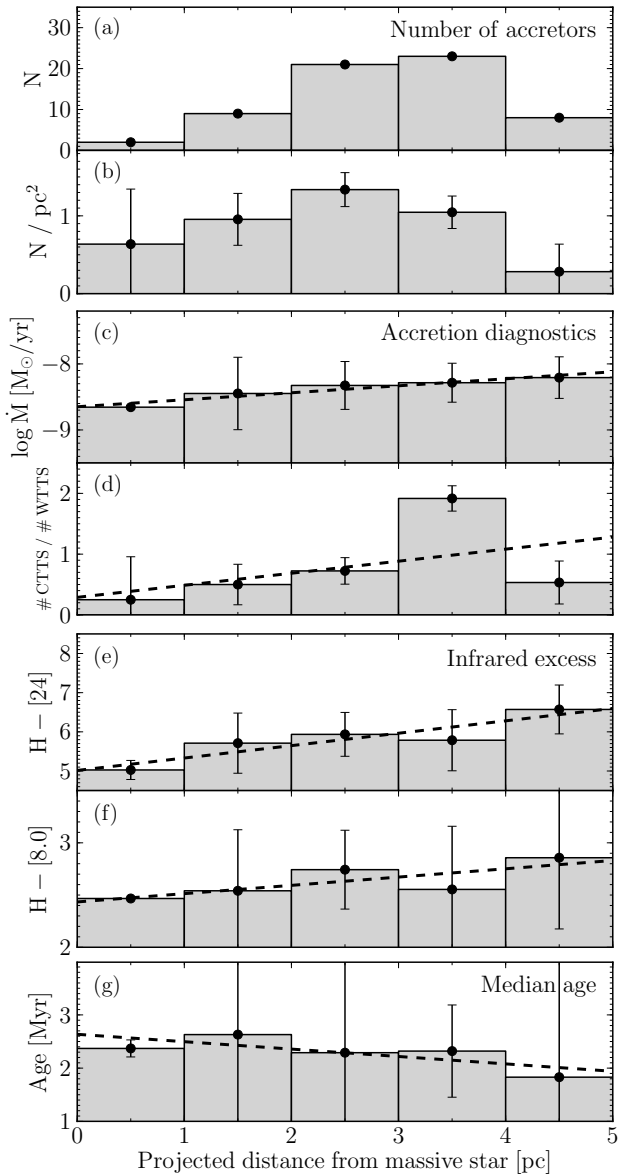
- (i) the stars near the shock fronts must be younger than HD 206267;
- (ii) there should be a small age spread and gradient; stars located further away from HD 206267 and closer to the shock fronts should be younger.

Condition (i) is consistent with the fact that the model ages for the majority of our candidates ( $\sim 2-3$  Myr) do not exceed the dynamical age of the expanding ring of molecular material (2-3 Myr, Patel et al. 1995). While there is no reliable estimate of the age of HD 206267 itself, we note that the main-sequence lifetime of an O6.5V-type star (4 Myr, Weidner & Vink 2010) does not contradict the picture.

Proposed evidence for condition (ii), the age gradient, comes from the spatial distribution of the youngest members. First, SA05 noted an increasing number of objects younger than 1 Myr towards globule A. Subsequently, Spitzer observations of clouds A, E and J (Sicilia-Aguilar et al. 2006b; Getman et al. 2007; Ikeda et al. 2008; Choudhury et al. 2010) showed that the embedded Class 0/I protostars are systematically located further away from HD 206267 than the more evolved optical Class II/III objects.

We investigate if such gradient is also apparent in our sample. In Fig. 17 we show the number of candidates as a function of the projected distance from HD 206267 (panel a)

<sup>7</sup> Sicilia-Aguilar et al. (2006a) found a standard deviation of  $3.6$  km s $^{-1}$  ( $= 3.7$  pc Myr $^{-1}$ ) in the radial velocity, but reported that individual errors contributed  $1-2$  km s $^{-1}$ .



**Figure 17.** Candidate properties in function of the projected distance from HD 206267, limited out to 5 pc where both Spitzer and literature data are available. With increasing distance from the massive star, we find higher mean accretion rates (panel c), a higher fraction of accretors to the number of non-accretors from SA06 (panel d), stronger mean infrared disc excess (panels e-f), and a lower median age (panel g). The dashed lines show linear fits. All panels are consistent with a spatio-temporal gradient where objects further away from the massive star have a younger appearance. The bright rim of molecular cloud A, the Elephant Trunk Nebula, is located at the distance of 3.8 pc.

and the corresponding surface density (panel b). We find a significant peak near  $\sim 3$  pc, which corresponds to cluster in front of cloud A (commonly known as the Elephant’s Trunk Nebula). We also show the mean accretion rate (panel c) and the number of weak-line (non-accreting) T-Tauri stars in the region obtained from SA06 (panel d). We find an apparent increase in the accretion activity further away from the hot star. This trend is consistent with an increase in the mean

mid- and far-infrared excess (panels e-f) and a decrease in the median stellar age (panel g).

Dashed lines show a linear regression of the mean properties as a function of radial distance. Although the statistical significance of the individual correlations is only moderate at best (the p-values being 0.01, 0.4, 0.03, 0.12, 0.14 for panels c, d, e, f, g), it is significant that all panels are mutually consistent with the picture of a spatio-temporal gradient towards cloud A.

#### 8.2.4 Can the age gradient be explained by photoevaporation instead?

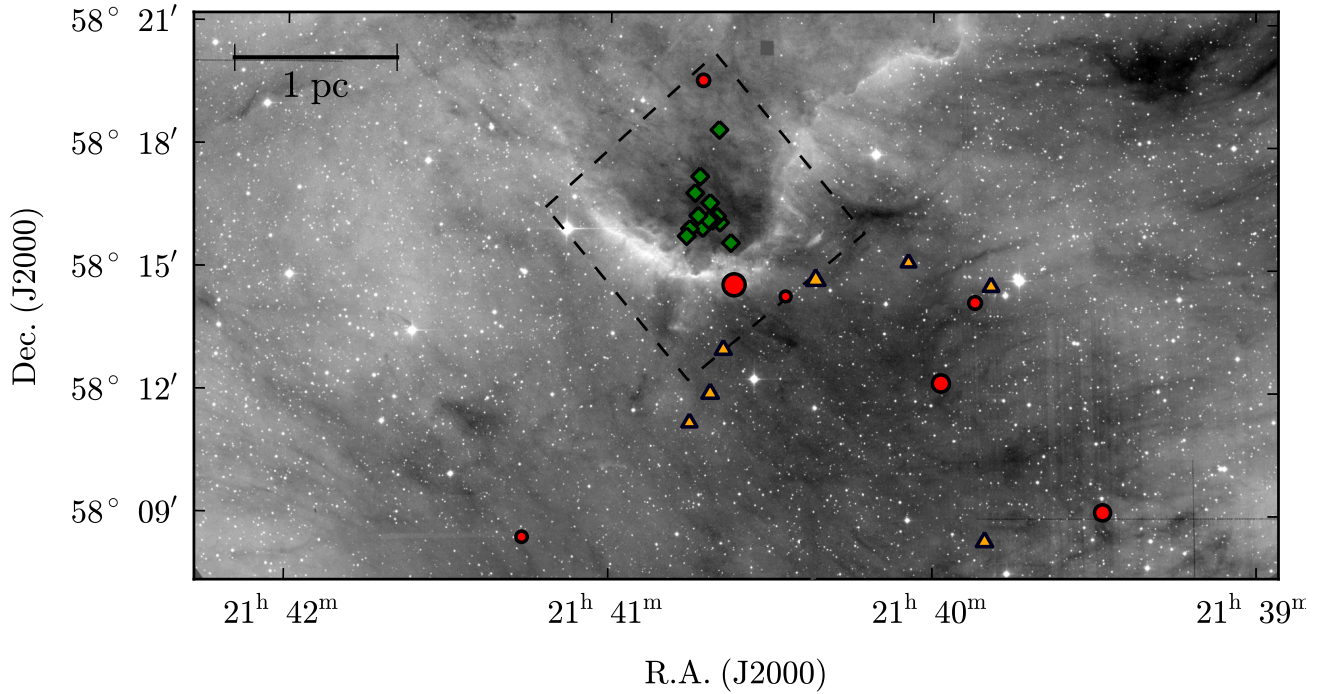
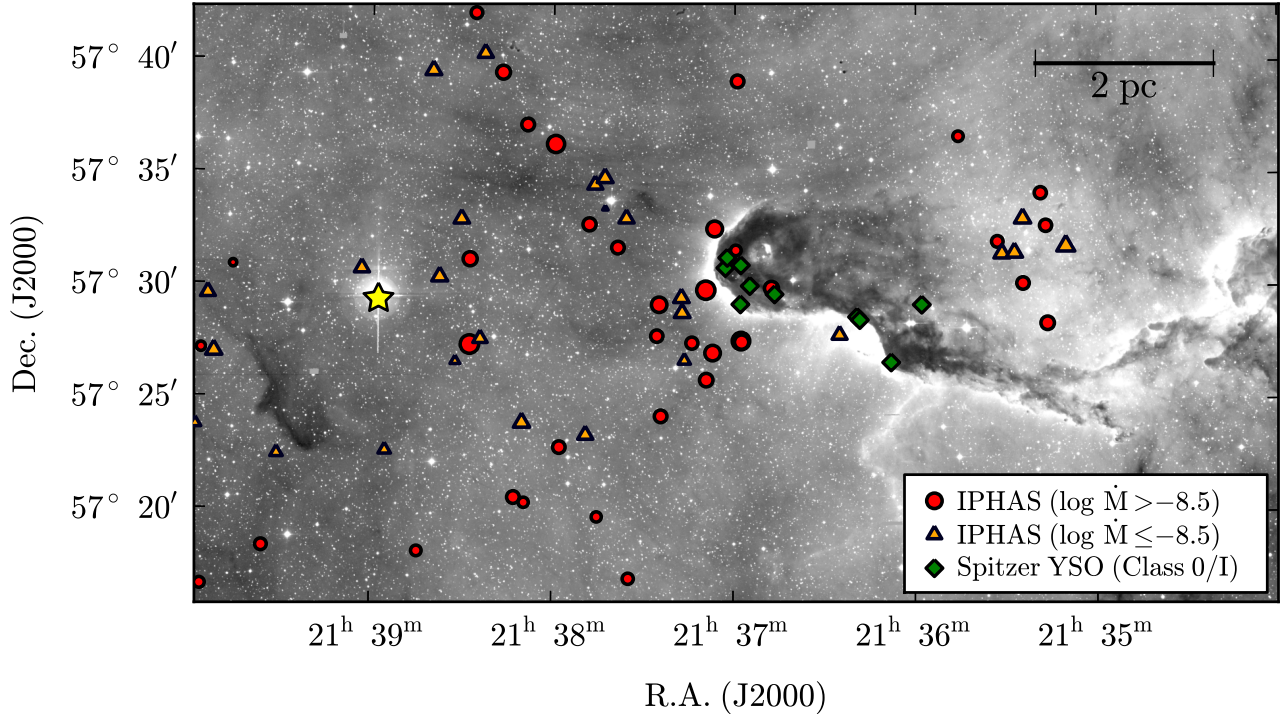
Photoevaporation denotes the process where radiation heats disc particles until they reach the escape velocity and “evaporate” into space. The effect is often discussed in the context of irradiation from the central star (“inside-out” photoevaporation, e.g. Ercolano et al. 2009), but here we consider it in the context of the external radiation from HD 206267 (“outside-in”). Beltrán et al. (2009) argued that the infrared excess gradients observed in IC 1396 would also be expected if such photoevaporation is progressively removing circumstellar material. The Spitzer-based discovery by Balog et al. (2006) of a tail on a proplyd-like object in the region, confirms that the effect certainly plays a role here.

However, the proplyd in question is located at only 0.5 pc from HD 206267, while clouds A, B and E are located significantly further away at  $\sim 4$ , 9, and 12 pc. The available Spitzer data does not show proplyd-like tails around the stars in front of these clouds, and SA05 did not find a clear absence of accreting stars near the O-star. Both models and observations (Richling & Yorke 1998; Balog et al. 2007) suggest that photoevaporation is perhaps only effective at removing disks within  $\sim 1$  pc from the radiation source.

Moreover, while a decrease of the infrared excess may be explained by photoevaporation, it is not obvious that the removal of circumstellar material from the outer disk would necessarily result in the lower mass accretion rates that are observed in Fig. 17. More importantly, it is hard to see how the apparent trend of decreasing model ages may be explained by photoevaporation (although the trend has a low statistical significance).

The strongest argument in favour of triggered star formation comes from the dispersal of objects in front of clouds A/B/E, shown close-up in Fig. 18. The clusters of our candidates in front of the ionized rims are significantly more dispersed than the embedded Class 0/I protostars inside the clouds, covering an area which is roughly 1 to 2 pc larger in diameter. If we were to assume that all objects in the region disperse from their birth at the rate of the cluster’s current spread in radial velocity ( $\sim 1$ -2 pc Myr $^{-1}$ , see earlier), there would be a dynamical argument to suggest that our Class II candidates are more dispersed because they formed roughly 0.5 Myr before the Class 0/I protostars, which is consistent with the apparent age gradient found in Fig. 17.

In summary, we found evidence for a spatio-temporal gradient which is unlikely to be explained by photoevaporation alone. This provides a strong indication that star formation in IC 1396 has, at least in part, been triggered sequentially by HD 206267 during the last  $\sim 1$  Myr.



**Figure 18.** Spatial distribution of our candidates (red circles and blue triangles) near the bright-rimmed clouds A/B (top panel) and cloud E (bottom panel). The size of the symbol denotes the age. We also show the position of Spitzer Class 0/I protostars (green diamonds) as found by Reach et al. (2009, top panel) and Choudhury et al. (2010, bottom panel). The dashed line in the bottom panel shows the footprint of the Spitzer data. HD 206267 is indicated by a yellow star symbol in the top panel, while it is located 12 pc south of the image in the bottom panel. The dispersed nature of our candidates in front of the ionized rims, compared to the compact clustering of the Spitzer protostars inside the clouds, suggests that our candidates were triggered at an earlier stage when the ionization shock was located closer to HD 206267.

## 9 CONCLUSIONS

We presented a photometric study of IC 1396, located at a distance of 870 pc in the Cepheus OB2 association (Contreras et al. 2002). We used IPHAS narrow-band H $\alpha$  imaging and broadband r'/i' photometry to identify 158 low-mass T Tauri candidates from a database of 1 million objects, using a homogeneous selection method which offers a precise handle on the selection effects.

Compared to previous searches by Sicilia-Aguilar et al. (2005, 2006a), we have tripled the number of classical T Tauri candidates in the region, in particular at low masses, with 56 per cent of our candidates below 0.5 M $_{\odot}$ . Our main findings are:

(i) We find a near-linear dependency of the accretion rate on stellar mass of  $\dot{M}_{\text{acc}} \propto M_{\star}^{1.1 \pm 0.2}$ , but the determination of the slope is affected by detectability limits and the limited mass range (0.2 - 2.0 M $_{\odot}$ ).

(ii) A significant fraction of the candidates are found to be part of dispersed clusters which are located in-between the bright-rimmed molecular globules A/B/E and the massive star system HD 206267 (O6.5V). We discovered a spatio-temporal gradient of increasing accretion rates, increasing disc excesses, and younger ages away from HD 206267 towards globule A. This provides a strong indication that the formation of these clusters has been sequentially triggered by the massive star (in support of the picture proposed by Patel et al. 1995). This finding is consistent with recent Spitzer-based discoveries of apparent age gradients inside the globules (Getman et al. 2007; Ikeda et al. 2008; Choudhury et al. 2010).

In future work, we aim to extend our sample by applying the method introduced in this paper to different regions in the Galactic Plane, allowing comparative studies between different star-forming environments to be made based on a uniform selection method and large sample sizes.

A relaxed version of the selection criteria used in this work may also serve as a basis for high-resolution spectroscopic follow-up work.

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This paper has been typeset from a T $_{\text{E}}\text{X}/\text{L}^{\text{A}}\text{T}_{\text{E}}\text{X}$  file prepared by the author.

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## APPENDIX A: SIMULATED IPHAS COLOURS FOR EMISSION-LINE STARS

**Table A1.** Synthetic tracks in the  $(r - H\alpha, r' - i')$  plane for objects with increasing levels of  $H\alpha$  emission equivalent width (EW). The underlying data are provided by the library of SEDs due to Pickles (1998), assuming a standard extinction law  $R_V = 3.1$  (see §2.3).

H $\alpha$ EW( $\text{\AA}$ )	$E(B - V) = 0.0$		$E(B - V) = 1.0$		Reddening $E(B - V) = 2.0$		$E(B - V) = 3.0$		$E(B - V) = 4.0$	
	$(r' - i')$	$(r' - H\alpha)$	$(r' - i')$	$(r' - H\alpha)$	$(r' - i')$	$(r' - H\alpha)$	$(r' - i')$	$(r' - H\alpha)$	$(r' - i')$	$(r' - H\alpha)$
O5V SED										
0	-0.179	0.076	0.509	0.302	1.182	0.486	1.841	0.629	2.486	0.733
-10	-0.185	0.174	0.502	0.399	1.174	0.582	1.831	0.724	2.476	0.827
-20	-0.190	0.262	0.495	0.486	1.166	0.668	1.822	0.809	2.466	0.911
-40	-0.201	0.419	0.482	0.640	1.150	0.819	1.804	0.958	2.446	1.059
-60	-0.212	0.552	0.468	0.771	1.134	0.949	1.786	1.085	2.426	1.184
-100	-0.234	0.773	0.442	0.987	1.103	1.160	1.751	1.293	2.388	1.388
-200	-0.286	1.161	0.379	1.364	1.030	1.527	1.668	1.650	2.297	1.739
-400	-0.383	1.618	0.263	1.803	0.896	1.948	1.519	2.056	2.137	2.133
-1000	-0.631	2.226	-0.025	2.370	0.572	2.480	1.166	2.559	1.762	2.614
-1000000	-6.963	3.291	-6.501	3.291	-6.013	3.291	-5.498	3.291	-4.958	3.292
A0V SED										
0	0.030	0.014	0.719	0.232	1.394	0.408	2.053	0.543	2.699	0.639
-10	0.025	0.102	0.714	0.319	1.387	0.494	2.046	0.628	2.691	0.723
-20	0.020	0.182	0.708	0.398	1.380	0.572	2.038	0.705	2.683	0.800
-40	0.011	0.325	0.697	0.539	1.367	0.711	2.023	0.843	2.666	0.936
-60	0.002	0.449	0.685	0.661	1.354	0.832	2.008	0.962	2.650	1.054
-100	-0.016	0.657	0.663	0.865	1.328	1.032	1.979	1.159	2.619	1.249
-200	-0.060	1.031	0.610	1.230	1.266	1.388	1.910	1.508	2.543	1.592
-400	-0.143	1.484	0.511	1.668	1.152	1.811	1.783	1.918	2.407	1.993
-1000	-0.360	2.108	0.258	2.256	0.868	2.368	1.474	2.450	2.079	2.506
-1000000	-6.562	3.290	-6.091	3.291	-5.593	3.291	-5.069	3.291	-4.520	3.291
F0V SED										
0	0.218	0.134	0.903	0.338	1.573	0.501	2.228	0.623	2.870	0.708
-10	0.213	0.225	0.896	0.428	1.565	0.590	2.219	0.711	2.861	0.796
-20	0.208	0.308	0.890	0.510	1.557	0.671	2.210	0.792	2.851	0.875
-40	0.197	0.455	0.877	0.655	1.542	0.814	2.194	0.933	2.833	1.015
-60	0.186	0.582	0.864	0.780	1.527	0.937	2.177	1.055	2.815	1.136
-100	0.165	0.793	0.839	0.988	1.498	1.141	2.144	1.255	2.780	1.334
-200	0.114	1.169	0.778	1.354	1.428	1.498	2.067	1.605	2.697	1.678
-400	0.018	1.617	0.665	1.785	1.301	1.915	1.927	2.009	2.548	2.073
-1000	-0.226	2.220	0.385	2.352	0.989	2.451	1.591	2.521	2.195	2.567
-1000000	-6.545	3.291	-6.065	3.291	-5.560	3.291	-5.028	3.291	-4.471	3.292
K0V SED										
0	0.461	0.306	1.133	0.491	1.792	0.635	2.437	0.740	3.070	0.809
-10	0.454	0.403	1.125	0.587	1.782	0.730	2.426	0.834	3.059	0.903
-20	0.447	0.491	1.117	0.674	1.773	0.816	2.416	0.919	3.048	0.988
-40	0.433	0.646	1.101	0.826	1.754	0.966	2.395	1.068	3.026	1.135
-60	0.419	0.779	1.085	0.956	1.736	1.094	2.375	1.194	3.005	1.260
-100	0.393	0.996	1.053	1.169	1.701	1.302	2.337	1.399	2.964	1.463
-200	0.329	1.374	0.979	1.537	1.617	1.661	2.245	1.750	2.867	1.808
-400	0.212	1.812	0.843	1.957	1.466	2.066	2.082	2.144	2.696	2.194
-1000	-0.080	2.378	0.516	2.488	1.109	2.567	1.703	2.622	2.302	2.657
-1000000	-6.565	3.291	-6.078	3.291	-5.564	3.291	-5.025	3.291	-4.462	3.292
K5V SED										
0	0.670	0.423	1.340	0.596	1.997	0.730	2.641	0.826	3.274	0.889
-10	0.662	0.522	1.331	0.695	1.986	0.827	2.629	0.923	3.262	0.985
-20	0.654	0.612	1.321	0.783	1.975	0.915	2.617	1.010	3.250	1.071
-40	0.638	0.770	1.303	0.938	1.955	1.067	2.595	1.160	3.226	1.220
-60	0.622	0.903	1.285	1.069	1.934	1.196	2.573	1.287	3.202	1.346
-100	0.592	1.121	1.249	1.282	1.894	1.405	2.529	1.493	3.157	1.549
-200	0.519	1.496	1.166	1.646	1.801	1.760	2.428	1.840	3.050	1.891
-400	0.387	1.925	1.015	2.057	1.635	2.155	2.251	2.224	2.865	2.267
-1000	0.067	2.466	0.659	2.562	1.251	2.632	1.846	2.680	2.446	2.709
-1000000	-6.508	3.291	-6.012	3.291	-5.490	3.291	-4.943	3.291	-4.373	3.292

**Table A2.** Table A1 – continued for M-type SEDs.

H $\alpha$ EW(Å)	Reddening									
	$E(B - V) = 0.0$		$E(B - V) = 1.0$		$E(B - V) = 2.0$		$E(B - V) = 3.0$		$E(B - V) = 4.0$	
	( $r' - i'$ )	( $r' - H\alpha$ )	( $r' - i'$ )	( $r' - H\alpha$ )	( $r' - i'$ )	( $r' - H\alpha$ )	( $r' - i'$ )	( $r' - H\alpha$ )	( $r' - i'$ )	( $r' - H\alpha$ )
M0V SED										
0	0.929	0.537	1.599	0.699	2.255	0.821	2.898	0.906	3.530	0.958
-10	0.920	0.635	1.589	0.796	2.243	0.917	2.885	1.001	3.517	1.053
-20	0.911	0.723	1.579	0.883	2.232	1.003	2.873	1.086	3.504	1.137
-40	0.894	0.878	1.559	1.035	2.210	1.153	2.849	1.234	3.479	1.284
-60	0.877	1.009	1.539	1.164	2.188	1.279	2.825	1.359	3.454	1.408
-100	0.843	1.223	1.500	1.372	2.145	1.484	2.779	1.561	3.406	1.607
-200	0.764	1.590	1.410	1.729	2.045	1.831	2.673	1.901	3.295	1.943
-400	0.621	2.007	1.249	2.127	1.870	2.215	2.486	2.274	3.102	2.310
-1000	0.280	2.525	0.874	2.612	1.468	2.673	2.066	2.714	2.670	2.738
-1000000	-6.354	3.291	-5.847	3.291	-5.314	3.291	-4.757	3.292	-4.177	3.292
M2V SED										
0	1.174	0.666	1.844	0.818	2.500	0.931	3.144	1.009	3.778	1.055
-10	1.164	0.763	1.833	0.913	2.487	1.025	3.130	1.102	3.764	1.148
-20	1.154	0.850	1.821	0.999	2.475	1.110	3.117	1.187	3.750	1.232
-40	1.134	1.003	1.799	1.149	2.450	1.258	3.090	1.332	3.722	1.377
-60	1.115	1.132	1.777	1.276	2.426	1.382	3.065	1.455	3.695	1.498
-100	1.078	1.342	1.734	1.481	2.379	1.583	3.015	1.653	3.643	1.694
-200	0.990	1.700	1.635	1.828	2.270	1.921	2.899	1.984	3.523	2.021
-400	0.833	2.103	1.460	2.213	2.081	2.291	2.699	2.344	3.317	2.375
-1000	0.464	2.595	1.059	2.672	1.654	2.726	2.255	2.761	2.863	2.781
-1000000	-6.239	3.291	-5.722	3.291	-5.180	3.292	-4.615	3.292	-4.027	3.292
M3V SED										
0	1.556	0.867	2.229	1.008	2.887	1.111	3.532	1.178	4.168	1.216
-10	1.544	0.963	2.215	1.102	2.872	1.204	3.516	1.270	4.151	1.307
-20	1.532	1.049	2.202	1.187	2.857	1.287	3.500	1.353	4.135	1.389
-40	1.508	1.198	2.175	1.333	2.828	1.431	3.469	1.495	4.103	1.530
-60	1.485	1.324	2.149	1.456	2.799	1.551	3.439	1.614	4.071	1.649
-100	1.441	1.527	2.099	1.654	2.744	1.745	3.381	1.804	4.012	1.837
-200	1.337	1.871	1.983	1.986	2.619	2.067	3.249	2.120	3.875	2.149
-400	1.155	2.249	1.782	2.346	2.404	2.413	3.024	2.457	3.645	2.480
-1000	0.742	2.697	1.338	2.762	1.936	2.806	2.541	2.835	3.153	2.850
-1000000	-6.064	3.291	-5.534	3.292	-4.979	3.292	-4.403	3.292	-3.806	3.292
M4V SED										
0	1.872	0.945	2.536	1.072	3.183	1.160	3.817	1.214	4.442	1.239
-10	1.860	1.036	2.522	1.162	3.168	1.249	3.801	1.303	4.426	1.327
-20	1.848	1.119	2.508	1.243	3.153	1.329	3.786	1.382	4.410	1.406
-40	1.823	1.262	2.480	1.384	3.123	1.468	3.755	1.519	4.378	1.543
-60	1.800	1.384	2.454	1.503	3.094	1.585	3.725	1.635	4.348	1.658
-100	1.754	1.581	2.403	1.694	3.039	1.773	3.667	1.821	4.289	1.843
-200	1.647	1.914	2.285	2.017	2.913	2.088	3.535	2.130	4.154	2.150
-400	1.460	2.283	2.081	2.369	2.697	2.427	3.311	2.462	3.926	2.478
-1000	1.039	2.718	1.631	2.776	2.227	2.814	2.829	2.836	3.438	2.847
-1000000	-5.784	3.291	-5.250	3.292	-4.693	3.292	-4.114	3.292	-3.514	3.292
M6V SED										
0	2.570	1.127	3.242	1.232	3.899	1.301	4.543	1.338	5.178	1.348
-10	2.553	1.227	3.224	1.331	3.879	1.398	4.523	1.435	5.158	1.445
-20	2.537	1.316	3.206	1.418	3.860	1.485	4.503	1.521	5.138	1.531
-40	2.504	1.468	3.170	1.568	3.823	1.632	4.464	1.667	5.099	1.677
-60	2.473	1.595	3.136	1.691	3.786	1.754	4.427	1.788	5.061	1.797
-100	2.414	1.795	3.071	1.886	3.717	1.945	4.356	1.977	4.989	1.986
-200	2.277	2.121	2.923	2.201	3.562	2.253	4.196	2.281	4.828	2.289
-400	2.047	2.464	2.677	2.528	3.305	2.569	3.933	2.591	4.563	2.597
-1000	1.555	2.842	2.161	2.883	2.773	2.908	3.393	2.922	4.020	2.926
-1000000	-5.404	3.292	-4.838	3.292	-4.252	3.292	-3.646	3.292	-3.023	3.292

## APPENDIX B: PHOTOMETRY OF T TAURI CANDIDATES TOWARDS IC 1396

Table B1: Photometry of T-Tauri candidates towards IC 1396, selected based on their position in the IPHAS ( $r - H\alpha$ ,  $r' - i'$ ) and ( $r'$ ,  $r' - i'$ ) diagrams. The IAU-registered naming convention for the objects is formed by prefixing "IPHAS" to the position string given in column 2. Calibrated IPHAS photometry is given in columns 3-5 and matched 2MASS data is given in columns 6-8. For objects which were previously identified in the survey by Sicilia-Aguilar et al. (2006a) we give the existing identifier in column 9.

	IPHAS ID J[RA(2000)+Dec(2000)]	IPHAS photometry			2MASS magnitudes			Name [SA06] <sup>1</sup>
		$r'$	$r' - i'$	$r' - H\alpha$	$J$	$H$	$K$	
1	J213008.57+575727.4	19.93±0.03	2.04±0.04	1.40±0.04	15.50±0.08	14.85±0.08	14.53±0.09	
2	J213237.11+583241.6	17.85±0.01	1.29±0.02	1.45±0.02	14.56±0.03	13.80±0.04	13.57±0.04	
3	J213317.77+574813.5	14.73±0.00	0.68±0.00	0.96±0.00	12.25±0.02	10.96±0.03	9.83±0.02	
4	J213324.81+575022.9	17.33±0.00	1.01±0.01	0.96±0.01	14.53±0.03	14.02±0.05	13.65±0.04	
5	J213450.88+583534.7	19.61±0.03	1.58±0.04	1.66±0.04	15.74±0.10	14.85±0.11	14.42±0.09	
6	J213451.70+575140.5	18.33±0.01	1.58±0.02	1.28±0.02	14.33±0.03	13.25±0.03	12.65±0.04	
7	J213453.08+575125.2	18.73±0.01	1.52±0.02	1.35±0.02	14.71±0.04	13.48±0.04	12.68±0.03	
8	J213510.14+565022.6	18.04±0.01	1.18±0.01	0.88±0.02	15.15±0.05	14.47±0.06	13.92±0.05	
9	J213510.21+573147.5	17.89±0.01	1.66±0.01	1.09±0.01	14.00±0.02	12.89±0.03	12.35±0.02	
10	J213516.28+572822.2	17.41±0.01	1.39±0.01	1.25±0.01	14.01±0.03	13.07±0.03	12.67±0.03	Tr37 72-1427
11	J213516.88+573242.3	17.98±0.01	1.43±0.01	1.81±0.01	14.44±0.04	13.56±0.04	13.07±0.04	
12	J213517.46+574822.3	16.90±0.00	1.30±0.01	1.04±0.01	13.40±0.03	12.47±0.03	12.04±0.03	Tr37 81-541
13	J213518.61+573409.2	15.76±0.00	1.01±0.00	0.73±0.00	13.22±0.02	12.30±0.03	11.83±0.02	Tr37 73-472
14	J213523.77+560600.3	17.65±0.01	1.97±0.01	1.22±0.01	13.68±0.03	13.13±0.03	12.81±0.02	
15	J213524.40+573008.5	18.22±0.01	1.60±0.01	1.51±0.01	14.22±0.03	13.29±0.04	12.89±0.03	
16	J213524.51+573301.2	17.43±0.01	1.38±0.01	0.98±0.01	14.01±0.03	13.11±0.04	12.67±0.03	Tr37 73-311
17	J213527.23+573130.1	18.24±0.01	1.67±0.01	1.09±0.02	14.26±0.03	13.31±0.03	12.75±0.03	
18	J213528.42+575823.1	18.87±0.02	1.41±0.03	1.93±0.02	15.15±0.06	14.21±0.05	13.60±0.04	
19	J213531.36+573127.8	18.85±0.01	1.90±0.02	1.28±0.02	> 14.53	13.63±0.05	13.22±0.04	
20	J213531.75+582800.4	18.90±0.01	1.54±0.02	1.29±0.02	15.04±0.03	14.01±0.03	13.66±0.03	
21	J213532.89+573159.6	18.51±0.01	1.65±0.02	1.47±0.02	14.64±0.03	13.81±0.04	13.51±0.04	
22	J213545.88+573640.1	18.37±0.01	1.41±0.02	1.90±0.01	14.55±0.04	13.68±0.03	13.14±0.03	
23	J213602.78+575512.8	18.15±0.01	1.74±0.01	1.24±0.01	13.98±0.03	13.12±0.03	12.74±0.03	
24	J213603.15+580204.5	19.30±0.03	2.32±0.03	1.54±0.04	-	-	-	
25	J213607.92+580410.9	18.83±0.02	1.83±0.02	1.23±0.03	14.77±0.05	13.99±0.06	13.69±0.06	
26	J213613.48+581922.6	14.06±0.00	0.64±0.00	0.54±0.00	12.34±0.02	12.08±0.03	11.79±0.02	
27	J213625.07+572750.2	19.35±0.02	1.97±0.03	1.69±0.03	14.95±0.04	14.03±0.05	13.52±0.04	J21362507+5727502
28	J213628.40+574251.7	18.46±0.01	1.63±0.01	1.29±0.02	14.25±0.02	13.15±0.03	12.49±0.02	
29	J213647.43+580442.6	19.31±0.02	2.16±0.02	1.35±0.03	14.47±0.03	13.60±0.04	13.30±0.04	
30	J213647.63+572954.0	17.36±0.00	1.44±0.01	1.14±0.01	> 13.57	12.34±0.04	11.65±0.03	J21364762+5729540
31	J213651.54+574806.7	19.08±0.01	1.62±0.02	1.25±0.03	14.96±0.05	13.85±0.04	13.26±0.03	
32	J213657.46+572730.3	18.69±0.01	1.97±0.02	1.73±0.02	-	-	-	
33	J213657.67+572733.1	15.88±0.00	1.20±0.00	0.92±0.00	12.44	11.33±0.03	10.64±0.02	Tr37 11-2146
34	J213659.09+573905.6	16.65±0.00	1.17±0.01	1.22±0.01	13.57±0.02	12.65±0.03	12.26±0.03	
35	J213659.47+573134.8	18.87±0.01	1.68±0.02	1.62±0.02	14.52±0.03	13.40±0.04	12.77±0.03	J21365947+5731349
36	J213706.50+573231.8	16.60±0.00	1.31±0.00	0.90±0.01	13.32±0.02	12.33±0.03	11.91±0.02	Tr37 14-287
37	J213707.03+572700.8	15.47±0.00	1.04±0.00	0.90±0.00	12.65±0.02	11.75±0.03	11.26±0.02	Tr37 11-2037
38	J213709.09+572548.6	18.33±0.01	1.75±0.01	1.55±0.01	14.56±0.03	13.79±0.04	13.55±0.05	
39	J213709.37+572948.4	17.45±0.01	1.84±0.01	1.12±0.01	13.39±0.02	12.32±0.05	11.84±0.03	
40	J213713.90+572727.1	18.31±0.01	1.57±0.01	1.72±0.01	14.27±0.04	13.15±0.04	12.40±0.03	
41	J213716.35+572640.4	17.74±0.01	1.31±0.01	0.96±0.01	14.10±0.02	12.95±0.03	12.31±0.02	
42	J213717.15+572847.3	19.94±0.04	2.35±0.04	1.55±0.05	15.02±0.05	14.20±0.05	13.93±0.07	
43	J213717.42+572927.3	17.75±0.01	1.45±0.01	0.99±0.01	14.12±0.04	13.14±0.04	12.58±0.04	
44	J213724.10+572411.5	17.34±0.01	1.30±0.01	1.48±0.01	14.20±0.06	13.29±0.08	12.82±0.05	J21372410+5724115
45	J213724.77+572909.0	18.76±0.01	2.15±0.02	1.41±0.02	14.14±0.03	13.37±0.03	13.00±0.03	
46	J213725.50+572745.3	18.09±0.01	1.55±0.01	1.40±0.01	14.35±0.05	13.57±0.05	> 13.12	
47	J213734.65+571657.8	18.25±0.01	1.51±0.01	1.39±0.01	14.76±0.04	13.77±0.05	13.36±0.04	
48	J213735.72+573258.4	19.41±0.02	2.08±0.03	1.36±0.03	14.86±0.04	14.02±0.05	13.62±0.06	
49	J213738.51+573140.8	16.28±0.00	1.13±0.00	0.83±0.00	13.30±0.03	12.23±0.03	11.67	Tr37 14-183
50	J213742.77+573325.1	16.95±0.00	0.85±0.01	0.82±0.01	12.56±0.02	11.31±0.02	10.39±0.02	J21374275+5733250
51	J213742.90+573446.3	19.89±0.04	2.34±0.04	1.69±0.05	14.96±0.04	14.18±0.05	13.76±0.05	
52	J213745.15+571942.3	16.53±0.00	1.09±0.01	0.92±0.01	13.90±0.02	12.97±0.03	12.58±0.03	Tr37 11-383
53	J213746.13+573428.1	19.81±0.03	2.18±0.03	1.78±0.04	14.92±0.05	14.03±0.05	13.64±0.06	
54	J213747.97+573242.3	18.02±0.01	1.59±0.01	1.37±0.01	14.15±0.04	13.14±0.03	12.77±0.03	
55	J213748.94+572321.0	18.03±0.01	1.50±0.01	1.21±0.01	14.66±0.05	13.68±0.06	13.25±0.05	J21374893+5723209
56	J213757.62+572247.7	15.45±0.00	0.96±0.00	0.70±0.00	13.18±0.02	12.22±0.03	11.67±0.02	Tr37 12-1091
57	J213759.26+573616.1	17.17±0.00	1.56±0.01	1.37±0.01	13.32±0.03	12.40±0.03	11.89±0.02	Tr37 13-1238
58	J213808.56+573707.5	18.51±0.01	1.76±0.01	1.66±0.01	14.31±0.03	13.43±0.03	12.90±0.03	Tr37 13-1426
59	J213809.24+572020.0	16.72±0.00	1.13±0.01	0.92±0.01	13.92±0.03	12.96±0.03	12.48±0.02	J21380924+5720198
60	J213810.00+572352.8	19.11±0.02	2.02±0.03	1.52±0.03	14.87±0.07	14.09±0.06	13.70±0.06	
61	J213811.36+585317.2	18.73±0.02	1.55±0.02	1.10±0.02	14.69±0.04	13.63±0.04	13.24±0.03	
62	J213812.64+572033.7	17.80±0.01	1.41±0.01	1.72±0.01	14.53±0.05	13.64±0.04	13.17±0.03	
63	J213816.87+573926.4	16.52±0.00	1.18±0.00	1.19±0.00	-	-	-	
64	J213822.81+574017.3	19.01±0.01	1.88±0.02	1.41±0.02	14.83±0.03	13.78±0.04	13.35±0.03	
65	J213823.39+565301.6	18.97±0.02	2.26±0.02	1.34±0.03	13.99±0.03	13.20±0.04	12.84±0.04	
66	J213823.91+565304.1	16.05±0.00	1.23±0.00	0.85±0.00	12.86±0.04	11.88±0.04	11.45±0.03	
67	J213823.94+572736.3	18.82±0.01	1.76±0.02	1.19±0.02	14.93±0.06	14.03±0.05	13.60±0.05	
68	J213825.90+574204.8	16.47±0.00	1.13±0.00	0.86±0.00	-	-	-	
69	J213827.41+572720.7	16.11±0.00	1.31±0.00	0.90±0.00	12.86±0.03	11.99±0.03	11.51±0.02	Tr37 12-2113
70	J213827.42+573108.2	14.92±0.00	0.90±0.00	0.96±0.00	12.35±0.03	11.36±0.04	10.77±0.02	Tr37 13-236
71	J213830.30+573255.2	19.09±0.02	1.95±0.02	1.31±0.03	14.96±0.05	14.09±0.05	13.61±0.04	

<sup>1</sup> [SA06]: identifier from Sicilia-Aguilar et al. (2006a).

**Table B1 – continued**

	IPHAS ID J[RA(2000)+Dec(2000)]	IPHAS photometry			2MASS magnitudes			Name [SA06] <sup>1</sup>
		$r'$	$r' - i'$	$r' - H\alpha$	$J$	$H$	$K$	
72	J213832.15+572636.0	19.13±0.02	1.54±0.03	1.31±0.03	14.81±0.04	13.86±0.04	13.18±0.03	J21383216+5726359
73	J213837.43+573020.7	18.04±0.01	1.56±0.01	1.10±0.01	14.67±0.04	13.76±0.04	13.50±0.04	
74	J213840.02+573930.4	17.91±0.01	1.49±0.01	1.13±0.02	14.51±0.03	13.60±0.04	13.22±0.03	Tr37 13-1891
75	J213844.45+571809.0	15.87±0.00	0.95±0.00	1.03±0.00	13.42±0.02	12.65±0.03	12.18±0.03	Tr37 54-1547
76	J213855.21+572237.5	19.95±0.03	1.97±0.04	1.83±0.04	15.83±0.09	15.42±0.14	14.57±0.10	
77	J213858.08+574334.4	16.26±0.00	1.12±0.00	0.92±0.00	13.39±0.03	12.52±0.04	12.04±0.02	Tr37 91-506
78	J213903.21+573042.1	17.76±0.01	1.43±0.01	1.03±0.01	14.34±0.04	13.50±0.05	13.15±0.04	
79	J213916.91+565524.3	19.97±0.05	1.65±0.06	2.22±0.06	15.69±0.07	14.84±0.10	14.46±0.10	
80	J213928.42+580905.7	16.00±0.00	1.17±0.00	0.91±0.00	12.64±0.02	11.61±0.03	10.99±0.02	
81	J213929.40+570630.7	18.28±0.01	1.64±0.02	1.20±0.02	14.50±0.03	13.73±0.04	13.33±0.03	
82	J213930.87+572227.3	18.94±0.02	1.66±0.03	1.39±0.03	14.47±0.04	13.53±0.05	12.86±0.03	
83	J213935.61+571821.9	17.81±0.01	1.40±0.01	1.39±0.01	14.50±0.04	13.58±0.04	13.22±0.04	Tr37 21-33
84	J213938.83+575451.4	19.79±0.05	1.67±0.06	2.19±0.05	14.76±0.04	13.62±0.04	13.03±0.03	
85	J213942.89+560036.9	16.69±0.00	0.89±0.01	0.69±0.01	14.39±0.05	> 13.56	> 13.27	
86	J213945.87+573051.8	18.05±0.01	1.11±0.02	1.90±0.01	14.91±0.04	13.89±0.04	13.15±0.03	
87	J213949.19+581436.7	19.52±0.04	2.16±0.04	1.39±0.05	14.59±0.02	13.61±0.04	12.93±0.02	
88	J213950.05+565350.4	19.12±0.03	1.65±0.03	1.18±0.04	15.39±0.06	14.60±0.06	14.02±0.07	
89	J213950.27+580822.7	18.59±0.02	1.80±0.03	1.20±0.03	14.54±0.04	13.72±0.05	13.35±0.05	
90	J213951.88+572658.3	19.52±0.04	2.22±0.04	1.48±0.05	14.93±0.04	14.16±0.05	13.56±0.04	
91	J213952.15+581413.0	17.31±0.01	1.31±0.01	1.35±0.01	13.90±0.03	12.86±0.04	12.29±0.03	
92	J213952.38+575618.9	16.87±0.00	1.07±0.01	1.23±0.01	> 12.85	11.79:	10.76±0.02	Tr37 93-361
93	J213954.07+572933.5	17.89±0.01	1.43±0.01	1.14±0.01	14.73±0.04	13.87±0.04	13.28±0.05	
94	J213955.69+571638.2	18.20±0.01	1.48±0.02	1.30±0.02	14.78±0.05	13.94±0.05	13.73±0.05	
95	J213956.13+572707.9	19.57±0.04	1.75±0.05	2.00±0.04	15.51±0.06	14.64±0.07	14.22±0.07	
96	J213957.70+572343.3	19.75±0.04	1.81±0.05	1.74±0.05	15.67±0.06	14.86±0.09	14.56±0.11	
97	J213958.44+581214.8	17.49±0.01	1.70±0.01	1.27±0.01	13.42±0.03	12.35±0.03	11.76±0.02	
98	J214004.53+572836.3	15.59±0.00	0.97±0.00	0.84±0.00	13.01±0.03	12.21±0.04	11.81±0.03	
99	J214004.54+581511.6	19.47±0.04	2.02±0.04	1.47±0.05	14.73±0.04	13.67±0.04	13.05±0.03	
100	J214011.35+573951.8	17.82±0.01	1.54±0.01	1.17±0.01	14.31±0.02	13.44±0.04	12.97±0.03	Tr37 24-1736
101	J214011.85+574012.0	17.00±0.00	1.11±0.01	1.28±0.01	13.88±0.03	13.08±0.04	12.73±0.04	Tr37 24-1796
102	J214020.04+572344.3	19.63±0.04	2.38±0.04	1.48±0.05	14.49±0.05	13.58±0.07	13.27±0.04	
103	J214021.30+572657.9	17.08±0.00	1.11±0.01	1.04±0.01	14.50±0.04	13.40±0.03	12.62±0.03	Tr37 22-2651
104	J214021.75+581445.5	19.42±0.04	2.34±0.04	1.47±0.05	14.78±0.03	13.77±0.05	13.04±0.03	
105	J214025.37+573416.2	18.65±0.02	1.62±0.02	1.72±0.02	14.90±0.03	14.03±0.04	13.61±0.05	
106	J214027.32+581421.2	18.37±0.01	1.50±0.02	1.59±0.02	14.30±0.04	13.30±0.04	12.88±0.04	
107	J214035.33+572309.6	16.64±0.00	1.18±0.01	1.12±0.01	-	-	-	
108	J214035.75+573455.1	16.15±0.00	1.08±0.01	0.75±0.01	13.48±0.03	12.56±0.03	12.19±0.03	Tr37 23-570
109	J214036.90+581437.9	15.02±0.00	1.16±0.00	1.01±0.00	11.90±0.02	10.89±0.03	10.23±0.02	
110	J214038.83+581302.8	19.64±0.04	2.26±0.05	1.48±0.05	14.77±0.04	13.87±0.04	13.38±0.04	
111	J214040.73+573214.8	19.85±0.06	2.29±0.07	1.69±0.07	15.02±0.04	14.32±0.05	13.78±0.06	
112	J214041.23+581158.5	18.62±0.02	1.84±0.02	1.33±0.02	> 13.59	> 12.85	13.21±0.04	
113	J214042.81+581937.4	19.16±0.03	1.96±0.03	1.76±0.03	13.94±0.03	12.55±0.04	11.64±0.02	
114	J214044.97+574813.8	18.92±0.02	1.67±0.03	2.08±0.03	14.92±0.06	14.03±0.06	13.65±0.06	
115	J214045.04+581115.7	18.95±0.02	1.85±0.03	1.33±0.03	14.82±0.04	13.88±0.05	13.41±0.04	
116	J214051.37+574643.8	18.14±0.01	1.58±0.02	1.12±0.02	-	-	-	
117	J214053.64+563052.4	19.14±0.01	2.01±0.02	1.47±0.02	14.47±0.04	13.65±0.04	13.21±0.04	
118	J214055.94+571759.2	16.92±0.01	1.30±0.01	1.02±0.01	13.81±0.03	12.83±0.03	12.24±0.02	Tr37 53-1561
119	J214059.80+574403.6	14.73±0.00	0.91±0.00	0.71±0.00	12.27±0.02	11.43±0.03	11.00±0.02	
120	J214100.84+572303.0	16.91±0.01	0.94±0.01	0.78±0.01	14.56±0.05	14.02±0.09	13.70±0.06	
121	J214114.98+573814.9	15.67±0.00	0.94±0.01	0.86±0.01	13.14±0.02	12.32±0.03	11.92±0.02	Tr37 23-969
122	J214115.97+580826.5	18.39±0.01	1.56±0.02	1.51±0.02	14.47±0.03	13.31±0.03	12.53±0.02	
123	J214118.67+575152.1	19.87±0.05	2.01±0.05	1.47±0.06	15.53±0.07	14.91±0.08	14.61±0.09	
124	J214128.66+573643.3	18.48±0.02	1.26±0.03	1.50±0.03	14.41±0.02	13.53±0.03	12.82±0.03	Tr37 23-798
125	J214138.29+572337.2	19.10±0.03	2.16±0.03	1.48±0.03	14.52±0.03	13.70±0.04	13.32±0.03	
126	J214140.67+572253.2	14.39±0.00	1.03±0.00	0.74±0.00	11.79±0.02	11.15±0.03	10.97±0.02	
127	J214202.97+564740.1	15.32±0.00	1.01±0.00	0.81±0.00	12.58±0.03	12.10±0.03	11.69±0.02	
128	J214213.53+565747.3	16.03±0.00	0.85±0.01	0.65±0.01	13.74±0.02	13.38±0.03	12.98±0.02	
129	J214216.79+573622.1	15.86±0.00	1.00±0.01	1.05±0.01	13.08±0.02	12.27±0.03	11.95±0.02	
130	J214217.63+565550.2	15.80±0.00	1.16±0.00	0.97±0.00	12.43±0.03	11.53±0.03	11.02±0.02	
131	J214226.57+572952.2	16.28±0.01	1.01±0.01	0.75±0.01	13.75±0.03	13.04±0.04	12.84±0.03	
132	J214253.50+575508.7	17.04±0.01	1.02±0.01	1.38±0.02	14.38±0.03	13.34±0.03	12.59±0.02	
133	J214253.88+574728.9	17.11±0.01	1.21±0.01	1.16±0.01	14.21±0.03	13.17:	12.58±0.05	
134	J214332.26+565930.0	15.56±0.00	1.10±0.00	0.78±0.00	12.54±0.03	11.89±0.03	11.45±0.02	
135	J214401.86+571352.5	17.50±0.01	1.41±0.01	1.09±0.02	13.81±0.02	12.62±0.03	11.83±0.02	
136	J214427.05+581954.7	16.79±0.00	1.22±0.01	1.03±0.01	13.64±0.04	> 12.56	12.28±0.04	
137	J214433.16+573732.7	18.15±0.01	1.47±0.02	1.07±0.02	14.34±0.03	13.44±0.03	12.86±0.02	
138	J214435.59+571805.7	19.43±0.02	2.24±0.03	1.57±0.03	14.24±0.03	13.50±0.04	13.03±0.02	
139	J214441.34+574611.9	17.46±0.01	1.27±0.01	1.02±0.01	14.17±0.02	13.09±0.03	12.39±0.02	
140	J214505.89+571138.7	17.74±0.01	1.99±0.01	1.39±0.01	13.31±0.03	12.01±0.03	11.24±0.02	
141	J214513.05+565307.0	16.64±0.00	1.03±0.01	0.89±0.01	14.08±0.03	13.48±0.03	13.15±0.03	
142	J214513.62+571041.0	17.70±0.01	1.86±0.01	1.68±0.01	13.30±0.03	12.26±0.04	11.75±0.02	
143	J214514.71+572801.8	15.98±0.00	1.01±0.00	0.77±0.01	13.38±0.06	12.86±0.06	12.56±0.05	
144	J214527.86+573039.2	17.52±0.01	1.40±0.02	1.06±0.02	14.06±0.03	13.23±0.04	12.94±0.03	
145	J214537.00+570201.3	14.71±0.00	1.03±0.00	0.87±0.00	11.29±0.02	10.37±0.03	9.65±0.02	
146	J214547.74+564845.8	19.12±0.03	2.15±0.03	1.67±0.04	14.33±0.03	13.40±0.04	12.82±0.02	
147	J214559.43+562549.3	17.48±0.01	1.25±0.01	0.88±0.01	14.17±0.03	13.51±0.04	13.03±0.03	
148	J214600.27+572309.6	17.35±0.01	1.51±0.01	1.09±0.01	12.65±0.03	11.46±0.03	10.63±0.02	
149	J214616.89+583630.6	16.48±0.01	1.62±0.01	1.08±0.01	12.98±0.03	12.33±0.03	12.08±0.03	
150	J214625.99+572828.9	18.51±0.02	1.31±0.03	2.40±0.03	14.07±0.03	13.06±0.03	12.30±0.03	
151	J214627.17+583113.1	13.19±0.00	0.59±0.00	0.53±0.00	11.38±0.03	10.89±0.03	10.77±0.02	
152	J214659.27+582449.0	15.60±0.00	1.07±0.00	0.95±0.00	12.96±0.04	> 12.08	> 11.85	
153	J214713.52+564929.8	14.70±0.00	0.85±0.00	0.66±0.00	12.86±0.02	12.48±0.03	12.06±0.02	



**Table B1 – continued**

IPHAS ID		IPHAS photometry			2MASS magnitudes			Name
J[RA(2000)+Dec(2000)]		$r'$	$r' - i'$	$r' - H\alpha$	$J$	$H$	$K$	[SA06] <sup>1</sup>
154	J214808.09+574626.8	16.05±0.00	1.38±0.01	1.03±0.01	12.61±0.02	11.99±0.03	11.45±0.02	
155	J214811.78+575941.6	17.01±0.01	1.33±0.01	1.16±0.01	12.92±0.04	11.66±0.04	10.84±0.03	
156	J214820.52+570431.0	17.20±0.01	1.26±0.01	0.96±0.01	13.97±0.03	13.27±0.03	12.78±0.03	
157	J214933.64+574502.2	16.82±0.01	1.42±0.01	1.06±0.01	13.14±0.03	12.43±0.03	11.86±0.03	
158	J214942.69+571820.9	15.68±0.00	1.02±0.01	0.75±0.01	12.99±0.03	12.44±0.03	11.98±0.03	

**APPENDIX C: DERIVED PARAMETERS OF IPHAS T TAURI CANDIDATES TOWARDS IC 1396**

Table C1: Derived parameters of the T Tauri candidates found towards IC 1396. Estimates for the age, mass and radius are based on the Siess et al. (2000) model tracks placed in the ( $r'$ ,  $r' - i'$ ) diagram (Fig. 7). The H $\alpha$  equivalent width (EW) is derived from the ( $r - H\alpha$ ,  $r' - i'$ ) diagram and was subsequently used to obtain the absolute line excess luminosity  $L(H\alpha)$  and a corresponding mass accretion rate  $\dot{M}$  (see §3). The parameters were obtained using the assumption that (i) all stars are placed at a distance of 870 pc, (ii) all stars have a reddening  $A_V = 1.56$ , and (iii) there is a direct relation between the H $\alpha$  luminosity and the accretion luminosity. The estimated 1 $\sigma$ -uncertainties that follow from these assumptions are the errors listed in this table, obtained using Monte Carlo simulations (see §7). In the case of the age, the 25% and 75% quartiles are given as error bounds.

	IPHAS ID J[RA(2000)+Dec(2000)]	Age [Myr]	Mass [ $M_{\odot}$ ]	Radius [ $R_{\odot}$ ]	EW(H $\alpha$ ) [Å]	log L(H $\alpha$ ) [ $L_{\odot}$ ]	log $\dot{M}$ [ $M_{\odot} \text{ yr}^{-1}$ ]
1	J213008.57+575727.4	4.0 $^{+0.5}_{-0.5}$	0.28 $\pm$ 0.04	0.91 $\pm$ 0.17	-52 $\pm$ 7	-3.8 $\pm$ 0.2	-9.2 $\pm$ 0.6
2	J213237.11+583241.6	4.0 $^{+0.5}_{+0.4}$	0.53 $\pm$ 0.10	1.13 $\pm$ 0.10	-114 $\pm$ 7	-2.7 $\pm$ 0.2	-8.2 $\pm$ 0.6
3	J213317.77+574813.5	12.9 $^{+6.8}_{+4.7}$	1.46 $\pm$ 0.09	1.69 $\pm$ 0.17	-81 $\pm$ 6	-1.8 $\pm$ 0.2	-7.4 $\pm$ 0.6
4	J213324.81+575022.9	18.7 $^{+4.7}_{+0.9}$	0.85 $\pm$ 0.12	0.98 $\pm$ 0.08	-52 $\pm$ 5	-2.9 $\pm$ 0.2	-8.7 $\pm$ 0.6
5	J213450.88+583534.7	9.2 $^{+1.2}_{+1.5}$	0.34 $\pm$ 0.05	0.71 $\pm$ 0.11	-136 $\pm$ 11	-3.3 $\pm$ 0.2	-8.9 $\pm$ 0.6
6	J213451.70+575140.5	2.6 $^{+0.4}_{+0.4}$	0.41 $\pm$ 0.05	1.15 $\pm$ 0.14	-60 $\pm$ 5	-3.1 $\pm$ 0.2	-8.6 $\pm$ 0.6
7	J213453.08+575125.2	4.9 $^{+0.5}_{+0.7}$	0.41 $\pm$ 0.05	0.94 $\pm$ 0.13	-76 $\pm$ 6	-3.2 $\pm$ 0.2	-8.7 $\pm$ 0.6
8	J213510.14+565022.6	17.1 $^{+3.3}_{+2.6}$	0.68 $\pm$ 0.12	0.86 $\pm$ 0.08	-31 $\pm$ 4	-3.4 $\pm$ 0.2	-9.2 $\pm$ 0.6
9	J213510.21+573147.5	1.6 $^{+0.1}_{+0.1}$	0.43 $\pm$ 0.05	1.43 $\pm$ 0.21	-29 $\pm$ 3	-3.3 $\pm$ 0.2	-8.6 $\pm$ 0.6
10	J213516.28+572822.2	1.8 $^{+0.1}_{+0.3}$	0.52 $\pm$ 0.09	1.43 $\pm$ 0.16	-68 $\pm$ 5	-2.8 $\pm$ 0.2	-8.1 $\pm$ 0.6
11	J213516.88+573242.3	2.2 $^{+0.2}_{+0.3}$	0.43 $\pm$ 0.05	1.22 $\pm$ 0.15	-190 $\pm$ 8	-2.6 $\pm$ 0.2	-7.9 $\pm$ 0.6
12	J213517.46+574822.3	1.6 $^{+0.1}_{+0.1}$	0.63 $\pm$ 0.12	1.60 $\pm$ 0.19	-44 $\pm$ 4	-2.8 $\pm$ 0.2	-8.1 $\pm$ 0.6
13	J213518.61+573409.2	2.3 $^{+0.2}_{+0.5}$	1.12 $\pm$ 0.26	1.85 $\pm$ 0.14	-24 $\pm$ 4	-2.6 $\pm$ 0.2	-8.2 $\pm$ 0.6
14	J213523.77+560600.3	0.5 $^{+0.2}_{+0.3}$	0.31 $\pm$ 0.03	2.60 $\pm$ 0.25	-30 $\pm$ 3	-3.1 $\pm$ 0.2	-8.0 $\pm$ 0.6
15	J213524.40+573008.5	2.1 $^{+0.3}_{+0.4}$	0.41 $\pm$ 0.05	1.23 $\pm$ 0.13	-99 $\pm$ 5	-2.9 $\pm$ 0.2	-8.2 $\pm$ 0.6
16	J213524.51+573301.2	2.1 $^{+0.1}_{+0.2}$	0.55 $\pm$ 0.10	1.38 $\pm$ 0.13	-31 $\pm$ 4	-3.1 $\pm$ 0.2	-8.6 $\pm$ 0.6
17	J213527.23+573130.1	2.1 $^{+0.3}_{+0.3}$	0.41 $\pm$ 0.05	1.25 $\pm$ 0.14	-29 $\pm$ 4	-3.4 $\pm$ 0.2	-8.8 $\pm$ 0.6
18	J213528.42+575823.1	5.7 $^{+0.7}_{+0.9}$	0.38 $\pm$ 0.05	0.87 $\pm$ 0.12	-233 $\pm$ 11	-2.8 $\pm$ 0.2	-8.3 $\pm$ 0.6
19	J213531.36+573127.8	2.2 $^{+0.2}_{+0.4}$	0.34 $\pm$ 0.04	1.25 $\pm$ 0.17	-42 $\pm$ 4	-3.5 $\pm$ 0.2	-8.8 $\pm$ 0.6
20	J213531.75+582800.4	5.6 $^{+0.6}_{+0.8}$	0.40 $\pm$ 0.05	0.89 $\pm$ 0.12	-64 $\pm$ 5	-3.4 $\pm$ 0.2	-8.9 $\pm$ 0.6
21	J213532.89+573159.6	2.6 $^{+0.3}_{+0.4}$	0.39 $\pm$ 0.05	1.14 $\pm$ 0.13	-87 $\pm$ 5	-3.1 $\pm$ 0.2	-8.4 $\pm$ 0.6
22	J213545.88+573640.1	3.3 $^{+0.5}_{+0.6}$	0.41 $\pm$ 0.05	1.06 $\pm$ 0.14	-224 $\pm$ 9	-2.7 $\pm$ 0.2	-8.0 $\pm$ 0.6
23	J213602.78+575512.8	1.6 $^{+0.1}_{+0.1}$	0.39 $\pm$ 0.05	1.46 $\pm$ 0.24	-45 $\pm$ 4	-3.2 $\pm$ 0.2	-8.4 $\pm$ 0.6
24	J213603.15+580204.5	1.4 $^{+0.3}_{+0.4}$	0.25 $\pm$ 0.04	1.88 $\pm$ 0.31	-63 $\pm$ 7	-3.4 $\pm$ 0.2	-8.4 $\pm$ 0.6
25	J213607.92+580410.9	2.5 $^{+0.4}_{+0.3}$	0.35 $\pm$ 0.05	1.19 $\pm$ 0.13	-38 $\pm$ 5	-3.5 $\pm$ 0.2	-8.9 $\pm$ 0.6
26	J213613.48+581922.6	13.8 $^{+2.9}_{+0.7}$	1.62 $\pm$ 0.12	1.80 $\pm$ 0.23	-25 $\pm$ 4	-2.0 $\pm$ 0.2	-7.7 $\pm$ 0.6
27	J213625.07+572750.2	2.6 $^{+0.3}_{+0.3}$	0.29 $\pm$ 0.04	1.14 $\pm$ 0.21	-111 $\pm$ 8	-3.2 $\pm$ 0.2	-8.5 $\pm$ 0.6
28	J213628.40+574251.7	2.7 $^{+0.3}_{+0.4}$	0.40 $\pm$ 0.05	1.13 $\pm$ 0.14	-59 $\pm$ 4	-3.2 $\pm$ 0.2	-8.6 $\pm$ 0.6
29	J213647.43+580442.6	2.0 $^{+0.3}_{+0.3}$	0.29 $\pm$ 0.04	1.32 $\pm$ 0.30	-40 $\pm$ 4	-3.6 $\pm$ 0.2	-8.9 $\pm$ 0.6
30	J213647.63+572954.0	1.6 $^{+0.1}_{+0.1}$	0.51 $\pm$ 0.07	1.52 $\pm$ 0.21	-48 $\pm$ 4	-2.9 $\pm$ 0.2	-8.2 $\pm$ 0.6
31	J213651.54+574806.7	5.5 $^{+0.7}_{+0.8}$	0.38 $\pm$ 0.05	0.88 $\pm$ 0.12	-53 $\pm$ 5	-3.5 $\pm$ 0.2	-9.0 $\pm$ 0.6
32	J213657.46+572730.3	1.5 $^{+0.2}_{+0.2}$	0.31 $\pm$ 0.03	1.56 $\pm$ 0.44	-119 $\pm$ 5	-3.0 $\pm$ 0.2	-8.1 $\pm$ 0.6
33	J213657.67+572733.1	0.9 $^{+0.1}_{+0.0}$	0.69 $\pm$ 0.16	2.40 $\pm$ 0.21	-34 $\pm$ 4	-2.5 $\pm$ 0.2	-7.7 $\pm$ 0.6
34	J213659.09+573905.6	2.0 $^{+0.2}_{+0.1}$	0.73 $\pm$ 0.15	1.58 $\pm$ 0.15	-80 $\pm$ 5	-2.4 $\pm$ 0.2	-7.8 $\pm$ 0.6
35	J213659.47+573134.8	3.3 $^{+0.5}_{+0.5}$	0.35 $\pm$ 0.05	1.05 $\pm$ 0.13	-117 $\pm$ 6	-3.1 $\pm$ 0.2	-8.4 $\pm$ 0.6
36	J213706.50+573231.8	1.2 $^{+0.2}_{+0.1}$	0.60 $\pm$ 0.10	1.92 $\pm$ 0.22	-25 $\pm$ 3	-2.9 $\pm$ 0.2	-8.2 $\pm$ 0.6
37	J213707.03+572700.8	1.2 $^{+0.3}_{+0.1}$	0.98 $\pm$ 0.30	2.33 $\pm$ 0.16	-42 $\pm$ 4	-2.3 $\pm$ 0.2	-7.6 $\pm$ 0.6
38	J213709.09+572548.6	1.8 $^{+0.1}_{+0.3}$	0.36 $\pm$ 0.04	1.42 $\pm$ 0.24	-97 $\pm$ 5	-2.9 $\pm$ 0.2	-8.1 $\pm$ 0.6
39	J213709.37+572948.4	0.8 $^{+0.1}_{+0.1}$	0.33 $\pm$ 0.04	2.60 $\pm$ 0.21	-25 $\pm$ 3	-3.1 $\pm$ 0.2	-8.1 $\pm$ 0.6
40	J213713.90+572727.1	2.3 $^{+0.3}_{+0.3}$	0.40 $\pm$ 0.05	1.19 $\pm$ 0.13	-151 $\pm$ 6	-2.8 $\pm$ 0.2	-8.1 $\pm$ 0.6
41	J213716.35+572640.4	4.3 $^{+0.3}_{+0.4}$	0.58 $\pm$ 0.12	1.14 $\pm$ 0.09	-33 $\pm$ 4	-3.2 $\pm$ 0.2	-8.8 $\pm$ 0.6
42	J213717.15+572847.3	2.3 $^{+0.2}_{+0.2}$	0.22 $\pm$ 0.04	1.46 $\pm$ 0.37	-63 $\pm$ 9	-3.7 $\pm$ 0.2	-8.8 $\pm$ 0.6
43	J213717.42+572927.3	2.3 $^{+0.2}_{+0.2}$	0.50 $\pm$ 0.08	1.29 $\pm$ 0.13	-28 $\pm$ 4	-3.3 $\pm$ 0.2	-8.7 $\pm$ 0.6
44	J213724.10+572411.5	2.1 $^{+0.2}_{+0.2}$	0.55 $\pm$ 0.10	1.39 $\pm$ 0.13	-120 $\pm$ 6	-2.5 $\pm$ 0.2	-7.9 $\pm$ 0.6
45	J213724.77+572909.0	1.3 $^{+0.3}_{+0.2}$	0.29 $\pm$ 0.03	1.79 $\pm$ 0.43	-49 $\pm$ 4	-3.3 $\pm$ 0.2	-8.4 $\pm$ 0.6
46	J213725.50+572745.3	2.2 $^{+0.2}_{+0.3}$	0.43 $\pm$ 0.05	1.23 $\pm$ 0.15	-82 $\pm$ 5	-2.9 $\pm$ 0.2	-8.3 $\pm$ 0.6
47	J213734.65+571657.8	2.8 $^{+0.3}_{+0.5}$	0.43 $\pm$ 0.05	1.14 $\pm$ 0.15	-84 $\pm$ 5	-3.0 $\pm$ 0.2	-8.4 $\pm$ 0.6
48	J213735.72+573258.4	2.5 $^{+0.4}_{+0.2}$	0.29 $\pm$ 0.04	1.18 $\pm$ 0.23	-44 $\pm$ 5	-3.6 $\pm$ 0.2	-9.0 $\pm$ 0.6
49	J213738.51+573140.8	2.1 $^{+0.2}_{+0.1}$	0.87 $\pm$ 0.20	1.71 $\pm$ 0.16	-28 $\pm$ 4	-2.7 $\pm$ 0.2	-8.2 $\pm$ 0.6
50	J213742.77+573325.1	32.2 $^{+2.9}_{+4.0}$	0.89 $\pm$ 0.07	0.92 $\pm$ 0.08	-44 $\pm$ 4	-2.9 $\pm$ 0.2	-8.7 $\pm$ 0.6
51	J213742.90+573446.3	2.2 $^{+0.3}_{+0.2}$	0.22 $\pm$ 0.04	1.53 $\pm$ 0.37	-91 $\pm$ 10	-3.5 $\pm$ 0.2	-8.6 $\pm$ 0.6
52	J213745.15+571942.3	3.5 $^{+0.4}_{+0.3}$	0.89 $\pm$ 0.20	1.49 $\pm$ 0.12	-42 $\pm$ 4	-2.7 $\pm$ 0.2	-8.2 $\pm$ 0.6
53	J213746.13+573428.1	2.6 $^{+0.3}_{+0.2}$	0.24 $\pm$ 0.05	1.31 $\pm$ 0.33	-119 $\pm$ 10	-3.4 $\pm$ 0.2	-8.5 $\pm$ 0.6
54	J213747.97+573242.3	1.9 $^{+0.2}_{+0.3}$	0.43 $\pm$ 0.05	1.30 $\pm$ 0.16	-75 $\pm$ 4	-2.9 $\pm$ 0.2	-8.3 $\pm$ 0.6
55	J213748.94+572321.0	2.5 $^{+0.3}_{+0.3}$	0.45 $\pm$ 0.06	1.22 $\pm$ 0.14	-55 $\pm$ 4	-3.1 $\pm$ 0.2	-8.5 $\pm$ 0.6
56	J213757.62+572247.7	2.1 $^{+0.3}_{+0.8}$	1.24 $\pm$ 0.28	2.00 $\pm$ 0.14	-22 $\pm$ 4	-2.5 $\pm$ 0.2	-8.1 $\pm$ 0.6
57	J213759.26+573616.1	1.0 $^{+0.1}_{+0.0}$	0.42 $\pm$ 0.05	2.11 $\pm$ 0.29	-76 $\pm$ 4	-2.6 $\pm$ 0.2	-7.7 $\pm$ 0.6
58	J213808.56+573707.5	2.0 $^{+0.3}_{+0.2}$	0.35 $\pm$ 0.04	1.31 $\pm$ 0.18	-120 $\pm$ 5	-2.9 $\pm$ 0.2	-8.2 $\pm$ 0.6

Table C1 – continued

	IPHAS ID J[RA(2000)+Dec(2000)]	Age [Myr]	Mass [ $M_{\odot}$ ]	Radius [ $R_{\odot}$ ]	EW(H $\alpha$ ) [Å]	log L(H $\alpha$ ) [ $L_{\odot}$ ]	log M [ $M_{\odot} \text{ yr}^{-1}$ ]
59	J213809.24+572020.0	$3.5^{+0.5}_{-0.3}$	$0.83 \pm 0.19$	$1.43 \pm 0.12$	$-39 \pm 4$	$-2.8 \pm 0.2$	$-8.3 \pm 0.6$
60	J213810.00+572352.8	$2.1^{+0.3}_{-0.3}$	$0.30 \pm 0.04$	$1.27 \pm 0.26$	$-73 \pm 6$	$-3.3 \pm 0.2$	$-8.6 \pm 0.6$
61	J213811.36+585317.2	$4.9^{+0.5}_{-0.6}$	$0.41 \pm 0.06$	$0.95 \pm 0.13$	$-36 \pm 5$	$-3.5 \pm 0.2$	$-9.1 \pm 0.6$
62	J213812.64+572033.7	$2.1^{+0.2}_{-0.3}$	$0.45 \pm 0.06$	$1.29 \pm 0.15$	$-168 \pm 7$	$-2.5 \pm 0.2$	$-7.9 \pm 0.6$
63	J213816.87+573926.4	$1.6^{+0.1}_{-0.1}$	$0.72 \pm 0.15$	$1.69 \pm 0.18$	$-75 \pm 5$	$-2.4 \pm 0.2$	$-7.8 \pm 0.6$
64	J213822.81+574017.3	$2.5^{+0.3}_{-0.3}$	$0.33 \pm 0.04$	$1.17 \pm 0.16$	$-62 \pm 5$	$-3.4 \pm 0.2$	$-8.7 \pm 0.6$
65	J213823.39+565301.6	$1.0^{+0.1}_{-0.4}$	$0.27 \pm 0.03$	$1.90 \pm 0.32$	$-35 \pm 4$	$-3.5 \pm 0.2$	$-8.6 \pm 0.6$
66	J213823.91+565304.1	$1.0^{+0.0}_{-0.4}$	$0.67 \pm 0.14$	$2.27 \pm 0.21$	$-24 \pm 4$	$-2.7 \pm 0.2$	$-7.9 \pm 0.6$
67	J213823.94+572736.3	$2.9^{+0.4}_{-0.4}$	$0.36 \pm 0.05$	$1.10 \pm 0.12$	$-36 \pm 4$	$-3.5 \pm 0.2$	$-8.9 \pm 0.6$
68	J213825.90+574204.8	$2.5^{+0.3}_{-0.1}$	$0.84 \pm 0.19$	$1.61 \pm 0.14$	$-31 \pm 4$	$-2.8 \pm 0.2$	$-8.3 \pm 0.6$
69	J213827.41+572720.7	$0.8^{+0.1}_{-0.1}$	$0.58 \pm 0.12$	$2.48 \pm 0.22$	$-25 \pm 3$	$-2.7 \pm 0.2$	$-7.8 \pm 0.6$
70	J213827.42+573108.2	$1.6^{+0.3}_{-0.5}$	$1.35 \pm 0.36$	$2.40 \pm 0.13$	$-63 \pm 5$	$-1.9 \pm 0.2$	$-7.3 \pm 0.6$
71	J213830.30+573255.2	$2.4^{+0.3}_{-0.3}$	$0.32 \pm 0.04$	$1.19 \pm 0.18$	$-44 \pm 4$	$-3.5 \pm 0.2$	$-8.9 \pm 0.6$
72	J213832.15+572636.0	$7.4^{+0.9}_{-1.1}$	$0.39 \pm 0.06$	$0.81 \pm 0.11$	$-68 \pm 7$	$-3.4 \pm 0.2$	$-9.0 \pm 0.6$
73	J213837.43+573020.7	$2.2^{+0.2}_{-0.3}$	$0.44 \pm 0.06$	$1.24 \pm 0.15$	$-36 \pm 4$	$-3.3 \pm 0.2$	$-8.7 \pm 0.6$
74	J213840.02+573930.4	$2.3^{+0.2}_{-0.3}$	$0.47 \pm 0.07$	$1.26 \pm 0.14$	$-44 \pm 4$	$-3.1 \pm 0.2$	$-8.6 \pm 0.6$
75	J213844.45+571809.0	$3.2^{+0.3}_{-0.9}$	$1.12 \pm 0.23$	$1.69 \pm 0.12$	$-68 \pm 5$	$-2.2 \pm 0.2$	$-7.8 \pm 0.6$
76	J213855.21+572237.5	$4.1^{+0.5}_{-0.2}$	$0.26 \pm 0.05$	$0.91 \pm 0.17$	$-146 \pm 11$	$-3.4 \pm 0.2$	$-8.7 \pm 0.6$
77	J213858.08+574334.4	$2.0^{+0.1}_{-0.5}$	$0.86 \pm 0.20$	$1.73 \pm 0.16$	$-40 \pm 4$	$-2.6 \pm 0.2$	$-8.1 \pm 0.6$
78	J213903.21+573042.1	$2.5^{+0.2}_{-0.2}$	$0.50 \pm 0.09$	$1.27 \pm 0.13$	$-35 \pm 4$	$-3.2 \pm 0.2$	$-8.6 \pm 0.6$
79	J213916.91+565524.3	$6.8^{+1.1}_{-0.2}$	$0.27 \pm 0.05$	$0.72 \pm 0.15$	$-335 \pm 34$	$-3.1 \pm 0.2$	$-8.6 \pm 0.6$
80	J213928.42+580905.7	$1.1^{+0.1}_{-0.2}$	$0.74 \pm 0.18$	$2.16 \pm 0.19$	$-35 \pm 4$	$-2.5 \pm 0.2$	$-7.8 \pm 0.6$
81	J213929.40+570630.7	$2.2^{+0.2}_{-0.3}$	$0.41 \pm 0.05$	$1.21 \pm 0.14$	$-44 \pm 4$	$-3.2 \pm 0.2$	$-8.6 \pm 0.6$
82	J213930.87+572227.3	$4.1^{+0.6}_{-0.6}$	$0.37 \pm 0.05$	$0.98 \pm 0.13$	$-74 \pm 6$	$-3.3 \pm 0.2$	$-8.7 \pm 0.6$
83	J213935.61+571821.9	$2.5^{+0.3}_{-0.2}$	$0.48 \pm 0.08$	$1.25 \pm 0.14$	$-92 \pm 6$	$-2.8 \pm 0.2$	$-8.2 \pm 0.6$
84	J213938.83+575451.4	$5.6^{+0.8}_{-1.0}$	$0.28 \pm 0.05$	$0.79 \pm 0.16$	$-317 \pm 29$	$-3.1 \pm 0.2$	$-8.5 \pm 0.6$
85	J213942.89+560036.9	$19.4^{+1.4}_{-4.5}$	$0.97 \pm 0.09$	$1.05 \pm 0.08$	$-25 \pm 4$	$-3.0 \pm 0.2$	$-8.8 \pm 0.6$
86	J213945.87+573051.8	$8.1^{+0.8}_{-1.0}$	$0.56 \pm 0.12$	$0.93 \pm 0.08$	$-270 \pm 12$	$-2.5 \pm 0.2$	$-8.1 \pm 0.6$
87	J213949.19+581436.7	$2.3^{+0.3}_{-0.3}$	$0.27 \pm 0.05$	$1.26 \pm 0.29$	$-46 \pm 8$	$-3.7 \pm 0.2$	$-8.9 \pm 0.6$
88	J213950.05+565350.4	$5.4^{+0.8}_{-0.7}$	$0.37 \pm 0.05$	$0.88 \pm 0.13$	$-42 \pm 6$	$-3.6 \pm 0.2$	$-9.2 \pm 0.6$
89	J213950.27+580822.7	$2.1^{+0.3}_{-0.3}$	$0.36 \pm 0.05$	$1.25 \pm 0.14$	$-36 \pm 5$	$-3.4 \pm 0.2$	$-8.8 \pm 0.6$
90	J213951.88+572658.3	$2.1^{+0.3}_{-0.3}$	$0.26 \pm 0.05$	$1.39 \pm 0.34$	$-56 \pm 8$	$-3.6 \pm 0.2$	$-8.8 \pm 0.6$
91	J213952.15+581413.0	$2.1^{+0.2}_{-0.2}$	$0.56 \pm 0.11$	$1.40 \pm 0.13$	$-92 \pm 5$	$-2.6 \pm 0.2$	$-8.0 \pm 0.6$
92	J213952.38+575618.9	$5.2^{+0.6}_{-0.6}$	$0.84 \pm 0.18$	$1.30 \pm 0.11$	$-93 \pm 6$	$-2.5 \pm 0.2$	$-8.1 \pm 0.6$
93	J213954.07+572933.5	$2.6^{+0.2}_{-0.3}$	$0.49 \pm 0.08$	$1.23 \pm 0.13$	$-48 \pm 4$	$-3.1 \pm 0.2$	$-8.5 \pm 0.6$
94	J213955.69+571638.2	$3.1^{+0.3}_{-0.5}$	$0.44 \pm 0.06$	$1.13 \pm 0.14$	$-71 \pm 5$	$-3.0 \pm 0.2$	$-8.5 \pm 0.6$
95	J213956.13+572707.9	$4.5^{+0.6}_{-0.6}$	$0.29 \pm 0.04$	$0.88 \pm 0.18$	$-220 \pm 18$	$-3.1 \pm 0.2$	$-8.5 \pm 0.6$
96	J213957.70+572343.3	$5.2^{+0.6}_{-0.9}$	$0.30 \pm 0.04$	$0.83 \pm 0.17$	$-134 \pm 14$	$-3.4 \pm 0.2$	$-8.8 \pm 0.6$
97	J213958.44+581214.8	$1.0^{+0.0}_{-0.3}$	$0.36 \pm 0.06$	$2.27 \pm 0.29$	$-51 \pm 4$	$-2.9 \pm 0.2$	$-7.9 \pm 0.6$
98	J214004.53+572836.3	$2.1^{+0.1}_{-0.6}$	$1.17 \pm 0.27$	$1.93 \pm 0.14$	$-39 \pm 4$	$-2.4 \pm 0.2$	$-7.9 \pm 0.6$
99	J214004.54+581511.6	$2.8^{+0.3}_{-0.4}$	$0.29 \pm 0.04$	$1.10 \pm 0.21$	$-65 \pm 8$	$-3.5 \pm 0.2$	$-8.9 \pm 0.6$
100	J214011.35+573951.8	$1.8^{+0.2}_{-0.3}$	$0.46 \pm 0.06$	$1.35 \pm 0.17$	$-46 \pm 4$	$-3.1 \pm 0.2$	$-8.4 \pm 0.6$
101	J214011.85+574012.0	$4.3^{+0.6}_{-0.5}$	$0.77 \pm 0.17$	$1.30 \pm 0.11$	$-97 \pm 6$	$-2.5 \pm 0.2$	$-8.0 \pm 0.6$
102	J214020.04+575044.3	$1.8^{+0.3}_{-0.2}$	$0.23 \pm 0.04$	$1.74 \pm 0.36$	$-50 \pm 8$	$-3.6 \pm 0.2$	$-8.7 \pm 0.6$
103	J214021.30+572657.9	$6.1^{+0.2}_{-0.8}$	$0.81 \pm 0.17$	$1.22 \pm 0.10$	$-58 \pm 5$	$-2.8 \pm 0.2$	$-8.4 \pm 0.6$
104	J214021.75+581445.5	$1.6^{+0.3}_{-0.3}$	$0.24 \pm 0.04$	$1.79 \pm 0.34$	$-50 \pm 7$	$-3.6 \pm 0.2$	$-8.6 \pm 0.6$
105	J214025.37+573416.2	$2.9^{+0.4}_{-0.5}$	$0.37 \pm 0.05$	$1.10 \pm 0.12$	$-147 \pm 8$	$-2.9 \pm 0.2$	$-8.2 \pm 0.6$
106	J214027.32+581421.2	$3.1^{+0.4}_{-0.5}$	$0.41 \pm 0.05$	$1.09 \pm 0.14$	$-125 \pm 7$	$-2.9 \pm 0.2$	$-8.3 \pm 0.6$
107	J214035.33+572309.6	$2.0^{+0.2}_{-0.1}$	$0.74 \pm 0.15$	$1.59 \pm 0.15$	$-63 \pm 5$	$-2.5 \pm 0.2$	$-8.0 \pm 0.6$
108	J214035.75+573455.1	$2.5^{+0.3}_{-0.1}$	$0.97 \pm 0.23$	$1.70 \pm 0.14$	$-22 \pm 4$	$-2.8 \pm 0.2$	$-8.4 \pm 0.6$
109	J214036.90+581437.9	$0.5^{+0.0}_{-0.0}$	$0.73 \pm 0.23$	$3.45 \pm 0.28$	$-48 \pm 4$	$-2.0 \pm 0.2$	$-7.0 \pm 0.6$
110	J214038.83+581302.8	$2.2^{+0.3}_{-0.3}$	$0.25 \pm 0.05$	$1.41 \pm 0.35$	$-56 \pm 9$	$-3.6 \pm 0.2$	$-8.8 \pm 0.6$
111	J214040.73+573214.8	$2.3^{+0.3}_{-0.3}$	$0.23 \pm 0.05$	$1.45 \pm 0.37$	$-92 \pm 15$	$-3.5 \pm 0.2$	$-8.6 \pm 0.6$
112	J214041.23+581158.5	$2.0^{+0.3}_{-0.2}$	$0.35 \pm 0.04$	$1.31 \pm 0.19$	$-53 \pm 5$	$-3.3 \pm 0.2$	$-8.6 \pm 0.6$
113	J214042.81+581937.4	$2.3^{+0.2}_{-0.3}$	$0.30 \pm 0.04$	$1.22 \pm 0.24$	$-128 \pm 9$	$-3.1 \pm 0.2$	$-8.4 \pm 0.6$
114	J214044.97+574813.8	$2.8^{+0.4}_{-0.4}$	$0.32 \pm 0.04$	$1.13 \pm 0.17$	$-262 \pm 14$	$-2.8 \pm 0.2$	$-8.1 \pm 0.6$
115	J214045.04+581115.7	$2.6^{+0.3}_{-0.4}$	$0.34 \pm 0.04$	$1.17 \pm 0.15$	$-51 \pm 6$	$-3.4 \pm 0.2$	$-8.8 \pm 0.6$
116	J214051.37+574643.8	$2.3^{+0.3}_{-0.2}$	$0.43 \pm 0.05$	$1.21 \pm 0.15$	$-38 \pm 4$	$-3.3 \pm 0.2$	$-8.7 \pm 0.6$
117	J214053.64+563052.4	$2.2^{+0.2}_{-0.3}$	$0.31 \pm 0.04$	$1.23 \pm 0.23$	$-66 \pm 4$	$-3.4 \pm 0.2$	$-8.7 \pm 0.6$
118	J214055.94+571759.2	$1.6^{+0.1}_{-0.1}$	$0.62 \pm 0.12$	$1.58 \pm 0.18$	$-41 \pm 4$	$-2.8 \pm 0.2$	$-8.2 \pm 0.6$
119	J214059.80+574403.6	$1.3^{+0.4}_{-0.7}$	$1.39 \pm 0.38$	$2.62 \pm 0.15$	$-26 \pm 4$	$-2.2 \pm 0.2$	$-7.6 \pm 0.6$
120	J214100.84+572303.0	$18.5^{+1.8}_{-1.1}$	$0.94 \pm 0.10$	$1.04 \pm 0.08$	$-33 \pm 4$	$-3.0 \pm 0.2$	$-8.8 \pm 0.6$
121	J214114.98+573814.9	$3.0^{+0.4}_{-1.1}$	$1.20 \pm 0.24$	$1.79 \pm 0.12$	$-44 \pm 5$	$-2.3 \pm 0.2$	$-7.9 \pm 0.6$
122	J214115.97+580826.5	$2.8^{+0.4}_{-0.5}$	$0.41 \pm 0.05$	$1.12 \pm 0.14$	$-103 \pm 6$	$-3.0 \pm 0.2$	$-8.3 \pm 0.6$
123	J214118.67+575152.1	$4.0^{+0.5}_{-0.5}$	$0.28 \pm 0.04$	$0.92 \pm 0.19$	$-65 \pm 11$	$-3.7 \pm 0.2$	$-9.1 \pm 0.6$
124	J214128.66+573643.3	$9.8^{+1.0}_{-1.6}$	$0.51 \pm 0.11$	$0.84 \pm 0.08$	$-128 \pm 9$	$-2.9 \pm 0.2$	$-8.6 \pm 0.6$

**Table C1 – continued**

	IPHAS ID J[RA(2000)+Dec(2000)]	Age [Myr]	Mass [ $M_{\odot}$ ]	Radius [ $R_{\odot}$ ]	EW(H $\alpha$ ) [Å]	log L(H $\alpha$ ) [ $L_{\odot}$ ]	log M [ $M_{\odot} \text{ yr}^{-1}$ ]
125	J214138.29+572337.2	$1.6^{+0.3}_{-0.2}$	$0.29 \pm 0.04$	$1.47 \pm 0.37$	$-60 \pm 6$	$-3.4 \pm 0.2$	$-8.6 \pm 0.6$
126	J214140.67+572253.2	$0.5^{+0.1}_{-0.0}$	$1.17 \pm 0.43$	$3.64 \pm 0.25$	$-23 \pm 4$	$-2.1 \pm 0.2$	$-7.3 \pm 0.6$
127	J214202.97+564740.1	$1.3^{+0.1}_{-0.3}$	$1.09 \pm 0.34$	$2.34 \pm 0.15$	$-33 \pm 4$	$-2.3 \pm 0.2$	$-7.7 \pm 0.6$
128	J214213.53+565747.3	$11.7^{+3.6}_{-2.0}$	$1.14 \pm 0.10$	$1.30 \pm 0.08$	$-23 \pm 3$	$-2.8 \pm 0.2$	$-8.6 \pm 0.6$
129	J214216.79+573622.1	$2.2^{+0.3}_{-0.2}$	$1.04 \pm 0.25$	$1.81 \pm 0.14$	$-68 \pm 5$	$-2.2 \pm 0.2$	$-7.7 \pm 0.6$
130	J214217.63+565550.2	$1.0^{+0.0}_{-0.0}$	$0.74 \pm 0.19$	$2.37 \pm 0.21$	$-44 \pm 4$	$-2.3 \pm 0.2$	$-7.6 \pm 0.6$
131	J214226.57+572952.2	$4.7^{+1.0}_{-0.5}$	$1.05 \pm 0.20$	$1.48 \pm 0.11$	$-25 \pm 4$	$-2.8 \pm 0.2$	$-8.5 \pm 0.6$
132	J214253.50+575508.7	$7.9^{+1.0}_{-1.0}$	$0.84 \pm 0.16$	$1.17 \pm 0.10$	$-128 \pm 8$	$-2.4 \pm 0.2$	$-8.0 \pm 0.6$
133	J214253.88+574728.9	$2.9^{+0.3}_{-0.2}$	$0.68 \pm 0.15$	$1.36 \pm 0.12$	$-68 \pm 5$	$-2.7 \pm 0.2$	$-8.1 \pm 0.6$
134	J214332.26+565930.0	$1.0^{+0.0}_{-0.0}$	$0.88 \pm 0.26$	$2.38 \pm 0.18$	$-23 \pm 4$	$-2.5 \pm 0.2$	$-7.9 \pm 0.6$
135	J214401.86+571352.5	$2.0^{+0.2}_{-0.2}$	$0.53 \pm 0.09$	$1.38 \pm 0.14$	$-43 \pm 4$	$-3.0 \pm 0.2$	$-8.4 \pm 0.6$
136	J214427.05+581954.7	$2.1^{+0.1}_{-0.2}$	$0.70 \pm 0.15$	$1.53 \pm 0.14$	$-48 \pm 4$	$-2.7 \pm 0.2$	$-8.1 \pm 0.6$
137	J214433.16+573732.7	$3.3^{+0.3}_{-0.3}$	$0.47 \pm 0.08$	$1.12 \pm 0.13$	$-37 \pm 5$	$-3.3 \pm 0.2$	$-8.8 \pm 0.6$
138	J214435.59+571805.7	$1.9^{+0.3}_{-0.3}$	$0.26 \pm 0.05$	$1.50 \pm 0.36$	$-71 \pm 6$	$-3.4 \pm 0.2$	$-8.6 \pm 0.6$
139	J214441.34+574611.9	$3.5^{+0.4}_{-0.4}$	$0.62 \pm 0.13$	$1.24 \pm 0.10$	$-44 \pm 4$	$-3.0 \pm 0.2$	$-8.5 \pm 0.6$
140	J214505.89+571138.7	$0.4^{+0.3}_{-0.3}$	$0.31 \pm 0.02$	$2.55 \pm 0.27$	$-54 \pm 3$	$-2.9 \pm 0.2$	$-7.8 \pm 0.6$
141	J214513.05+565307.0	$6.3^{+0.7}_{-0.7}$	$0.96 \pm 0.18$	$1.32 \pm 0.11$	$-42 \pm 4$	$-2.7 \pm 0.2$	$-8.4 \pm 0.6$
142	J214513.62+571041.0	$0.9^{+0.2}_{-0.1}$	$0.31 \pm 0.02$	$2.50 \pm 0.29$	$-117 \pm 5$	$-2.6 \pm 0.2$	$-7.5 \pm 0.6$
143	J214514.71+572801.8	$3.0^{+0.2}_{-0.6}$	$1.08 \pm 0.24$	$1.70 \pm 0.12$	$-28 \pm 4$	$-2.7 \pm 0.2$	$-8.2 \pm 0.6$
144	J214527.86+573039.2	$2.1^{+0.2}_{-0.2}$	$0.53 \pm 0.09$	$1.36 \pm 0.13$	$-39 \pm 4$	$-3.0 \pm 0.2$	$-8.5 \pm 0.6$
145	J214537.00+570201.3	$0.6^{+0.1}_{-0.0}$	$1.08 \pm 0.36$	$3.16 \pm 0.24$	$-39 \pm 4$	$-2.0 \pm 0.2$	$-7.2 \pm 0.6$
146	J214547.74+564845.8	$1.6^{+0.2}_{-0.3}$	$0.28 \pm 0.04$	$1.52 \pm 0.38$	$-94 \pm 8$	$-3.2 \pm 0.2$	$-8.3 \pm 0.6$
147	J214559.43+562549.3	$4.5^{+0.5}_{-0.6}$	$0.66 \pm 0.14$	$1.18 \pm 0.10$	$-27 \pm 4$	$-3.2 \pm 0.2$	$-8.8 \pm 0.6$
148	J214600.27+572309.6	$1.3^{+0.1}_{-0.3}$	$0.48 \pm 0.05$	$1.62 \pm 0.27$	$-37 \pm 4$	$-3.0 \pm 0.2$	$-8.3 \pm 0.6$
149	J214616.89+583630.6	$0.2^{+0.1}_{-0.3}$	$0.40 \pm 0.04$	$2.93 \pm 0.31$	$-30 \pm 4$	$-2.7 \pm 0.2$	$-7.6 \pm 0.6$
150	J214625.99+572828.9	$3.2^{+0.5}_{-0.5}$	$0.35 \pm 0.05$	$1.06 \pm 0.13$	$-525 \pm 27$	$-2.4 \pm 0.2$	$-7.7 \pm 0.6$
151	J214627.17+583113.1	$8.6^{+0.4}_{-2.2}$	$1.91 \pm 0.19$	$2.34 \pm 0.26$	$-28 \pm 4$	$-1.6 \pm 0.2$	$-7.2 \pm 0.6$
152	J214659.27+582449.0	$1.2^{+0.2}_{-0.1}$	$0.90 \pm 0.26$	$2.30 \pm 0.17$	$-47 \pm 4$	$-2.3 \pm 0.2$	$-7.6 \pm 0.6$
153	J214713.52+564929.8	$2.3^{+0.6}_{-1.7}$	$1.54 \pm 0.29$	$2.32 \pm 0.15$	$-25 \pm 4$	$-2.2 \pm 0.2$	$-7.8 \pm 0.6$
154	J214808.09+574626.8	$0.6^{+0.1}_{-0.1}$	$0.50 \pm 0.09$	$2.89 \pm 0.31$	$-37 \pm 4$	$-2.5 \pm 0.2$	$-7.5 \pm 0.6$
155	J214811.78+575941.6	$1.6^{+0.1}_{-0.0}$	$0.58 \pm 0.10$	$1.59 \pm 0.20$	$-58 \pm 5$	$-2.7 \pm 0.2$	$-8.0 \pm 0.6$
156	J214820.52+570431.0	$2.8^{+0.3}_{-0.2}$	$0.65 \pm 0.14$	$1.35 \pm 0.12$	$-36 \pm 4$	$-3.0 \pm 0.2$	$-8.5 \pm 0.6$
157	J214933.64+574502.2	$1.0^{+0.1}_{-0.2}$	$0.50 \pm 0.06$	$1.99 \pm 0.29$	$-38 \pm 4$	$-2.8 \pm 0.2$	$-8.0 \pm 0.6$
158	J214942.69+571820.9	$1.9^{+0.1}_{-0.3}$	$1.09 \pm 0.28$	$1.96 \pm 0.15$	$-25 \pm 4$	$-2.6 \pm 0.2$	$-8.1 \pm 0.6$

## APPENDIX D: SPITZER IRAC AND MIPS PHOTOMETRY

Table D1: Spitzer IRAC and MIPS magnitudes for the T Tauri candidates in the central region (Spitzer data does not cover the entire IC 1396 region). The uncertainties listed include the background noise, but not the  $\sim 10$  per cent uncertainty in the absolute calibration which is typical for Spitzer. See §4 for details on the data reduction.

IPHAS ID	3.6 $\mu\text{m}$	4.5 $\mu\text{m}$	5.8 $\mu\text{m}$	8 $\mu\text{m}$	24 $\mu\text{m}$
J213531.36+573127.8	12.396 $\pm$ 0.003	12.151 $\pm$ 0.002	12.037 $\pm$ 0.008	11.731 $\pm$ 0.008	8.465 $\pm$ 0.037
J213532.89+573159.6	13.030 $\pm$ 0.002	12.653 $\pm$ 0.002	12.138 $\pm$ 0.006	11.119 $\pm$ 0.005	8.310 $\pm$ 0.032
J213625.07+572750.2	12.430 $\pm$ 0.002	11.862 $\pm$ 0.003	11.395 $\pm$ 0.004	10.346 $\pm$ 0.003	7.537 $\pm$ 0.016
J213659.09+573905.6	11.619 $\pm$ 0.002	11.151 $\pm$ 0.001	10.562 $\pm$ 0.002	9.482 $\pm$ 0.001	5.455 $\pm$ 0.003
J213659.47+573134.8	11.699 $\pm$ 0.002	11.269 $\pm$ 0.001	10.765 $\pm$ 0.003	9.598 $\pm$ 0.003	-
J213706.50+573231.8	10.803 $\pm$ 0.002	10.369 $\pm$ 0.001	9.984 $\pm$ 0.002	9.386 $\pm$ 0.001	6.901 $\pm$ 0.010
J213707.03+572700.8	10.684 $\pm$ 0.002	10.388 $\pm$ 0.001	10.165 $\pm$ 0.002	9.310 $\pm$ 0.001	5.833 $\pm$ 0.003
J213709.09+572548.6	13.057 $\pm$ 0.002	12.972 $\pm$ 0.002	12.829 $\pm$ 0.011	12.536 $\pm$ 0.015	9.395 $\pm$ 0.085
J213709.37+572948.4	11.061 $\pm$ 0.001	10.608 $\pm$ 0.001	10.142 $\pm$ 0.001	9.259 $\pm$ 0.001	6.381 $\pm$ 0.006
J213713.90+572727.1	11.415 $\pm$ 0.002	10.935 $\pm$ 0.001	10.489 $\pm$ 0.002	9.808 $\pm$ 0.002	6.310 $\pm$ 0.005
J213716.35+572640.4	11.259 $\pm$ 0.002	10.739 $\pm$ 0.002	10.288 $\pm$ 0.002	9.664 $\pm$ 0.002	5.552 $\pm$ 0.003
J213717.15+572847.3	13.297 $\pm$ 0.002	12.953 $\pm$ 0.002	12.410 $\pm$ 0.005	11.654 $\pm$ 0.007	8.854 $\pm$ 0.054
J213717.42+572927.3	11.686 $\pm$ 0.002	11.324 $\pm$ 0.002	11.054 $\pm$ 0.003	10.244 $\pm$ 0.003	7.088 $\pm$ 0.011
J213724.10+572411.5	11.761 $\pm$ 0.004	11.303 $\pm$ 0.003	10.856 $\pm$ 0.003	9.806 $\pm$ 0.002	6.357 $\pm$ 0.006
J213724.77+572909.0	12.418 $\pm$ 0.008	12.062 $\pm$ 0.002	11.656 $\pm$ 0.004	10.914 $\pm$ 0.003	8.604 $\pm$ 0.042
J213725.50+572745.3	12.757 $\pm$ 0.001	12.648 $\pm$ 0.002	12.559 $\pm$ 0.007	12.569 $\pm$ 0.016	-
J213734.65+571657.8	12.585 $\pm$ 0.002	12.214 $\pm$ 0.002	12.005 $\pm$ 0.005	11.354 $\pm$ 0.006	7.822 $\pm$ 0.020
J213735.72+573258.4	13.076 $\pm$ 0.001	12.723 $\pm$ 0.002	12.333 $\pm$ 0.004	11.588 $\pm$ 0.005	8.021 $\pm$ 0.025
J213738.51+573140.8	10.925 $\pm$ 0.002	10.647 $\pm$ 0.001	10.491 $\pm$ 0.002	10.049 $\pm$ 0.002	6.504 $\pm$ 0.006
J213742.77+573325.1	9.310 $\pm$ 0.002	9.023 $\pm$ 0.002	8.374 $\pm$ 0.002	7.496 $\pm$ 0.002	4.418 $\pm$ 0.001
J213742.90+573446.3	13.115 $\pm$ 0.002	12.773 $\pm$ 0.002	12.334 $\pm$ 0.007	11.597 $\pm$ 0.007	7.092 $\pm$ 0.011
J213745.15+571942.3	11.403 $\pm$ 0.002	10.926 $\pm$ 0.001	10.485 $\pm$ 0.002	9.582 $\pm$ 0.002	6.688 $\pm$ 0.007
J213746.13+573428.1	12.930 $\pm$ 0.002	12.582 $\pm$ 0.002	12.116 $\pm$ 0.004	11.532 $\pm$ 0.006	-
J213747.97+573242.3	12.011 $\pm$ 0.002	11.676 $\pm$ 0.002	11.349 $\pm$ 0.003	10.465 $\pm$ 0.003	6.508 $\pm$ 0.006
J213748.94+572321.0	12.336 $\pm$ 0.001	11.985 $\pm$ 0.001	11.599 $\pm$ 0.003	10.908 $\pm$ 0.004	7.886 $\pm$ 0.021
J213757.62+572247.7	10.852 $\pm$ 0.002	10.535 $\pm$ 0.002	10.205 $\pm$ 0.002	9.411 $\pm$ 0.002	6.393 $\pm$ 0.006
J213759.26+573616.1	11.133 $\pm$ 0.001	10.743 $\pm$ 0.001	10.370 $\pm$ 0.002	9.566 $\pm$ 0.002	6.540 $\pm$ 0.006
J213808.56+573707.5	11.462 $\pm$ 0.002	11.029 $\pm$ 0.002	10.774 $\pm$ 0.003	10.215 $\pm$ 0.003	7.536 $\pm$ 0.016
J213809.24+572020.0	11.909 $\pm$ 0.007	11.616 $\pm$ 0.005	11.246 $\pm$ 0.004	10.608 $\pm$ 0.003	7.857 $\pm$ 0.021
J213810.00+572352.8	12.729 $\pm$ 0.009	12.286 $\pm$ 0.004	11.765 $\pm$ 0.005	10.890 $\pm$ 0.004	8.609 $\pm$ 0.041
J213812.64+572033.7	12.435 $\pm$ 0.002	12.099 $\pm$ 0.002	11.744 $\pm$ 0.005	10.831 $\pm$ 0.004	7.153 $\pm$ 0.011
J213822.81+574017.3	12.359 $\pm$ 0.002	12.014 $\pm$ 0.001	11.746 $\pm$ 0.003	11.231 $\pm$ 0.003	8.412 $\pm$ 0.034
J213823.94+572736.3	13.006 $\pm$ 0.002	12.739 $\pm$ 0.002	12.448 $\pm$ 0.008	11.876 $\pm$ 0.008	9.260 $\pm$ 0.074
J213825.90+574204.8	10.712 $\pm$ 0.002	10.644 $\pm$ 0.002	10.468 $\pm$ 0.002	9.809 $\pm$ 0.002	6.682 $\pm$ 0.007
J213827.41+572720.7	10.406 $\pm$ 0.002	9.914 $\pm$ 0.002	9.416 $\pm$ 0.002	8.539 $\pm$ 0.002	5.821 $\pm$ 0.003
J213827.42+573108.2	10.299 $\pm$ 0.002	9.645 $\pm$ 0.001	9.307 $\pm$ 0.002	8.788 $\pm$ 0.001	6.944 $\pm$ 0.009
J213830.30+573255.2	13.162 $\pm$ 0.002	12.961 $\pm$ 0.004	12.929 $\pm$ 0.008	12.751 $\pm$ 0.017	8.284 $\pm$ 0.031
J213832.15+572636.0	12.303 $\pm$ 0.002	11.869 $\pm$ 0.002	11.610 $\pm$ 0.004	11.163 $\pm$ 0.005	8.217 $\pm$ 0.029
J213837.43+573020.7	12.936 $\pm$ 0.003	12.601 $\pm$ 0.001	12.122 $\pm$ 0.005	11.297 $\pm$ 0.004	8.491 $\pm$ 0.037
J213840.02+573930.4	12.379 $\pm$ 0.001	12.073 $\pm$ 0.002	11.719 $\pm$ 0.003	11.295 $\pm$ 0.005	-
J213844.45+571809.0	11.369 $\pm$ 0.002	11.086 $\pm$ 0.002	10.799 $\pm$ 0.003	10.077 $\pm$ 0.003	6.626 $\pm$ 0.007
J213855.21+572237.5	14.282 $\pm$ 0.003	14.112 $\pm$ 0.004	14.053 $\pm$ 0.028	14.204 $\pm$ 0.066	-
J213858.08+574334.4	11.366 $\pm$ 0.002	11.070 $\pm$ 0.002	10.800 $\pm$ 0.003	10.200 $\pm$ 0.003	6.829 $\pm$ 0.008
J213903.21+573042.1	12.488 $\pm$ 0.002	12.118 $\pm$ 0.002	11.678 $\pm$ 0.003	11.031 $\pm$ 0.004	8.717 $\pm$ 0.045
J213930.87+572227.3	11.916 $\pm$ 0.002	11.567 $\pm$ 0.002	11.408 $\pm$ 0.004	11.107 $\pm$ 0.005	8.387 $\pm$ 0.034
J213935.61+571821.9	12.628 $\pm$ 0.002	12.273 $\pm$ 0.002	11.913 $\pm$ 0.005	11.208 $\pm$ 0.004	8.703 $\pm$ 0.044
J213945.87+573051.8	12.129 $\pm$ 0.018	11.526 $\pm$ 0.015	11.337 $\pm$ 0.010	10.825 $\pm$ 0.004	7.318 $\pm$ 0.013
J213951.88+572658.3	12.757 $\pm$ 0.002	12.453 $\pm$ 0.002	12.155 $\pm$ 0.006	11.533 $\pm$ 0.006	8.643 $\pm$ 0.043
J213954.07+572933.5	12.366 $\pm$ 0.001	12.067 $\pm$ 0.002	11.870 $\pm$ 0.004	11.451 $\pm$ 0.006	7.077 $\pm$ 0.010
J213955.69+571638.2	13.344 $\pm$ 0.002	13.285 $\pm$ 0.003	13.254 $\pm$ 0.013	12.518 $\pm$ 0.014	8.180 $\pm$ 0.028
J213956.13+572707.9	13.477 $\pm$ 0.002	13.070 $\pm$ 0.002	12.829 $\pm$ 0.010	12.241 $\pm$ 0.010	-
J213957.70+572343.3	13.883 $\pm$ 0.002	13.591 $\pm$ 0.003	13.208 $\pm$ 0.011	12.498 $\pm$ 0.014	9.470 $\pm$ 0.088
J214004.53+572836.3	10.895 $\pm$ 0.002	10.518 $\pm$ 0.002	10.032 $\pm$ 0.002	9.120 $\pm$ 0.002	6.520 $\pm$ 0.006
J214011.35+573951.8	12.044 $\pm$ 0.002	11.706 $\pm$ 0.001	11.482 $\pm$ 0.004	10.976 $\pm$ 0.003	-
J214011.85+574012.0	12.117 $\pm$ 0.002	11.971 $\pm$ 0.002	11.825 $\pm$ 0.005	11.430 $\pm$ 0.005	-
J214021.30+572657.9	11.465 $\pm$ 0.001	11.004 $\pm$ 0.001	10.810 $\pm$ 0.002	10.407 $\pm$ 0.002	7.721 $\pm$ 0.018
J214025.37+573416.2	12.827 $\pm$ 0.001	12.507 $\pm$ 0.002	12.116 $\pm$ 0.060	11.878 $\pm$ 0.025	8.458 $\pm$ 0.052
J214035.33+572309.6	10.233 $\pm$ 0.001	10.153 $\pm$ 0.001	9.993 $\pm$ 0.002	9.811 $\pm$ 0.002	6.087 $\pm$ 0.004
J214035.75+573455.1	11.404 $\pm$ 0.002	11.033 $\pm$ 0.002	10.656 $\pm$ 0.002	10.051 $\pm$ 0.002	-
J214040.73+573214.8	12.987 $\pm$ 0.002	12.637 $\pm$ 0.002	12.335 $\pm$ 0.007	11.781 $\pm$ 0.008	-
J214100.84+572303.0	12.884 $\pm$ 0.002	12.690 $\pm$ 0.002	12.478 $\pm$ 0.007	11.981 $\pm$ 0.009	-



**APPENDIX E: REJECTED CANDIDATES**

Table E1: IPHAS and 2MASS photometry of objects towards IC 1396 which, based on their outlier position in the ( $r - H\alpha$ ,  $r' - i'$ ) diagram, are likely to be  $H\alpha$  emission-line objects. However, their position in the ( $r'$ ,  $r' - i'$ ) suggests that they are not related to IC 1396 itself, so they have not been considered as T Tauri candidates in our work. See §2.4 for details.

	IPHAS ID J[RA(2000)+Dec(2000)]	IPHAS photometry			2MASS magnitudes		
		$r'$	$r' - i'$	$r' - H\alpha$	$J$	$H$	$K$
1	J213412.45+562519.6	17.83±0.01	0.05±0.02	0.44±0.01	-	-	-
2	J213415.13+564734.3	18.91±0.02	1.14±0.03	1.65±0.02	15.61±0.07	14.66±0.08	13.79±0.06
3	J213454.21+562406.8	15.16±0.00	0.33±0.00	0.41±0.00	13.23±0.02	12.95±0.04	12.86±0.03
4	J213459.66+580325.9	18.02±0.01	0.86±0.01	0.69±0.02	15.62±0.09	15.43±0.16	15.05±0.14
5	J213502.78+572747.1	16.79±0.00	0.68±0.01	0.57±0.01	14.89±0.03	14.67±0.07	14.36±0.09
6	J213520.74+561049.3	16.60±0.00	0.02±0.01	0.20±0.01	16.41±0.14	> 16.49	> 16.19
7	J213557.95+572532.2	18.50±0.01	0.88±0.02	0.77±0.02	16.11±0.11	15.49±0.14	15.42±0.19
8	J213626.91+570213.7	15.86±0.00	0.65±0.00	0.61±0.00	14.15±0.03	13.78±0.04	13.56±0.04
9	J213636.44+572812.2	19.59±0.03	0.97±0.05	1.02±0.05	-	-	-
10	J213646.48+562716.6	16.59±0.00	0.63±0.01	0.96±0.01	14.88±0.03	14.56±0.06	14.41±0.08
11	J213710.08+564137.8	16.19±0.00	0.78±0.01	0.63±0.01	14.23±0.03	13.89±0.04	13.53±0.05
12	J213733.15+571812.9	15.72±0.00	0.61±0.00	0.61±0.00	14.48±0.04	14.26±0.06	14.16±0.08
13	J213738.67+563731.5	16.80±0.00	0.71±0.01	0.58±0.01	15.01±0.04	14.53±0.06	14.38±0.10
14	J213745.43+564657.6	18.21±0.01	0.86±0.02	0.72±0.02	> 15.80	15.19±0.10	> 14.77
15	J213745.74+565322.8	15.85±0.00	0.55±0.01	0.49±0.01	14.15±0.04	13.88±0.05	13.65±0.05
16	J213748.05+563637.7	17.53±0.01	0.76±0.01	0.76±0.01	15.66±0.09	15.31±0.13	14.94±0.14
17	J213800.10+573855.8	16.05±0.00	0.62±0.00	0.73±0.00	14.43±0.04	13.92±0.05	13.57±0.04
18	J213808.15+564425.2	18.58±0.01	0.10±0.04	0.36±0.03	-	-	-
19	J213910.98+565115.1	15.08±0.00	0.67±0.00	0.57±0.00	13.33±0.02	13.00±0.03	12.67±0.02
20	J213944.06+574138.0	17.28±0.01	0.17±0.02	0.35±0.01	14.95±0.05	14.41±0.06	14.12±0.07
21	J214049.80+583915.1	15.65±0.00	0.72±0.00	0.72±0.00	13.53±0.03	13.02±0.03	12.68±0.03
22	J214102.28+585229.3	19.49±0.03	0.85±0.06	0.80±0.05	-	-	-
23	J214202.26+561403.1	19.16±0.03	1.20±0.04	1.02±0.04	-	-	-
24	J214310.68+563644.8	19.81±0.04	1.17±0.06	1.73±0.05	15.87±0.07	14.35±0.06	13.26±0.03
25	J214330.29+575433.0	17.40±0.01	0.76±0.01	0.63±0.01	15.05±0.04	14.60±0.05	14.19±0.05
26	J214443.70+571419.8	18.41±0.02	0.93±0.03	1.13±0.03	16.35±0.12	15.05±0.08	13.81±0.05
27	J214513.93+572241.7	14.73±0.00	0.13±0.00	0.21±0.00	14.60±0.08	14.62±0.09	14.70±0.11
28	J214524.56+575548.8	17.25±0.01	0.76±0.01	0.62±0.01	14.39±0.06	13.01±0.04	11.67±0.02
29	J214624.10+583327.5	14.82±0.00	0.58±0.01	0.51±0.01	13.13±0.03	12.75±0.04	12.60±0.03
30	J214930.22+570323.1	18.68±0.03	1.27±0.04	0.98±0.04	15.31±0.05	14.68±0.06	14.20±0.05